Last lecture (7)

- Particle motion in magnetosphere
- Aurora

Today’s lecture (8)

- Aurora on other planets
- How to measure currents in space
- Magnetospheric dynamics

Today

Activity	Date	Time	Room	Subject	Litterature
L1	31/8	13-15	V22	Course description, Introduction, The Sun 1, Plasma physics 1	CGF Ch 1, 5, (p 110-113)
L2	3/9	15-17	Q36	The Sun 2, Plasma physics 2	CGF Ch 5 (p 114-121), 6.3
L3	7/9	13-15	Q36	Solar wind, The ionosphere and atmosphere 1, Plasma physics 3	CGF Ch 6.1, 2.1-2.6, 3.1-3.2, 3.5, LL Ch III, Extra material
T1	10/9	15-17	Q36	Mini-group work 1	
L4	14/9	13-15	E2	The ionosphere 2, Plasma physics 4	CGF Ch 3.4, 3.7, 3.8
T2	17/9	8-10	Q31	Mini-group work 2	
L5	17/9	15-17	L52	The Earth's magnetosphere 1, Plasma physics 5	CGF 4.1-4.3, LL Ch I, II, IV.A
L6	21/9	13-15	L52	The Earth's magnetosphere 2, Other magnetospheres	CGF Ch 4.6-4.9, LL Ch V.
T3	24/9	16-18	Q36	Mini-group work 3	
L7	28/9	13-15	Q36	Aurora, Measurement methods in space plasmas and data analysis 1	CGF Ch 4.5, 10, LL Ch VI, Extra material
T4	1/10	15-17	V22	Mini-group work 4	
L8	5/10	13-15	M33	Space weather and geomagnetic storms	CGF Ch 4.4, LL Ch IV.B-C, VII.A-C
L9	6/10	8-10	Q36	Interstellar and intergalactic plasma, Cosmic radiation,	CGF Ch 7-9
T5	8/10	15-17	Q34	Mini-group work 5	
L10	12/10	13-15	Q36	Swedish and international space physics research.	
T6	15/10	15-17	Q33	Round-up.	
Written examination	28/10	8-13	Q21, Q26		

Mini-groupwork 4

a)

$$\rho_{SW} v_{SW}^2 = \left[\frac{\mu_0 a}{4\pi} \frac{1}{r^3} \right]^2 / 2\mu_0 \quad \Rightarrow$$

$$r = \left(\frac{\mu_0 a}{4\pi} \right)^{1/3} \left(2\mu_0 \rho_{SW} v_{SW}^2 \right)^{-1/6}$$

Assuming the solar wind consists of protons

$$\rho_{SW} = n_{e,SW} m_p = 1.7 \cdot 10^{-22}\ kg\, m^{-3}$$

Thus

$r = 2.7 \cdot 10^9$ m ≈ 38 R_J

Mini-groupwork 4

b)

$$\rho_{SW} v_{SW}^2 = \left[\frac{\mu_0 a}{4\pi} \frac{1}{r^3} \right]^2 / 2\mu_0 + 2 n_e k_B T \quad \Rightarrow$$

$$\rho_{SW} v_{SW}^2 = \left[\frac{\mu_0 a}{4\pi} \frac{1}{r^3} \right]^2 / 2\mu_0 + 2 n_{e0} \left(\frac{R_J}{r} \right)^3 k_B T$$

Substitute $x = 1/r^3$. This gives you an equation on the form

$ax^2 + bx + c = 0$

with

$$a = \left[\frac{\mu_0 a}{4\pi} \right]^2 / 2\mu_0 = 1.02 \cdot 10^{46}$$

$$b = 2 n_{e0} R_J^3 k_B T = 3.6 \times 10^{18}$$

$$c = -\rho_{SW} v_{SW}^2 = -2.7 \cdot 10^{-11}$$

$$x = \frac{-b}{2a} \pm \sqrt{\frac{b^2}{4a^2} - \frac{c}{a}} = -1.8 \cdot 10^{-28} + \sqrt{3.24 \cdot 10^{-56} + 2.635 \cdot 10^{-57}} =$$

$$= -1.8 \cdot 10^{-28} + 1.87 \cdot 10^{-28} = 7.18 \cdot 10^{-30}$$

From this you get $r \approx 73\ \mathrm{R_J}$

Planetary magnetospheres

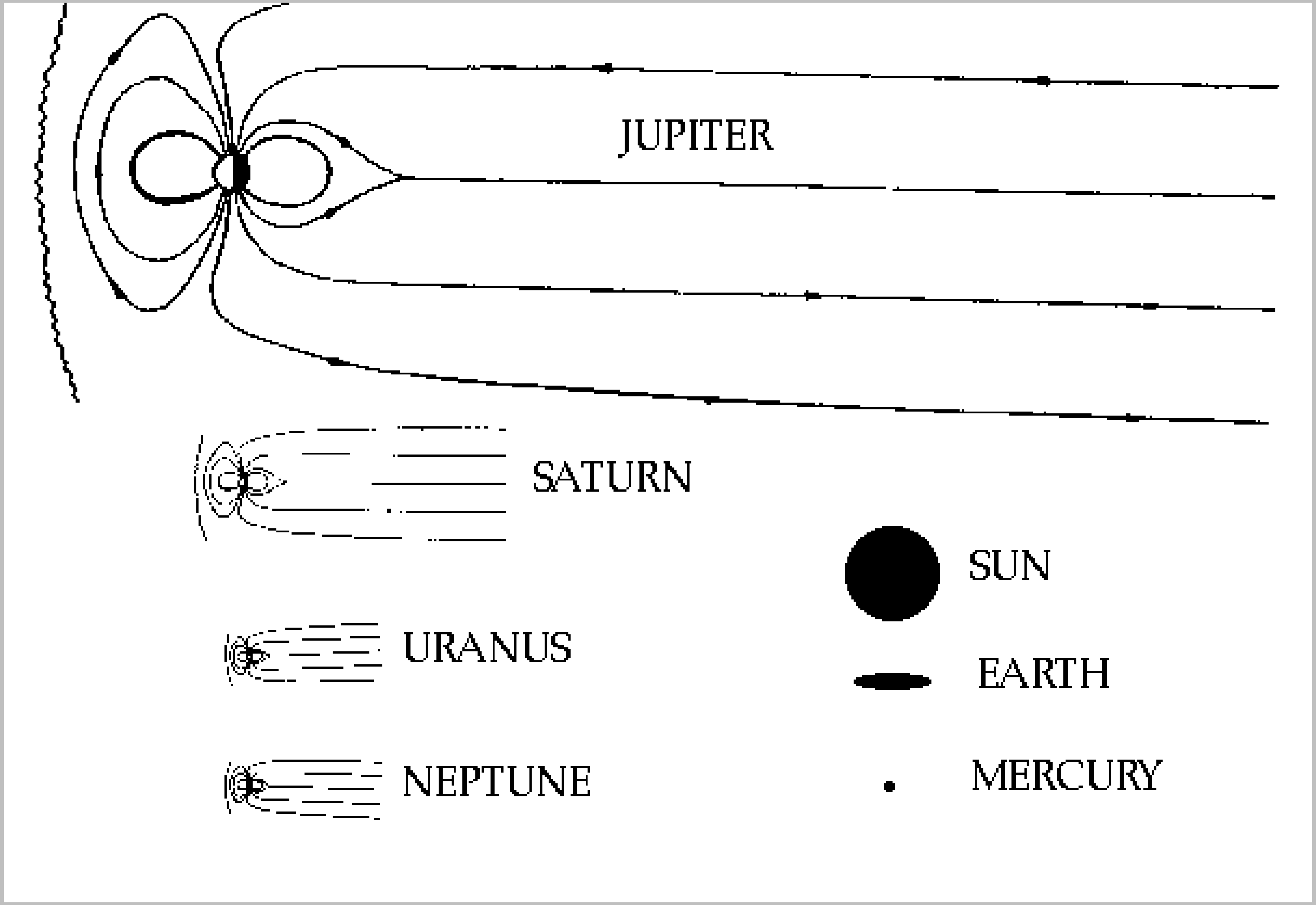
	Radius Earth radii	Spin period (days)	Equatorial field strength (μT)	Magnetic axis direction relative to spin axis	Polarity relative to Earth´s	Typical magneto-pause distance (planetary radii)
Mercury	0.38	58.6	0.35	10^o	Same	1.1
Venus	0.95	243	**< 0.03**	-	-	1.1
Earth	1.0	1	31	11.5^o	Same	10
Mars	0.53	1.02	**0.065**	-	Opposite	?
Jupiter	11.18	0.41	410	10^o	Opposite	60-100
Saturn	9.42	0.44	40	$<1^o$	Opposite	20-25
Uranus	3.84	0.72	23	60^o	Opposite	18-25
Neptune	3.93	0.74	20-150*)	47^o	Opposite	26**)

*) The magnetic field differs greatly from a dipole field. The numbers represent maximum and minimum strength at the planetary surface

**) Based on single passage

Very weak magnetic fields

Relative size of the magnetospheres

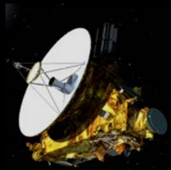


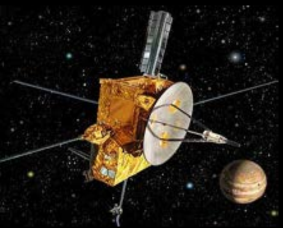
Comparative magnetospheres

In situ observations

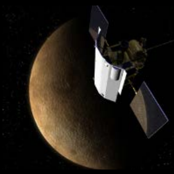

Galileo


Mariner 10


New Horizons


Ulysses

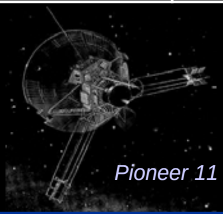

Rosetta

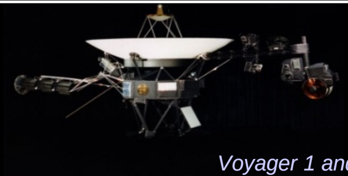

Messenger

Space probe	Celestial body	Observations
Mariner 10	Mercury	*1974 – 1975*
*Messenger **	Mercury	*2008 – 2015*
Pioneer 10,11	Jupiter, Saturn	*1973 – 1979*
Voyager 1,2	Jupiter, Saturn, Uranus, Neptune	*1977 – 1989*
Ulysses	Jupiter	*1992*
*Galileo**	Jupiter	*1995 – 2003*
*Cassini**	Jupiter, Saturn	*2004 –*
New Horizons	Jupiter	*2007*
Rosetta	Churymov-Gerasimenko	*2014 - 2016*

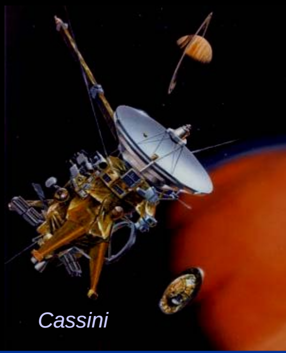
* Orbiters


Pioneer 10


Pioneer 11



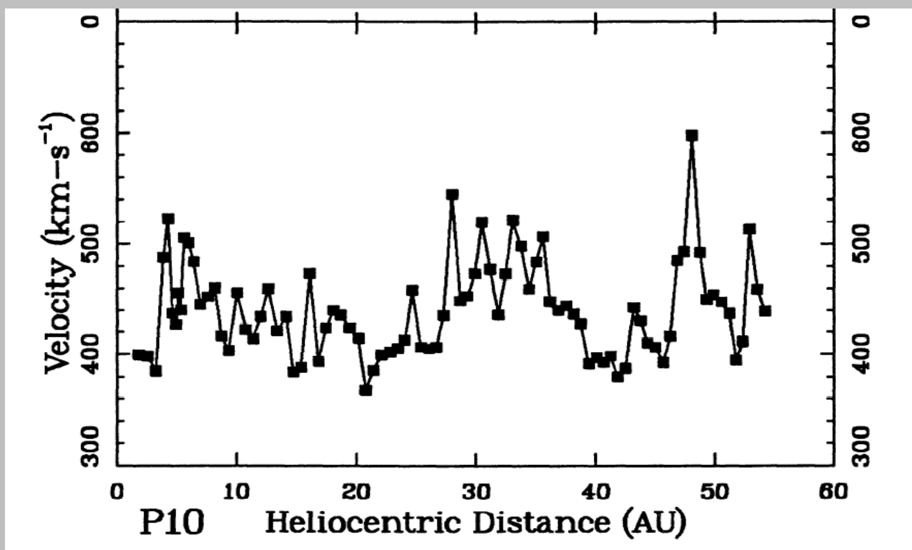

Voyager 1 and 2


Cassini

Comparative magnetospheres

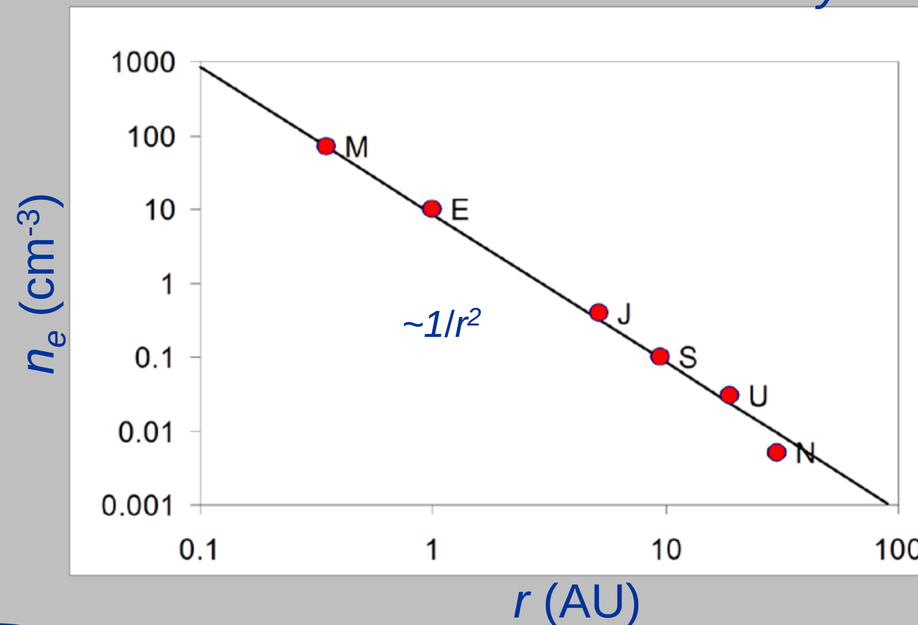
Solar wind properties

Solar wind velocity

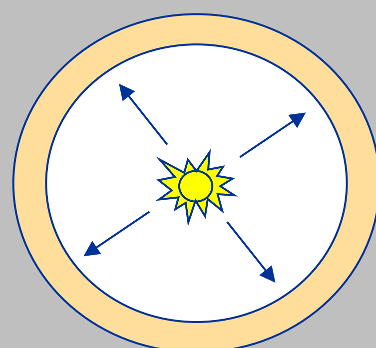


Pioneer 10, measurements
[Grazin et al., 1994]

Solar wind electron density



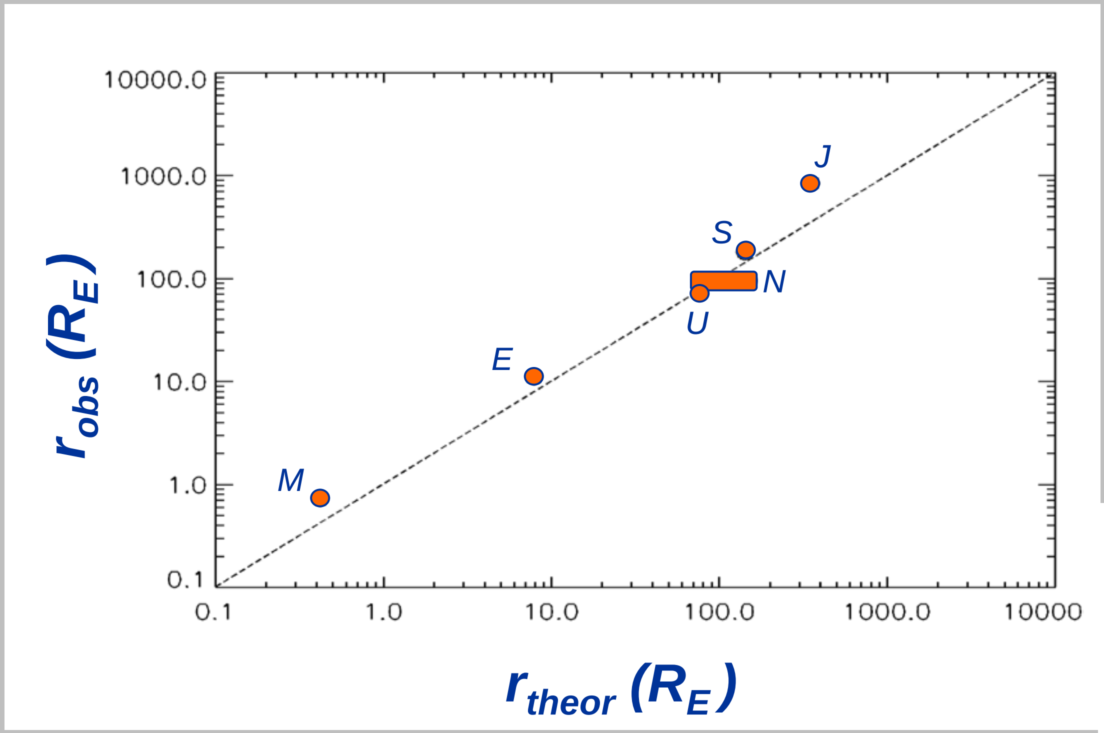
[Blanc et al., 2005]



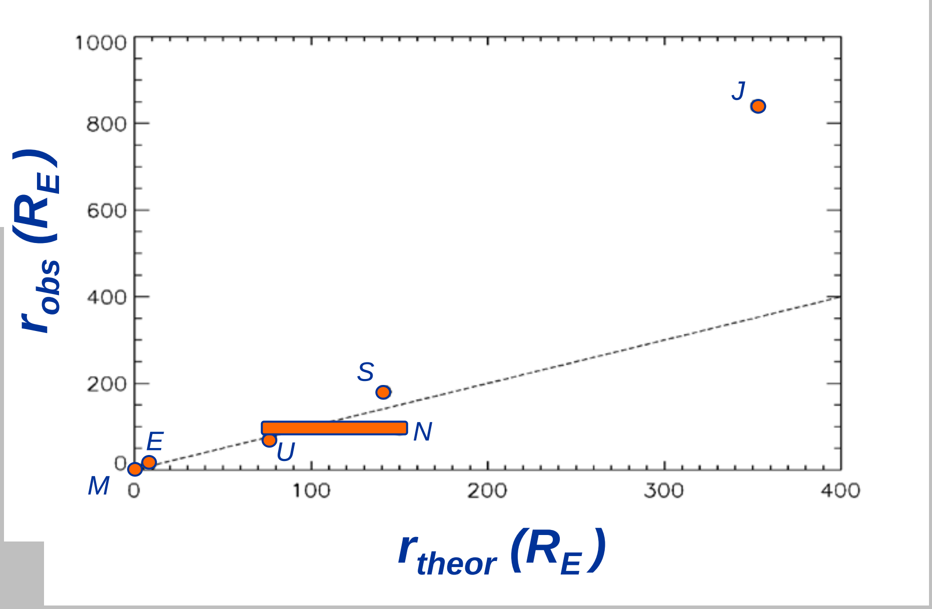
$$dV = 4\pi r^2 dr$$

Comparative magnetospheres

Observed vs. theoretical standoff-distance



$$r_{theor} = \left(\frac{\mu_0 a}{4\pi}\right)^{1/3} \left(2\mu_0 \rho_{SW} v_{SW}^2\right)^{-1/6}$$

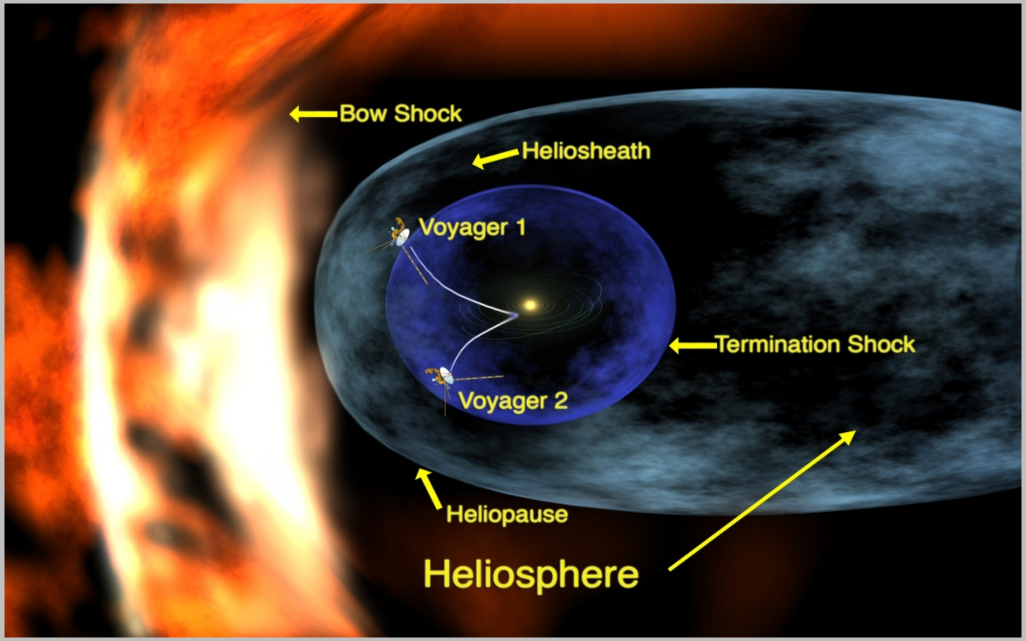


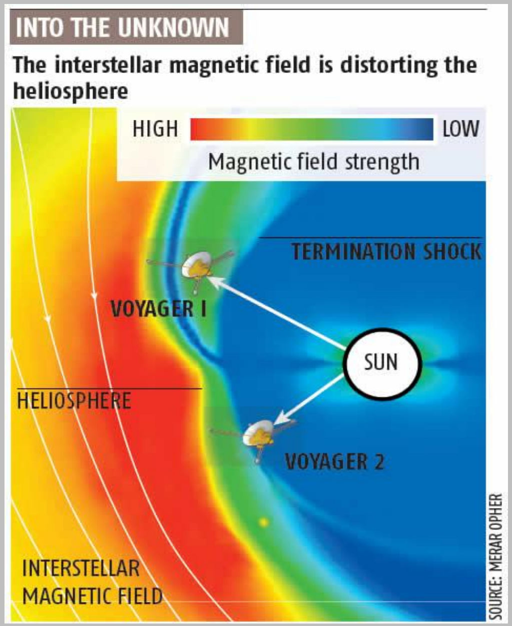
- *Model reasonably valid over three orders of magnitude*
- *Size of Jupiter's (and maybe Saturn's) magnetosphere underestimated*

Other other magnetospheres

Heliosphere

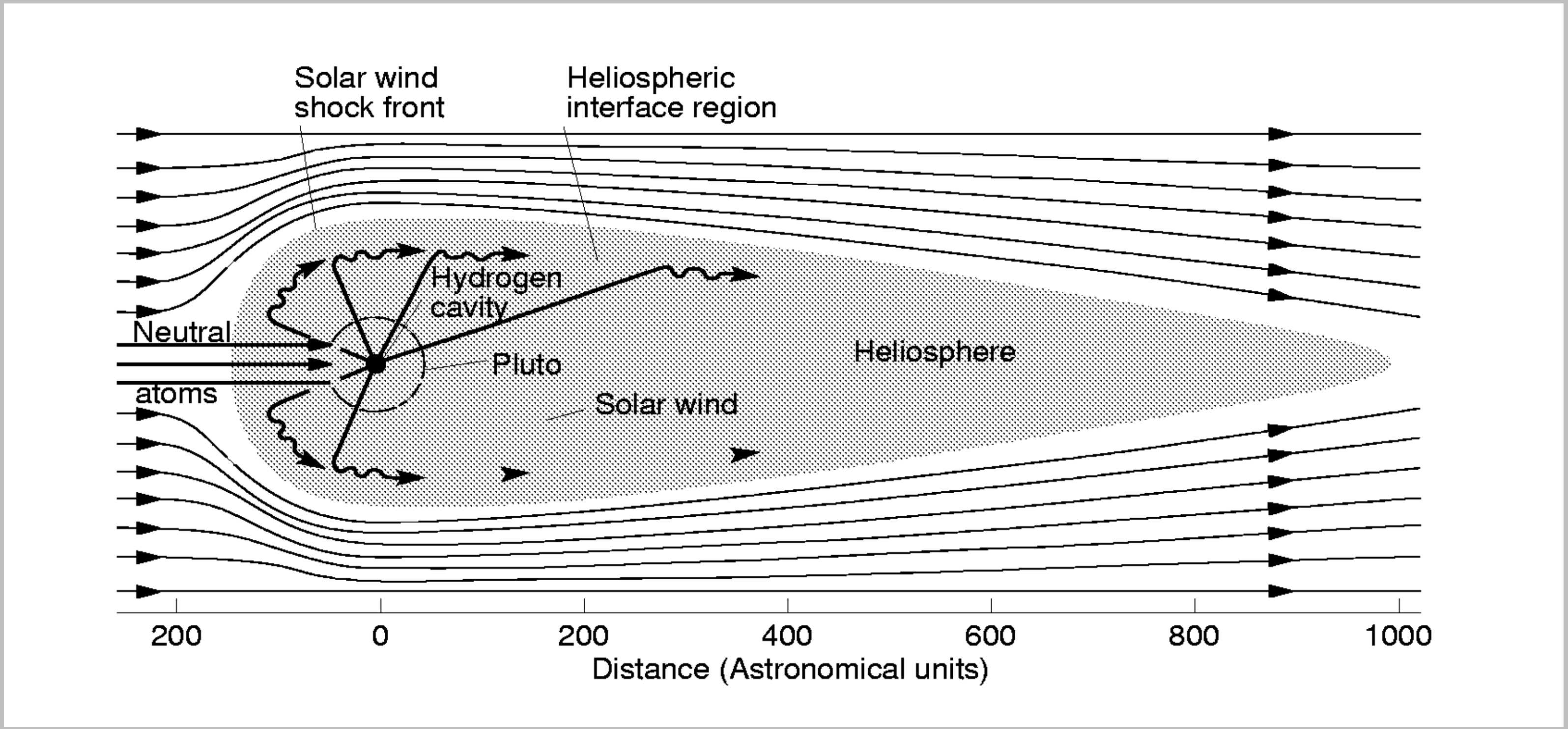
Interstellar wind





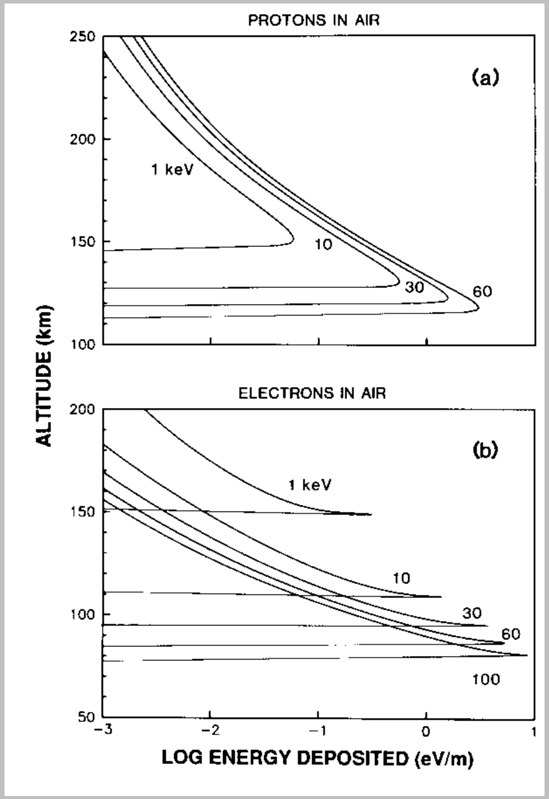
[Opher, 2007]

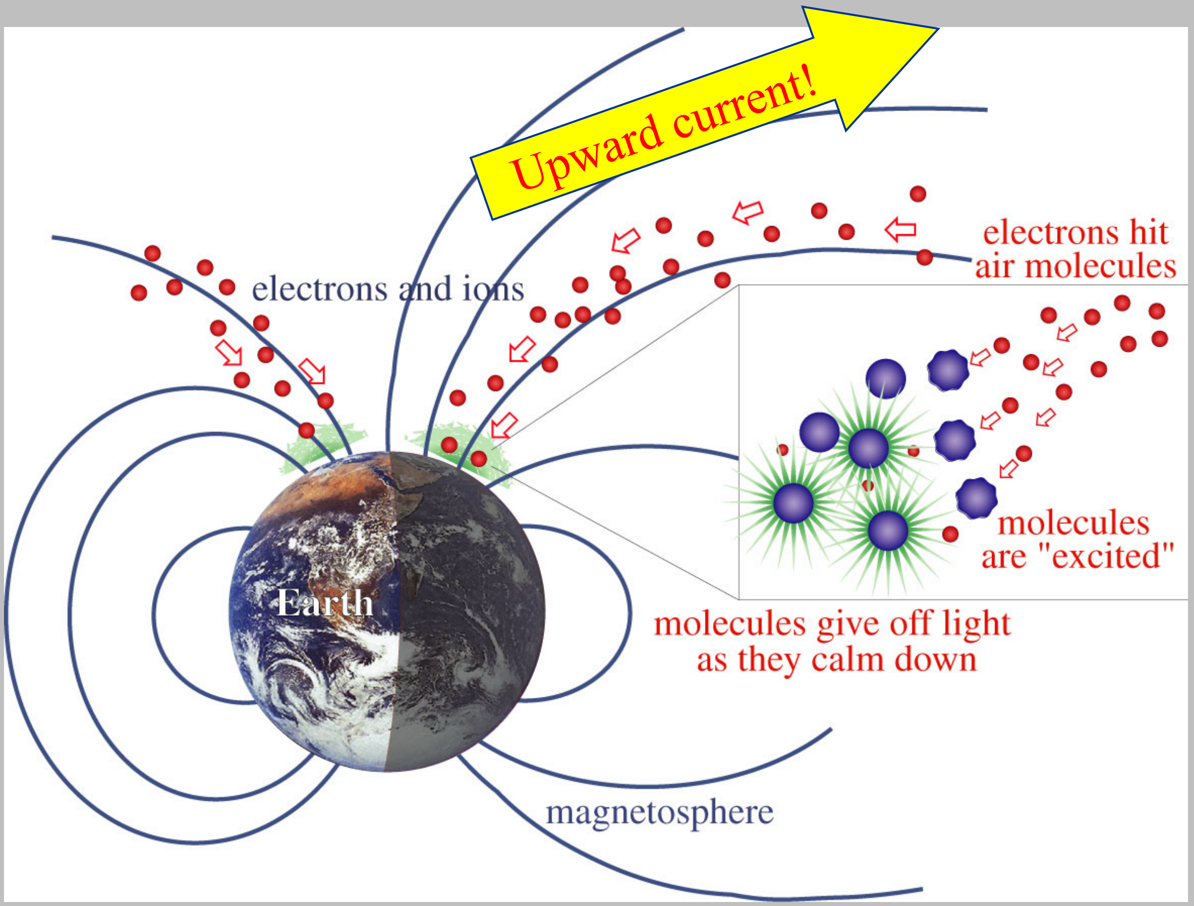
Heliosphere



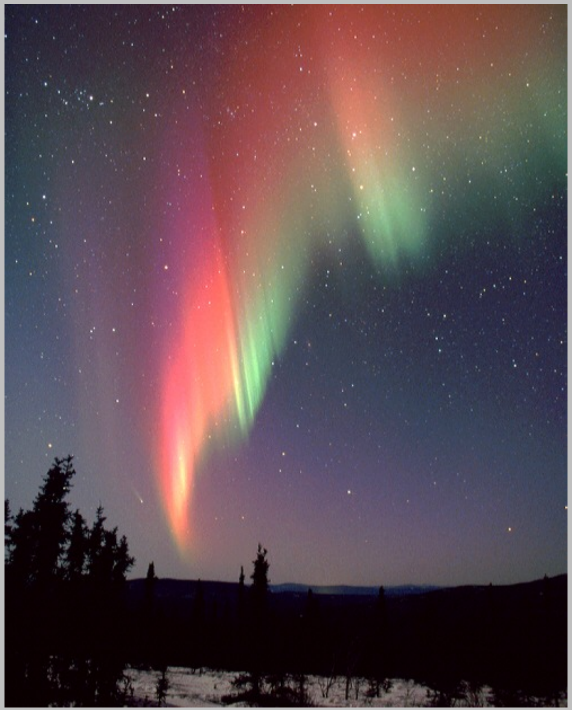
- Reaches approximately 100 AU into space (=1.5x10^{13} m)
- Voyager sonds are approaching/encountering the heliopause right now

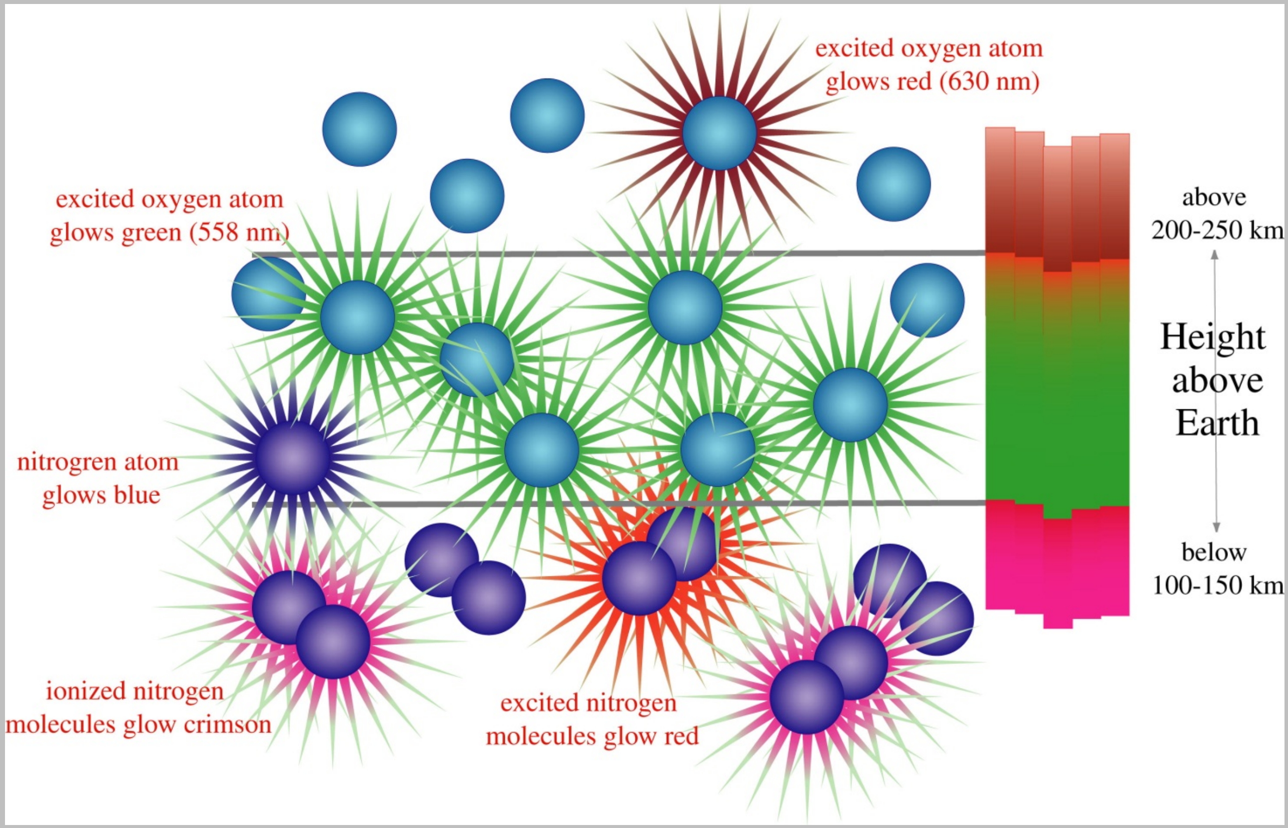
Collisions - emissions





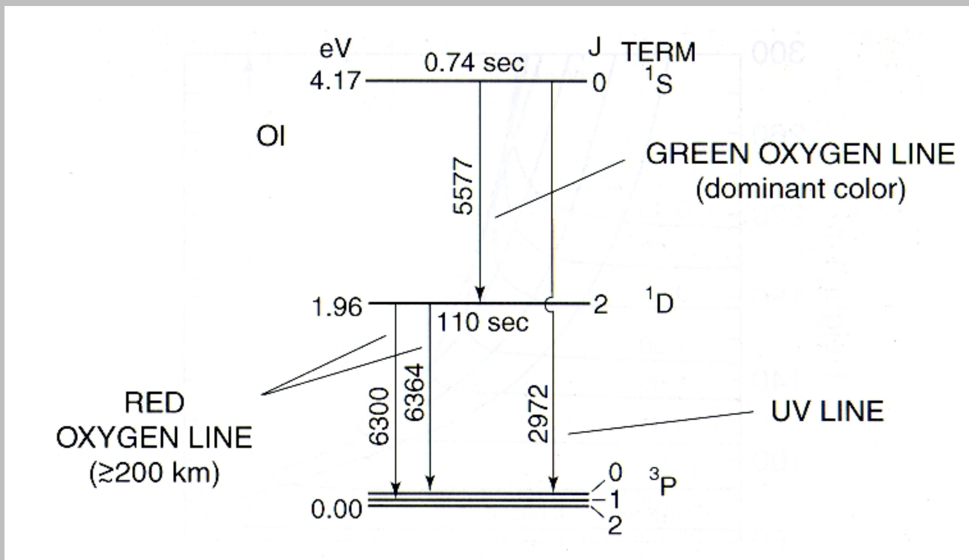
Emissions



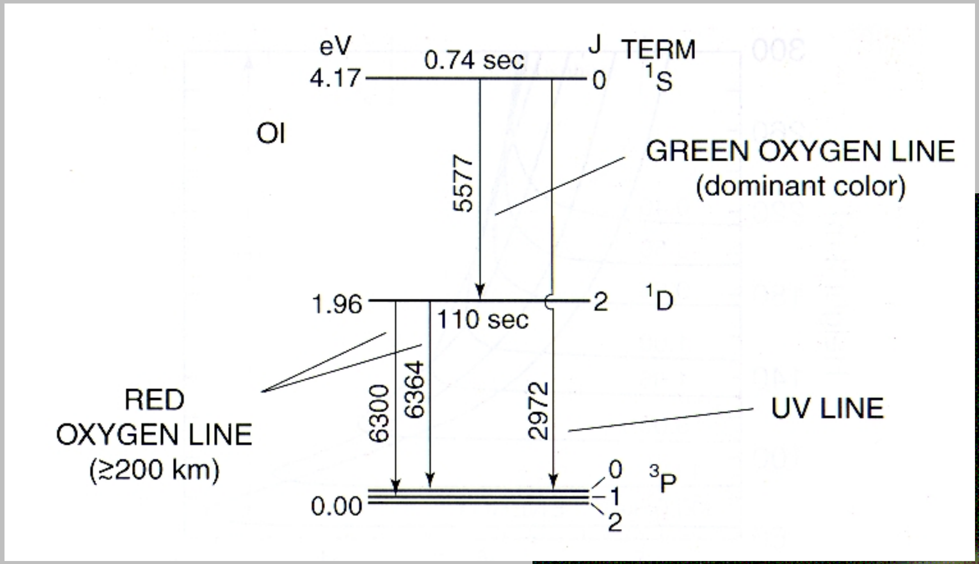


excited oxygen atom
glows red (630 nm)
excited oxygen atom
glows green (558 nm)
nitrogren atom
glows blue
ionized nitrogen
molecules glow crimson
excited nitrogen
molecules glow red
above
200-250 km
Height
above
Earth
below
100-150 km

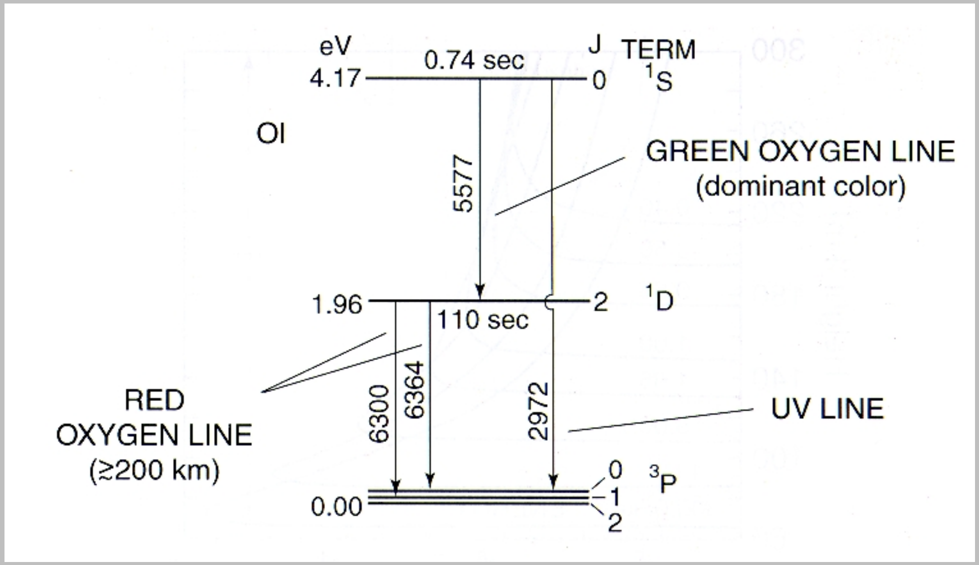
Oxygen emissions

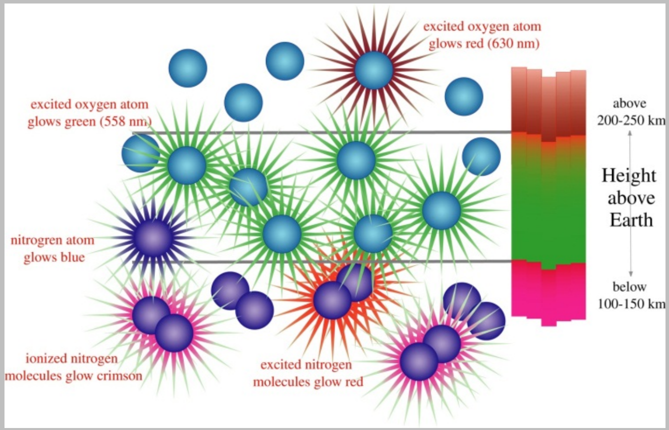


Why is there no red emissions at lower altitude?

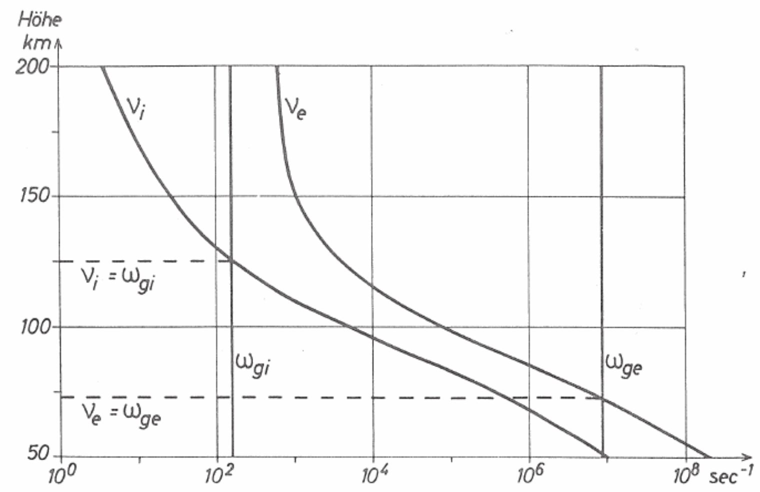


Oxygen emissions





The red emission line is suppressed by collisions at lower altitudes due the its long transition time. (When an excited atom collides with another atom, is is de-excited without any emission.)



Larger scales

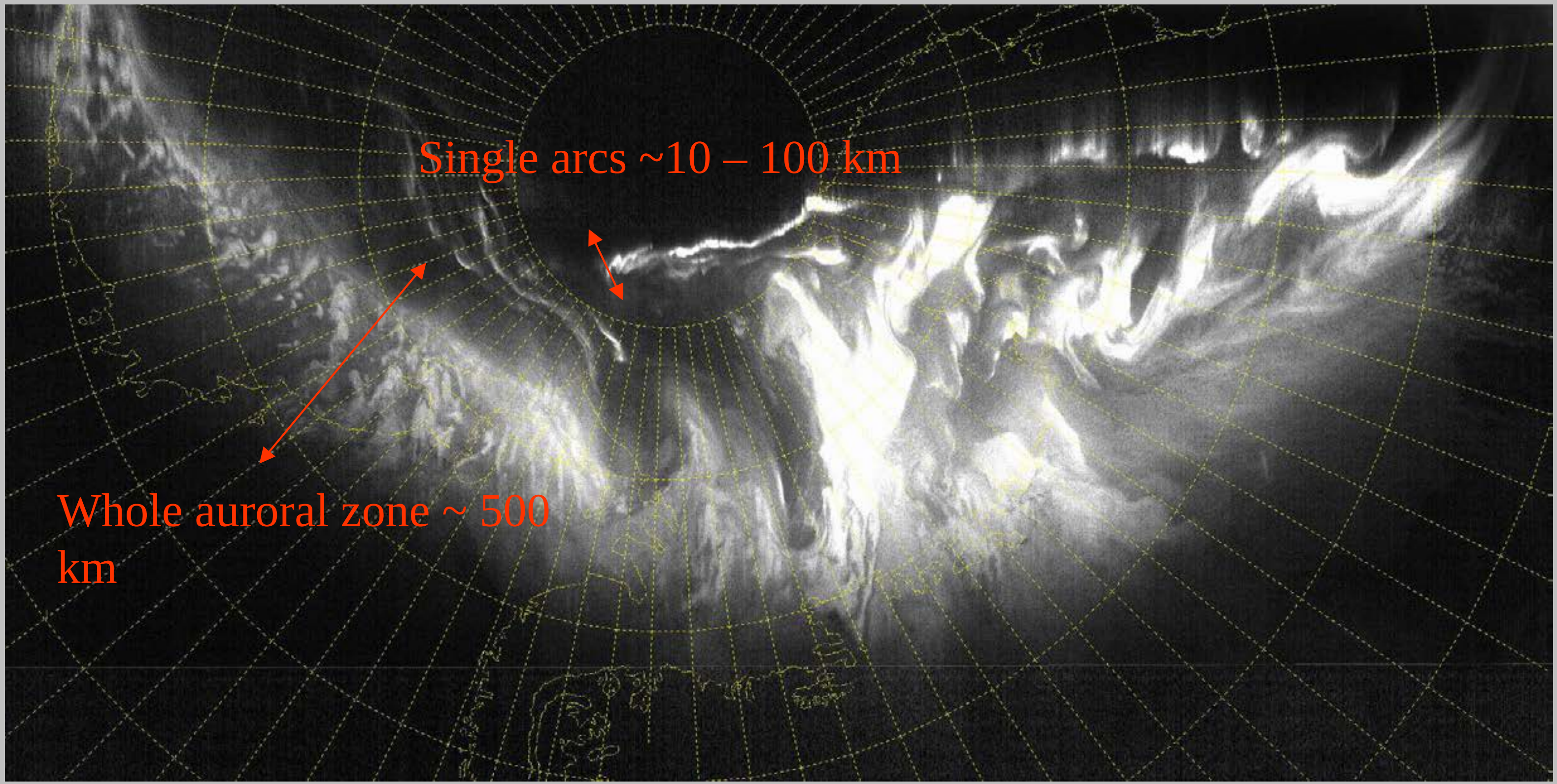
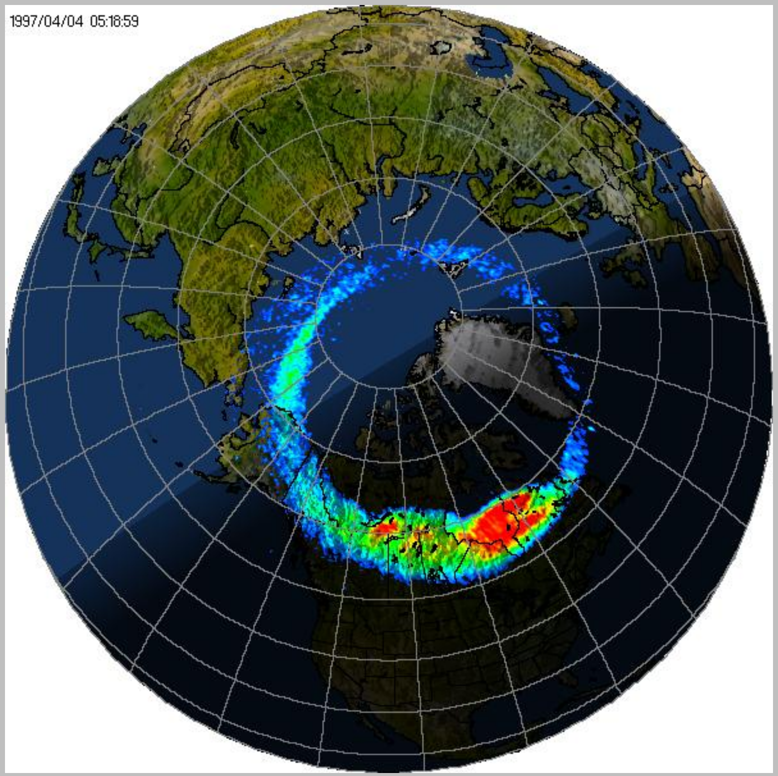
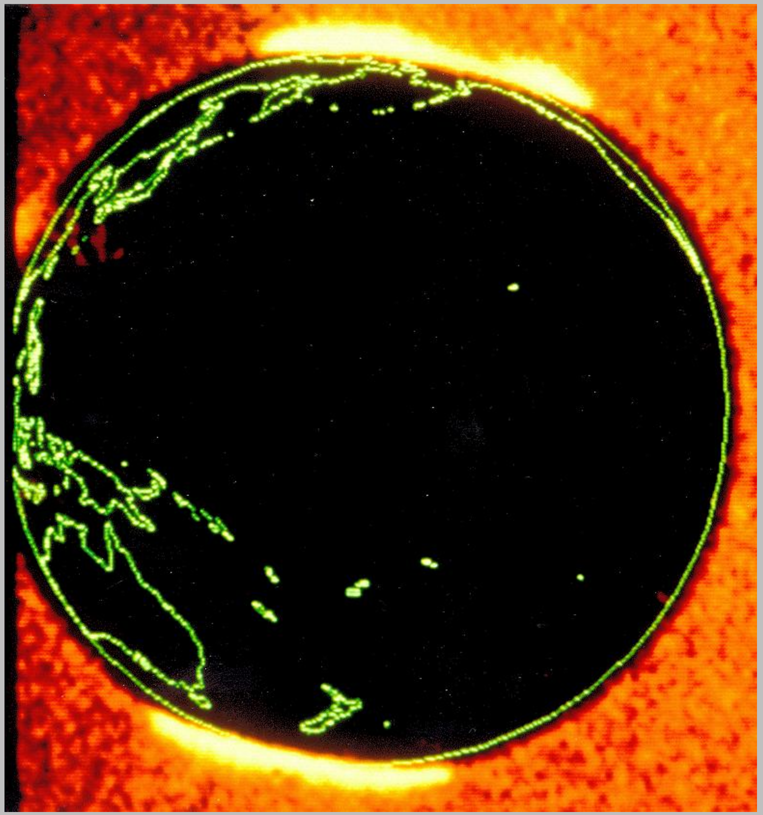


Foto från DMSP-satelliten

Auroral ovals



Polar

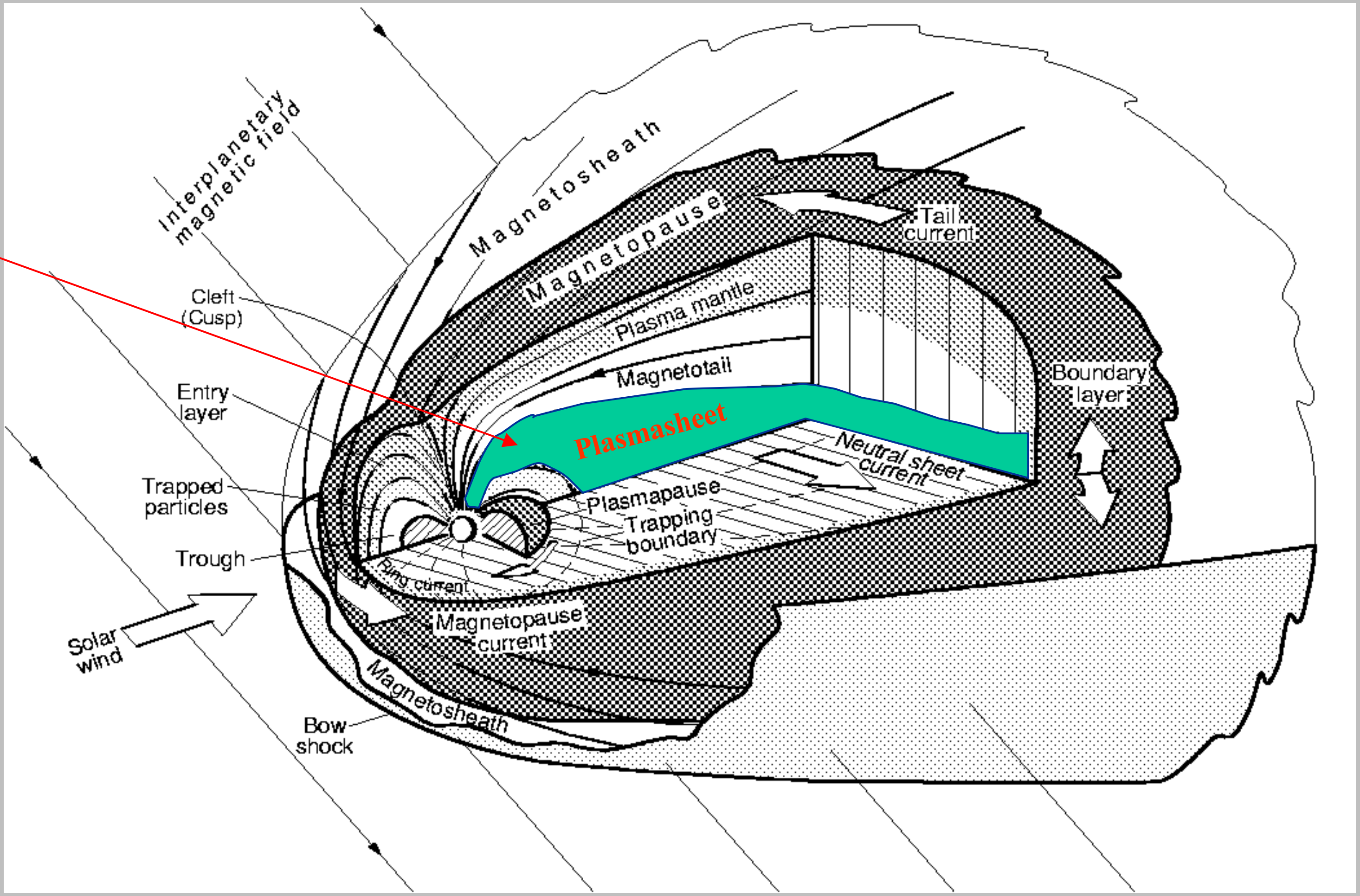


Dynamics Explorer

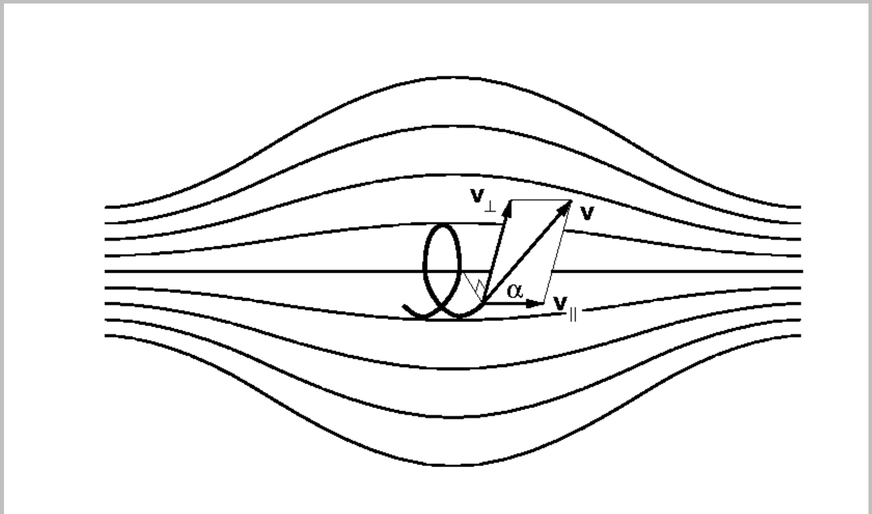
The auroral oval is the projection of the plasmasheet onto the atmosphere

Mystery!

The particles in the plasmasheet do not have high enough energy to create aurora visible to the eye.



Magnetic mirror



$mv^2/2$ constant (energy conservation) ➡

$$\frac{\sin^2\alpha}{B} = konst$$

particle turns when α = 90° ➡

$$B_{turn} = B/\sin^2\alpha$$

If maximal B-field is B_{max} a particle with pitch angle α can only be turned around if

$$B_{turn} = B/\sin^2\alpha \le B_{max}$$ ➡

$$\alpha > \alpha_{fl} = \arcsin\sqrt{B/B_{max}}$$

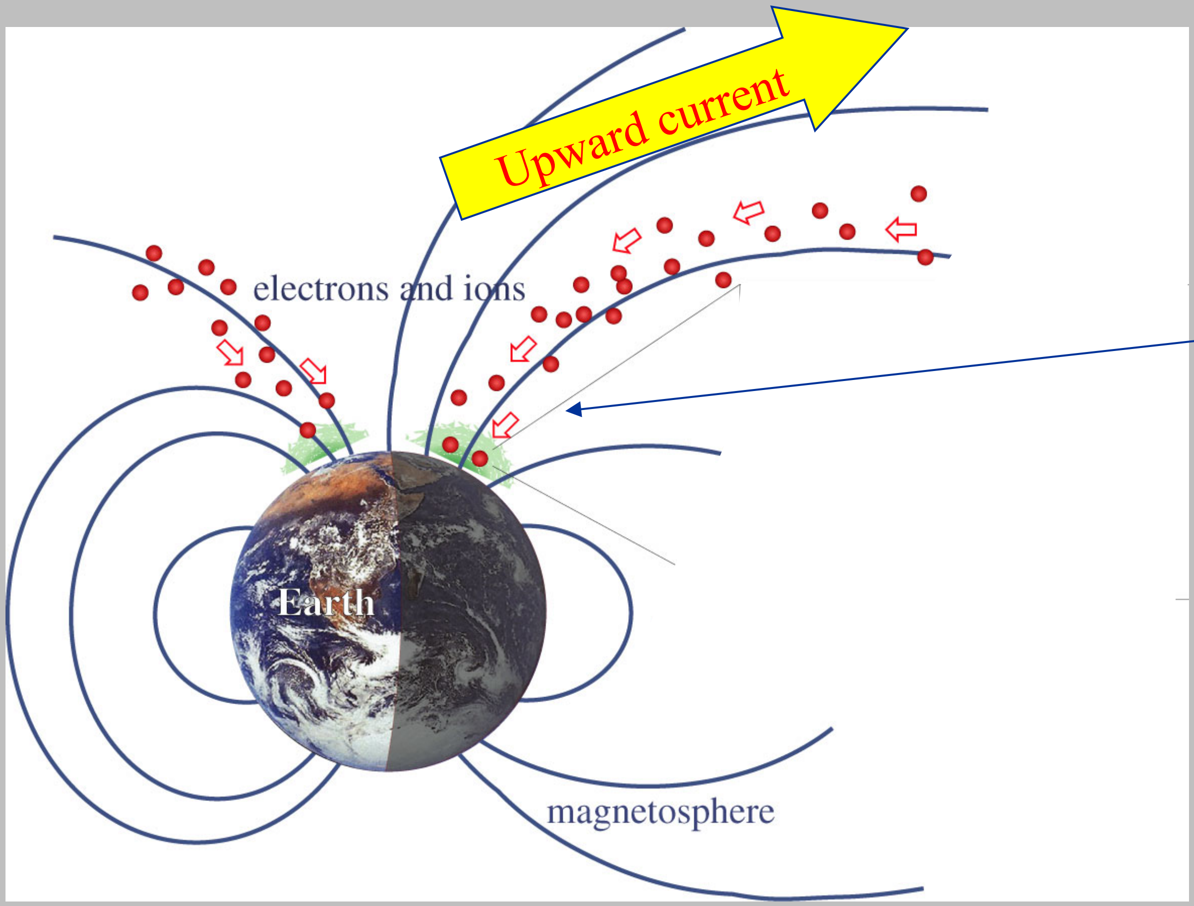
Particles in ***loss cone*** :

$$\alpha < \alpha_{fl}$$

The magnetic moment μ is an ***adiabatic invariant***.

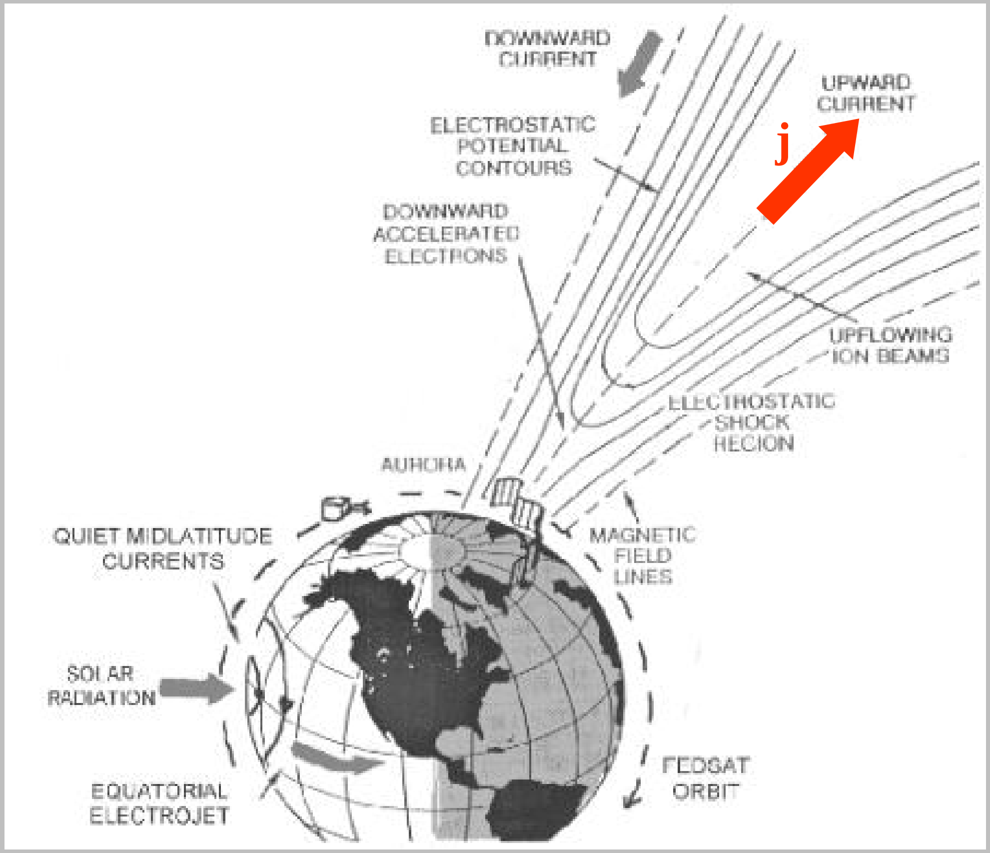
$$\mu = \frac{mv_\perp^2}{2B} = \frac{mv^2\sin^2\alpha}{2B}$$

Why particle acceleration?



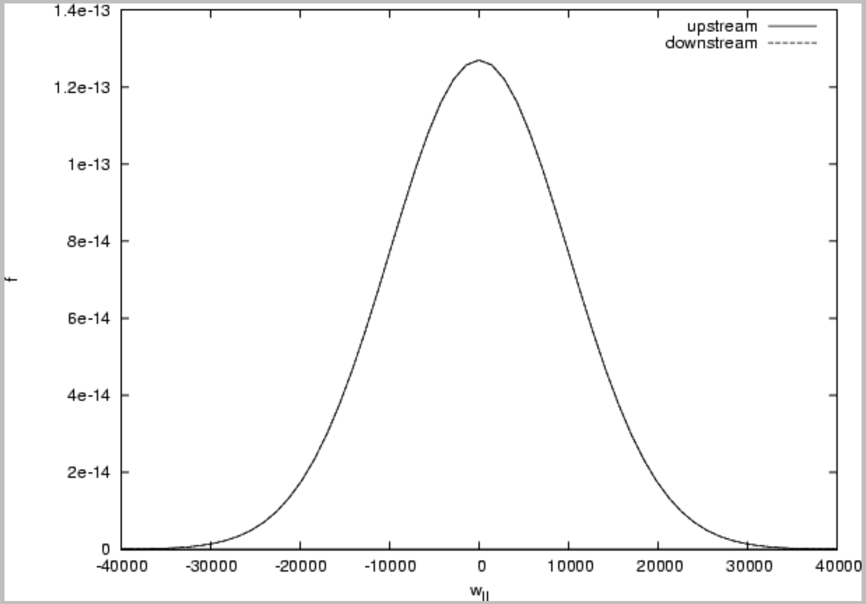
- The magnetosphere often seems to act as a current generator.
- The lower down you are on the field line, the more particles have been reflected by the magnetic mirror.
- At low altitudes there are not enough electrons to carry the current.

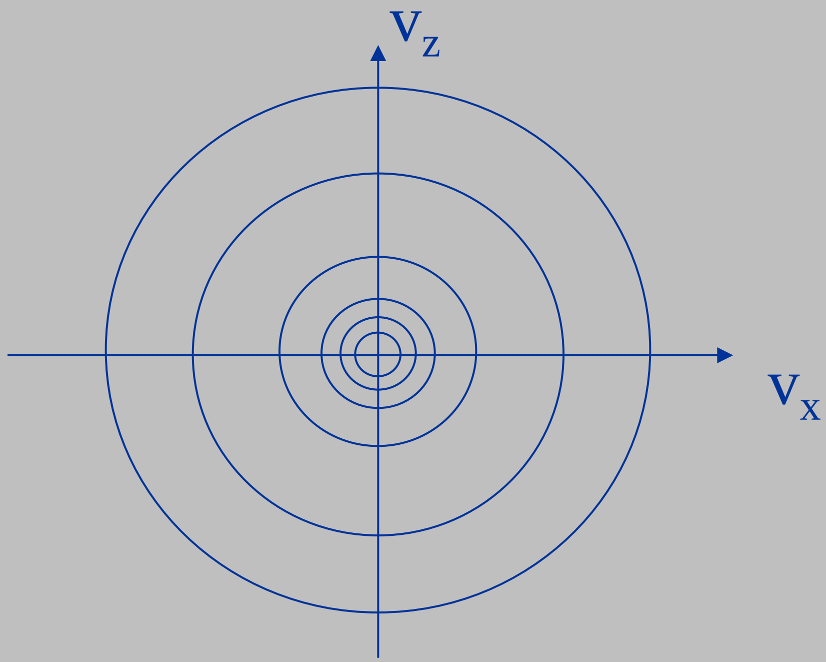
Why particle acceleration?



- Electrons are accelerated downwards by upward E-field.
- This increases the pitch-angle of the electrons, and more electrons can reach the ionosphere, where the current can be closed.

Distribution function

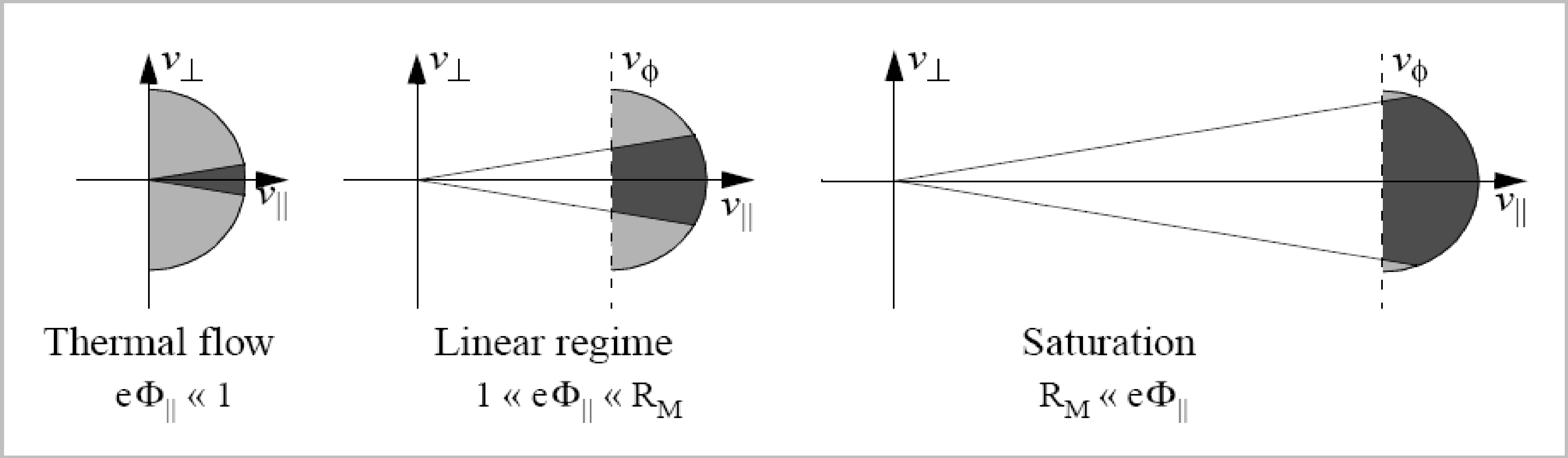




Example:
Maxwellian distribution

$$f = \frac{n}{\sqrt{(2\pi RT)^3}} \exp\left(-\frac{m(v_x^2+v_y^2+v_z^2)}{2kT}\right)$$

Why particle acceleration?



- Electrons are accelerated downwards by upward E-field.
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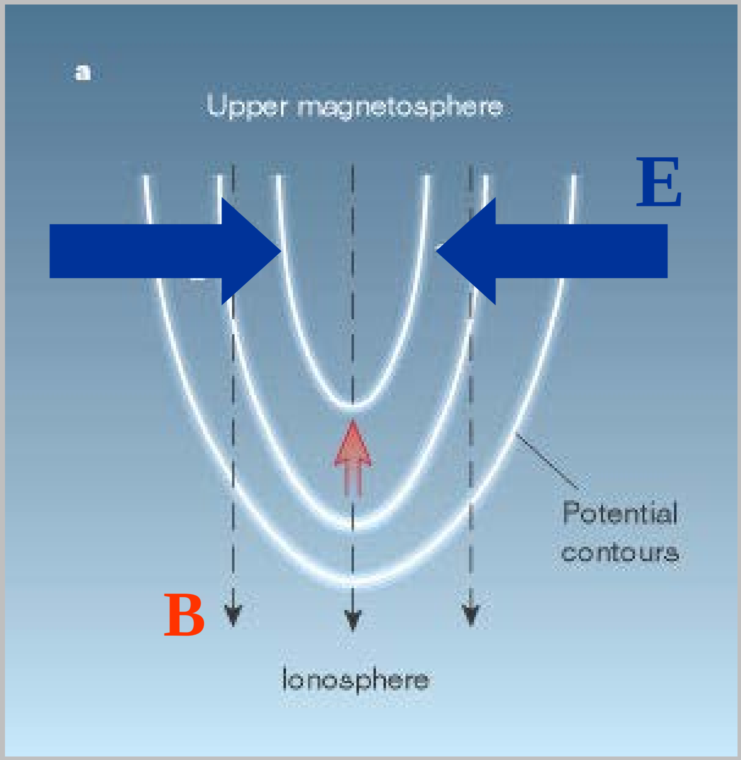
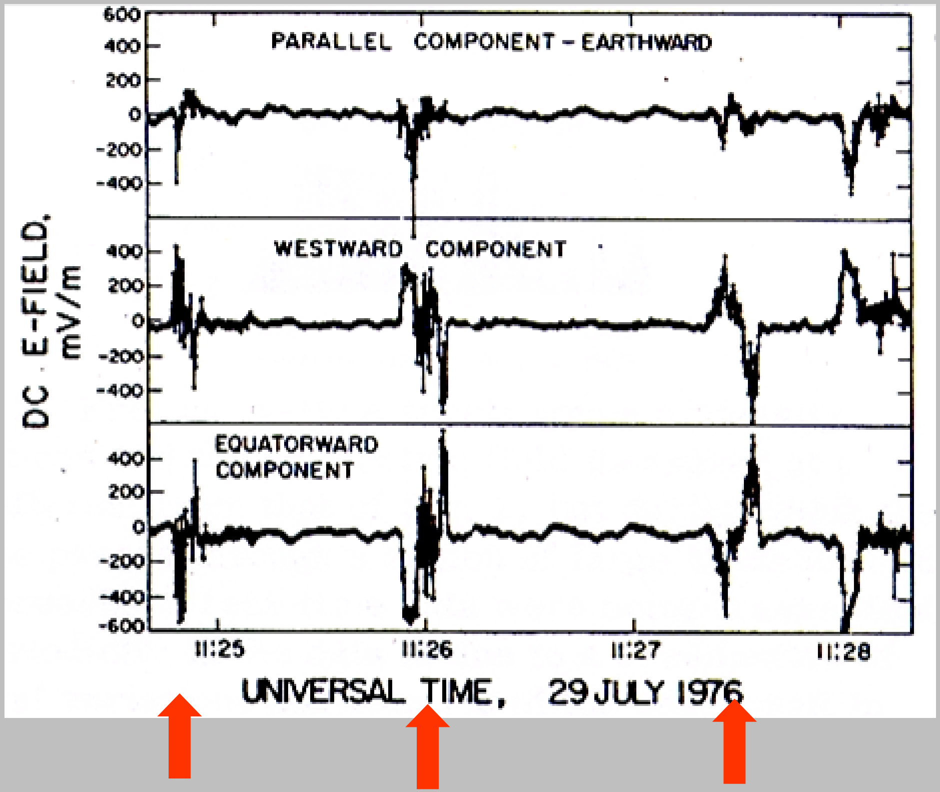
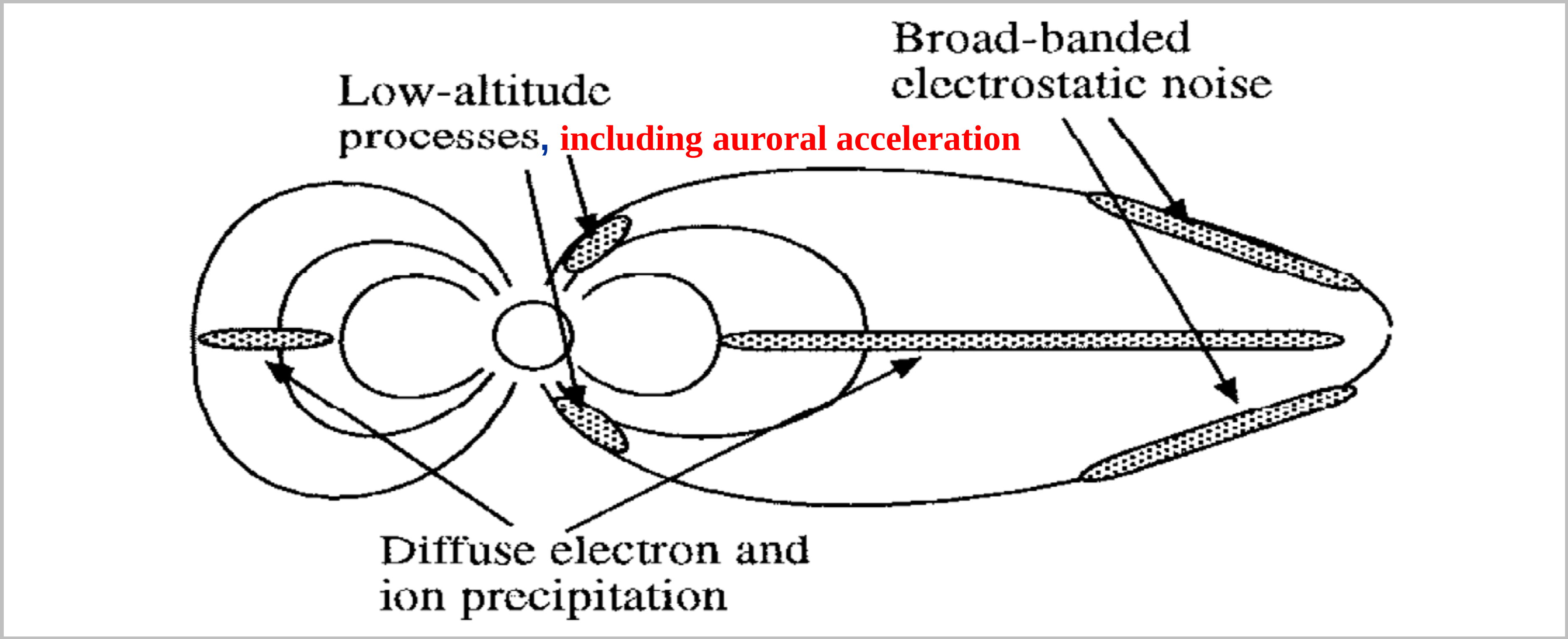
Satellite signatures of U potential





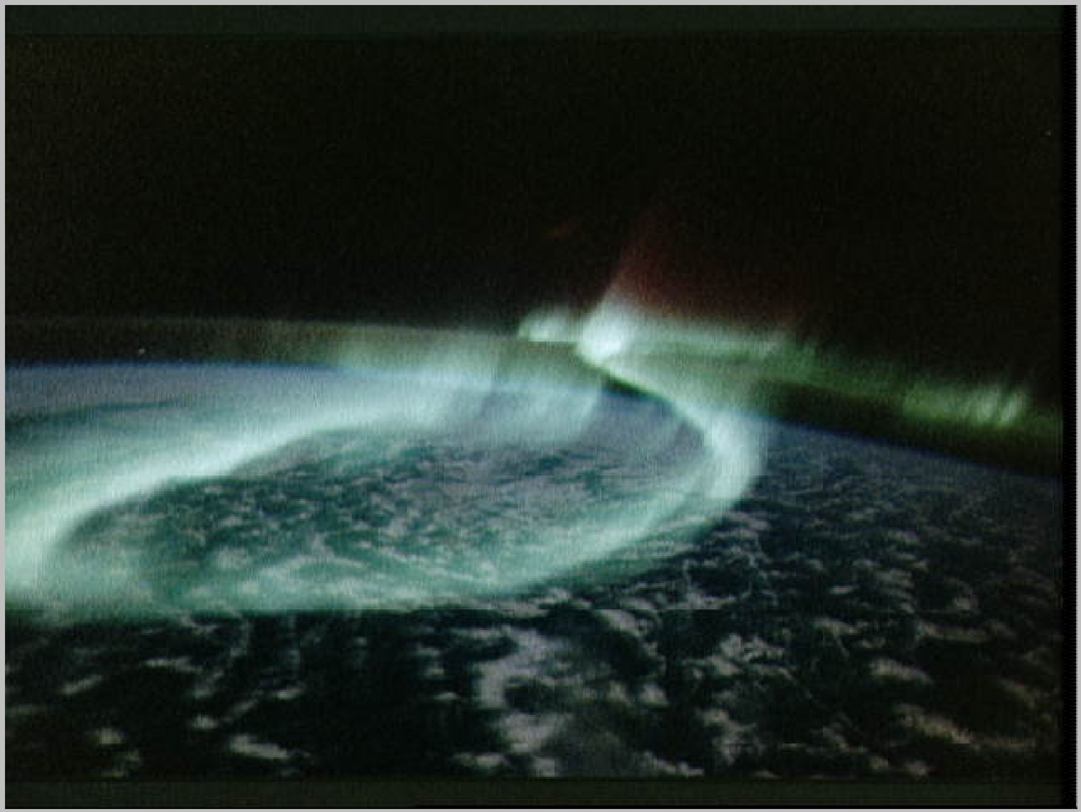
Measurements made by the ISEE satellite (Mozer et al., 1977)

Acceleration regions



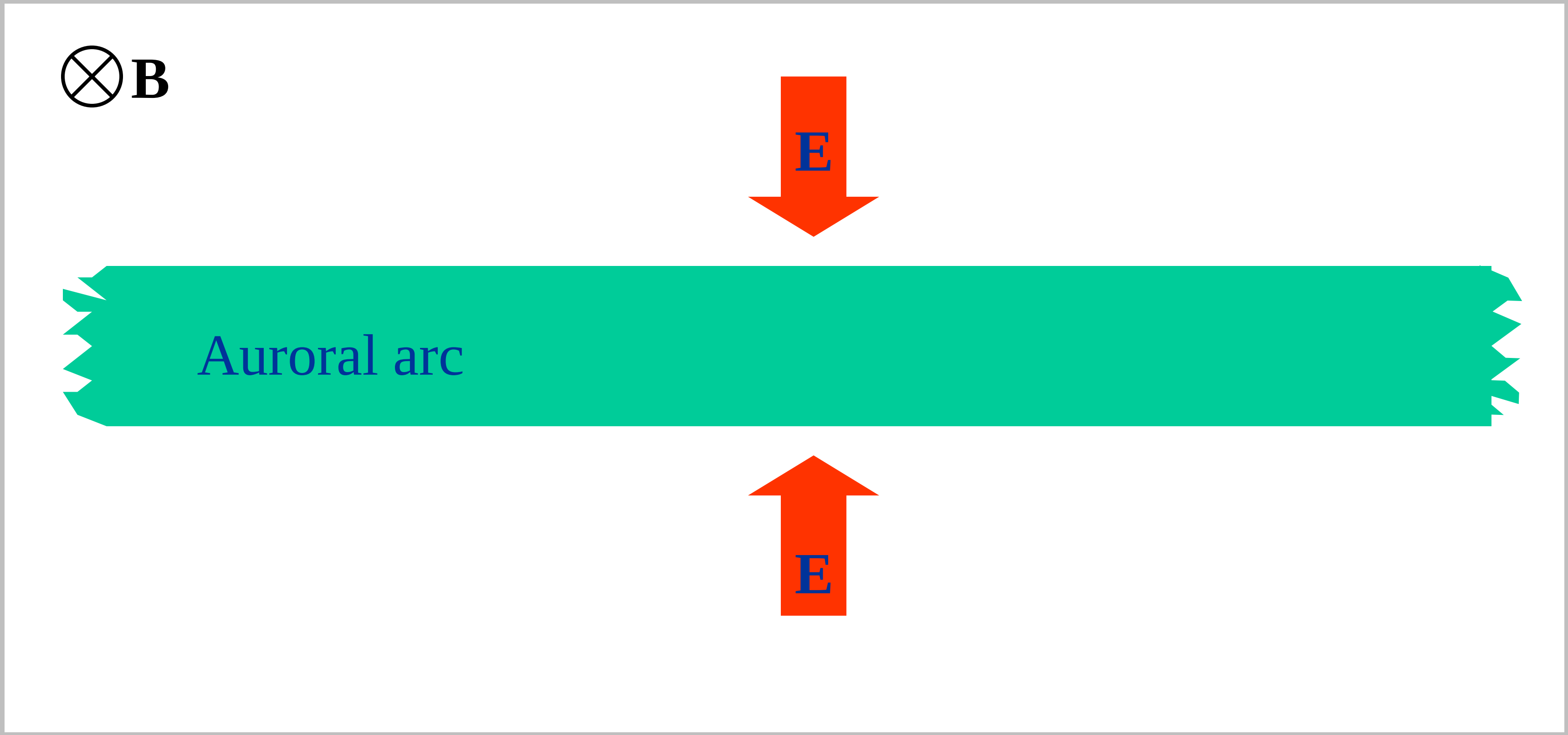
Auroral acceleration region typically situated at altitude of 1-3 R_E

Auroral spirals

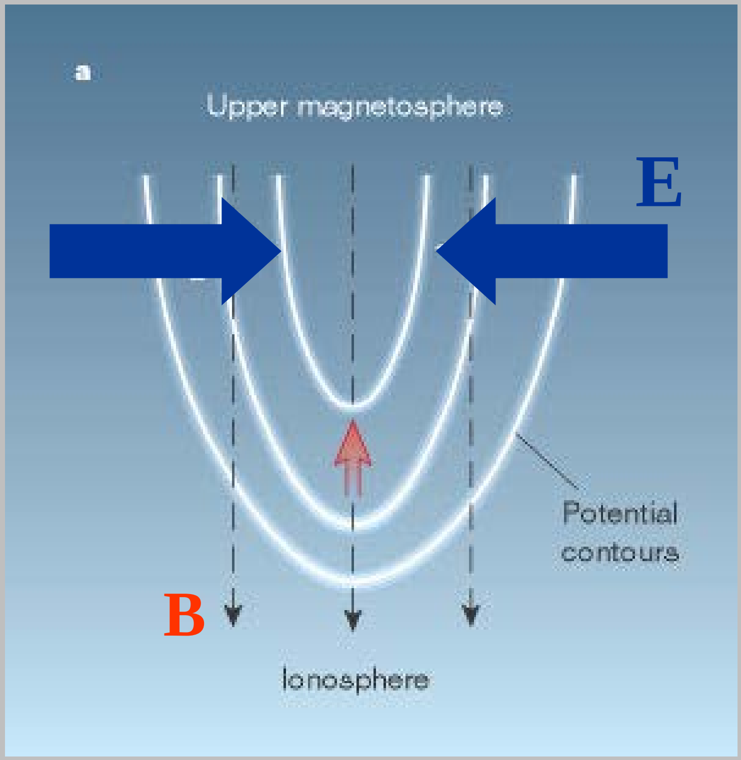


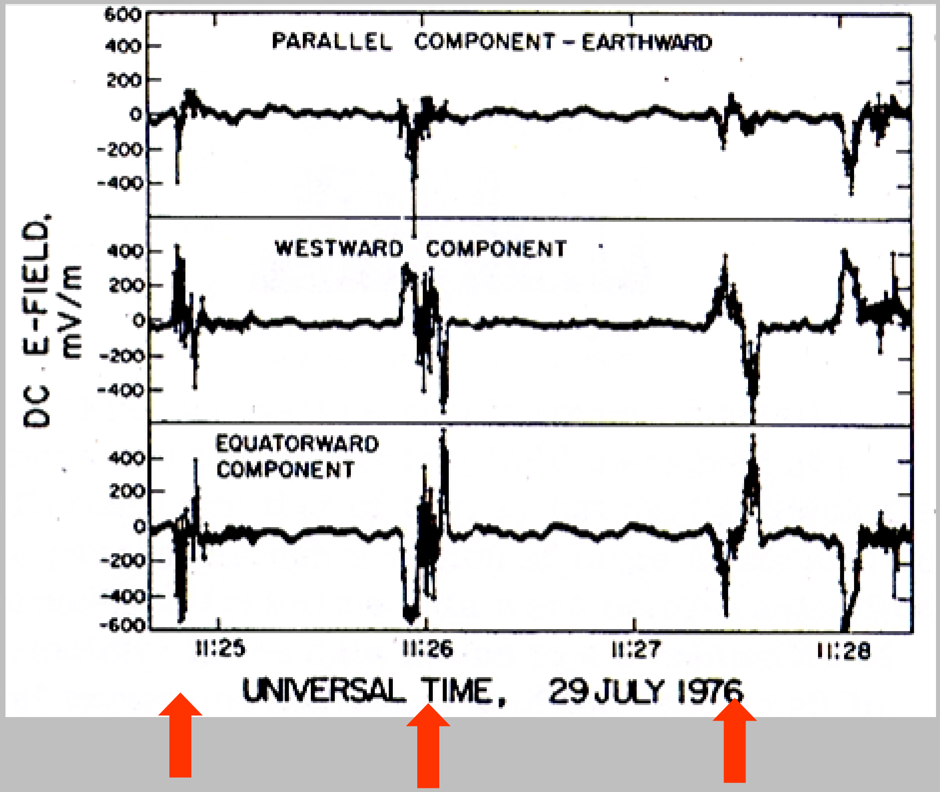
Develop when arcs become unstable

Spirals – Kelvin-Helmholz instability



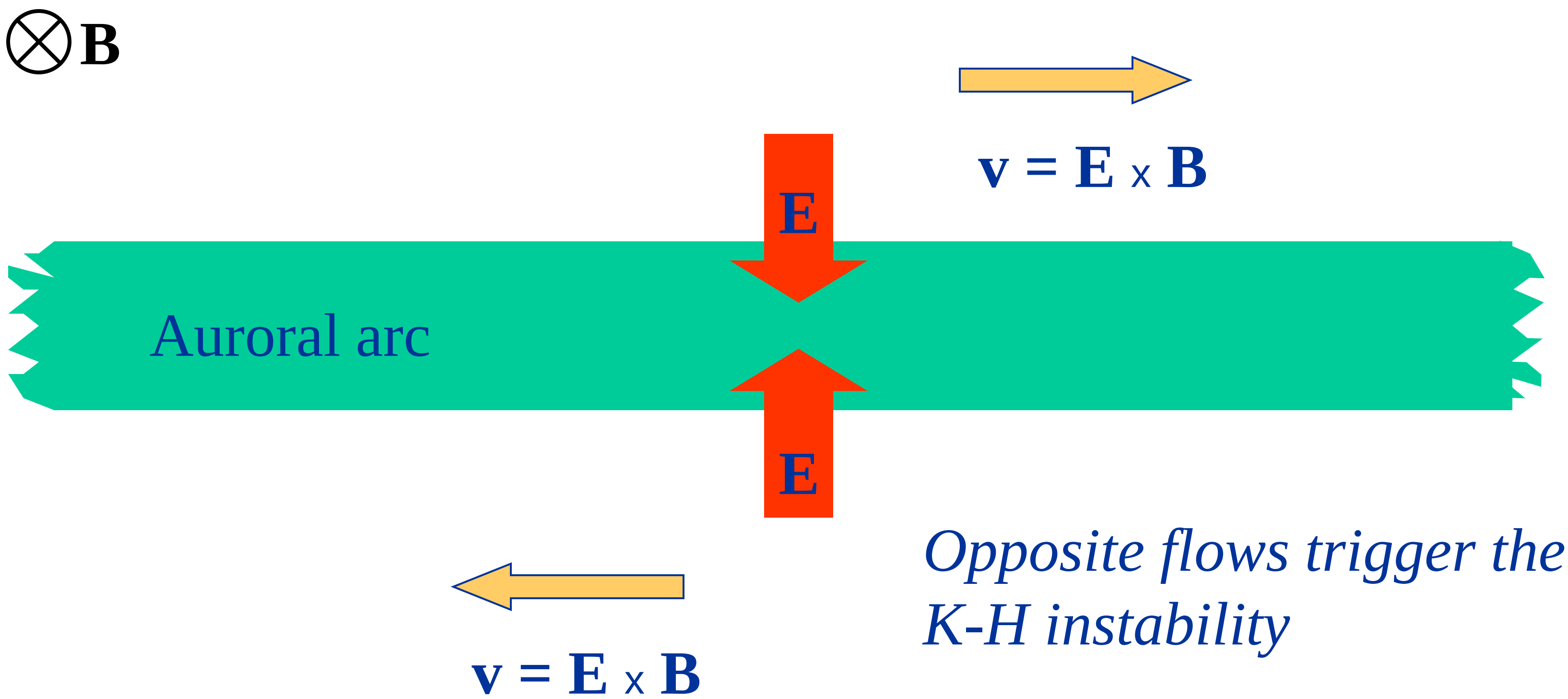
Satellite signatures of U potential





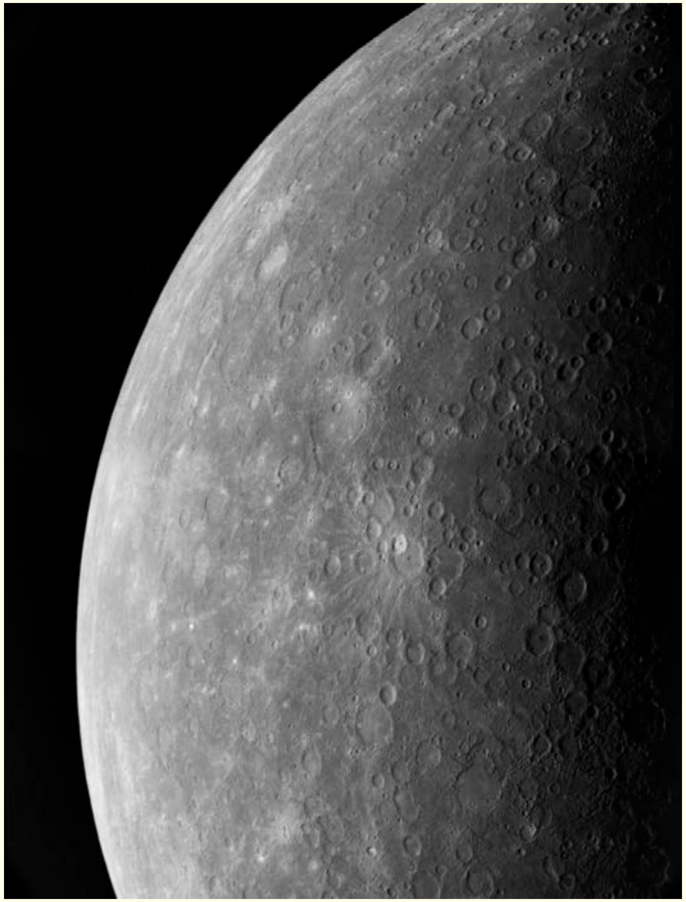
Measurements made by the ISEE satellite (Mozer et al., 1977)

Spirals – Kelvin-Helmholz instability



Opposite flows trigger the K-H instability

Mercury



- No atmosphere
- X-ray aurora???
  Can possibly be created by electrons colliding directly with the planetary surface and lose their energy in one single collision.

Jupiter aurora

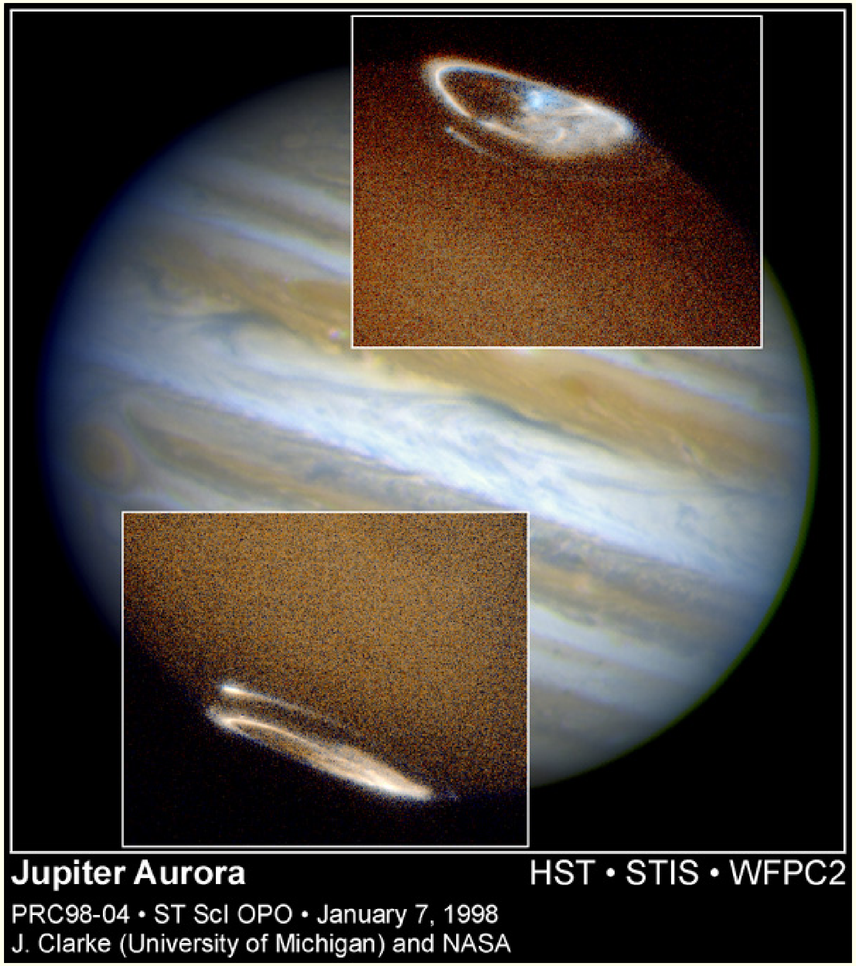
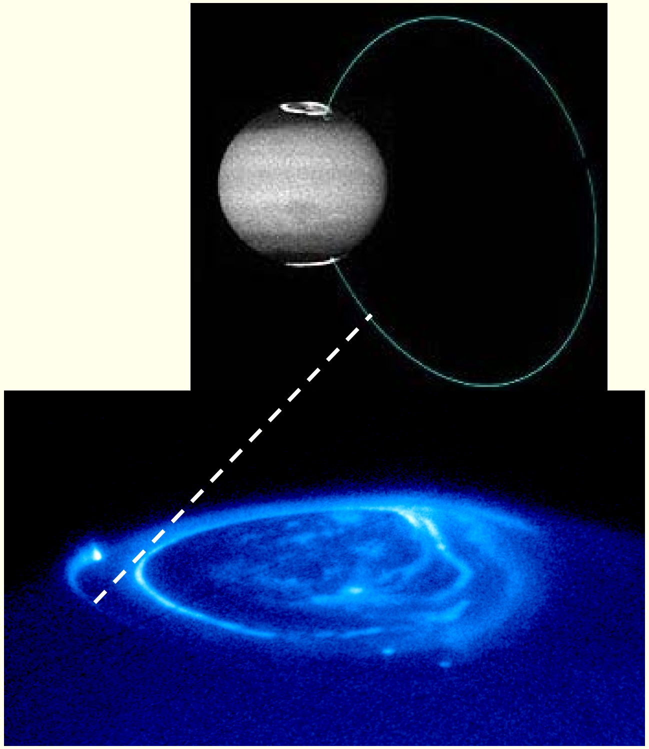


Foto från Hubble Space Telescope

- Jupiter's aurora has a power of ~1000 TW *(compare Earth: ~100 GW, nuclear power plant: ~1 GW)*
- Note the "extra" oval on Io's flux tube!



Jupiter and Io

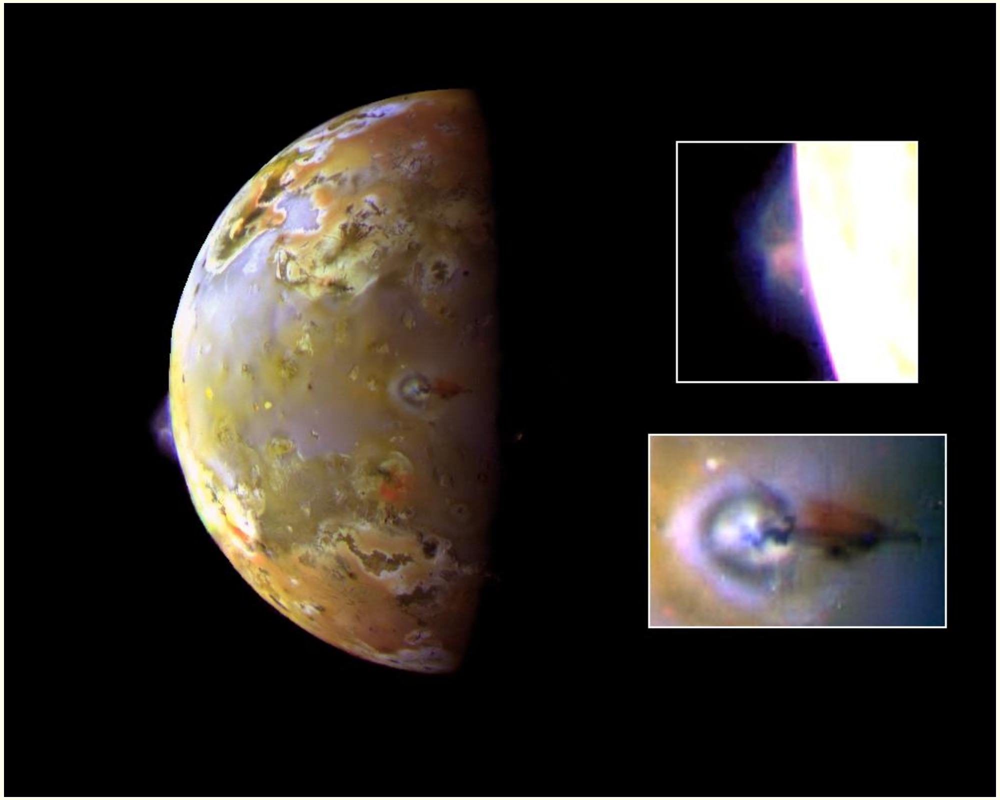
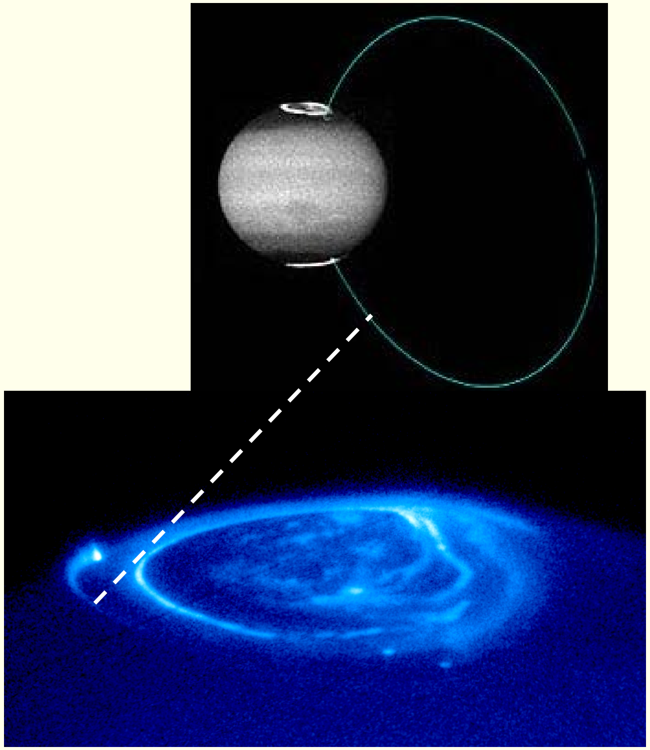


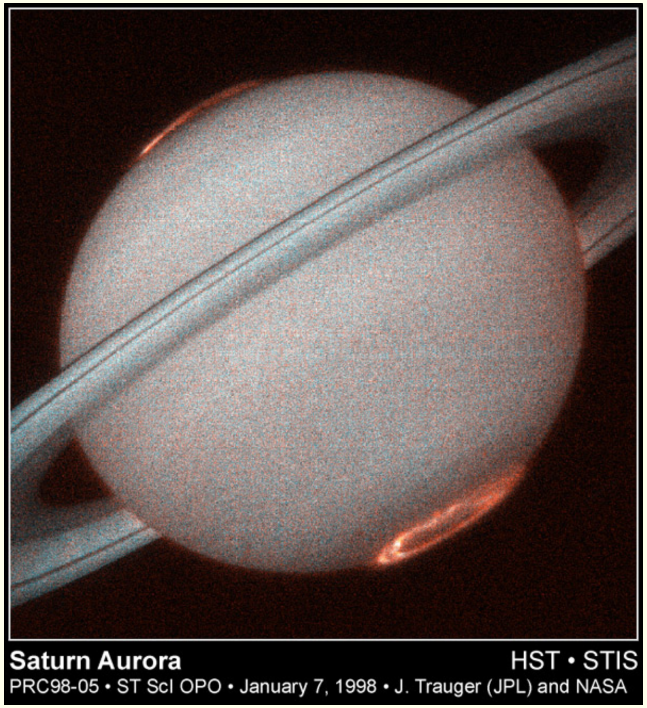
Photo from rymdsonden Galileo



The Jupiter moon Io is very volcanically active, and deposes large amounts of dust and gas in Jupiter's magnetosphere. This is ionized by the sunlight, and the charged plasma partícles follow Jupiter's magnetic field lines towards the atmosphere and cause auroral emissions.

Aurora of the other planets

Saturn



Saturnus' aurora: not noticeably different from Jupiter's, but much weaker. (Total power about the same as Earth's aurora.)

Uranus: *Auora detected in UV. Probably associated with Uranus' ring current/radiotion belts and not very dynamic.*

Neptunus: *weak UV aurora detected.*

Mars, Venus: *No aurora.*

Prerequisites for…

Life

- Energy source (sun)
- Atmosphere
- Magnetic field
- Water

Aurora

- Energy source (sun)
- Atmosphere
- Magnetic field

At what planets do you expect aurora to exist?

Blue — Earth, Mercury, Jupiter, Saturn

Yellow — Earth, Venus, Jupiter, Saturn, Uranus, Neptune

Green — Earth, Mars, Jupiter, Saturn, Uranus, Neptune

Red — Earth, Jupiter, Saturn, Uranus, Neptune

What do we need to have an aurora?

- Magnetic field (to guide the plasma particles towards the planet)
- Atmosphere (to create emissions)

At what planets do you expect aurora to exist?

Red

Earth, Jupiter, Saturn, Uranus, Neptune

On space weather and viewing aurora

Some space weather sites

http://spaceweather.com/

http://www.esa-spaceweather.net/

http://sunearthday.nasa.gov/swac/

http://www.noaawatch.gov/themes/space.php

http://www.windows2universe.org/spaceweather/more_details.html

Kiruna

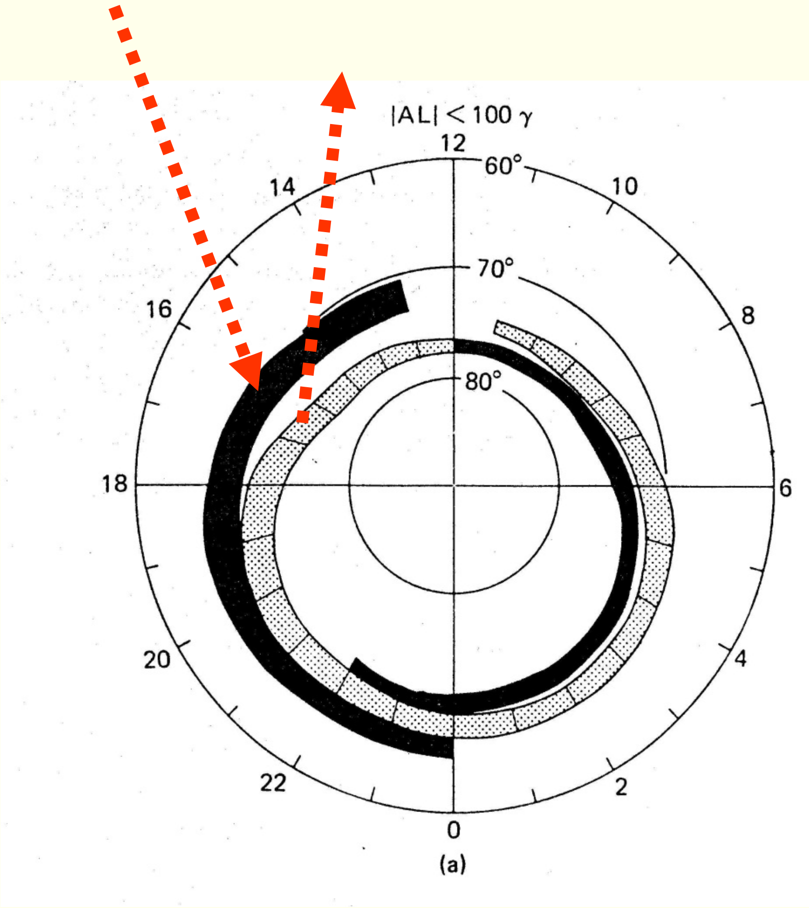
Kiruna all-sky camera:
http://www.irf.se/allsky/rtasc.php

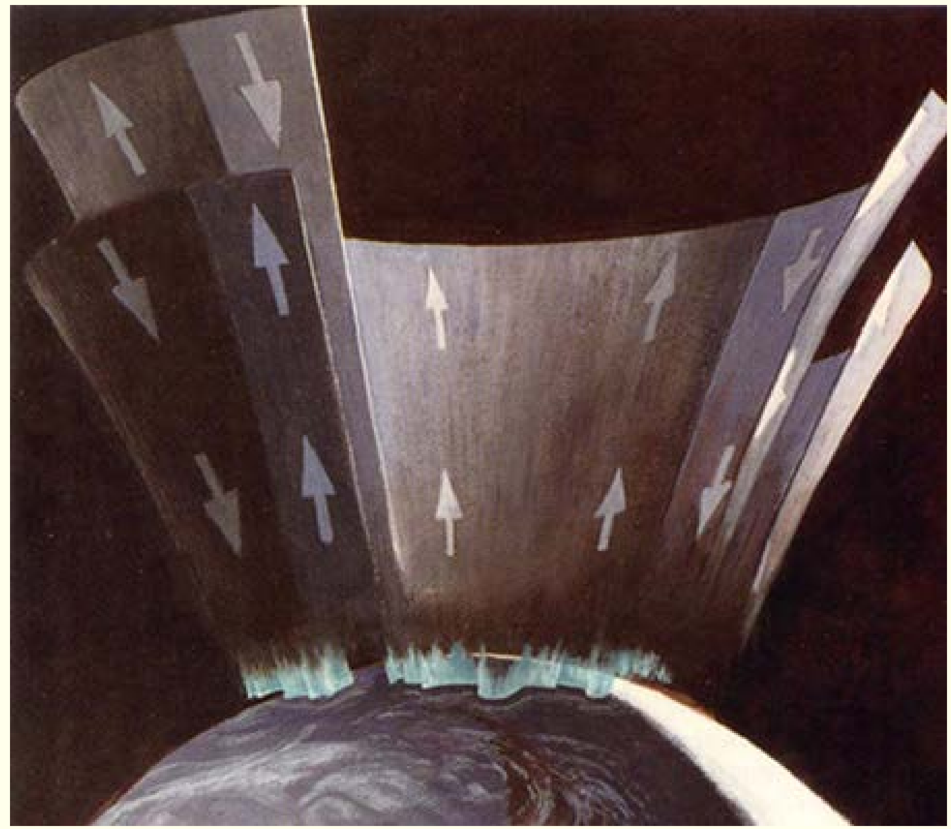
http://sunearthday.nasa.gov/swac/tutorials/aur_kiruna.php

Forecasts:
http://flare.lund.irf.se/rwc/aurora/

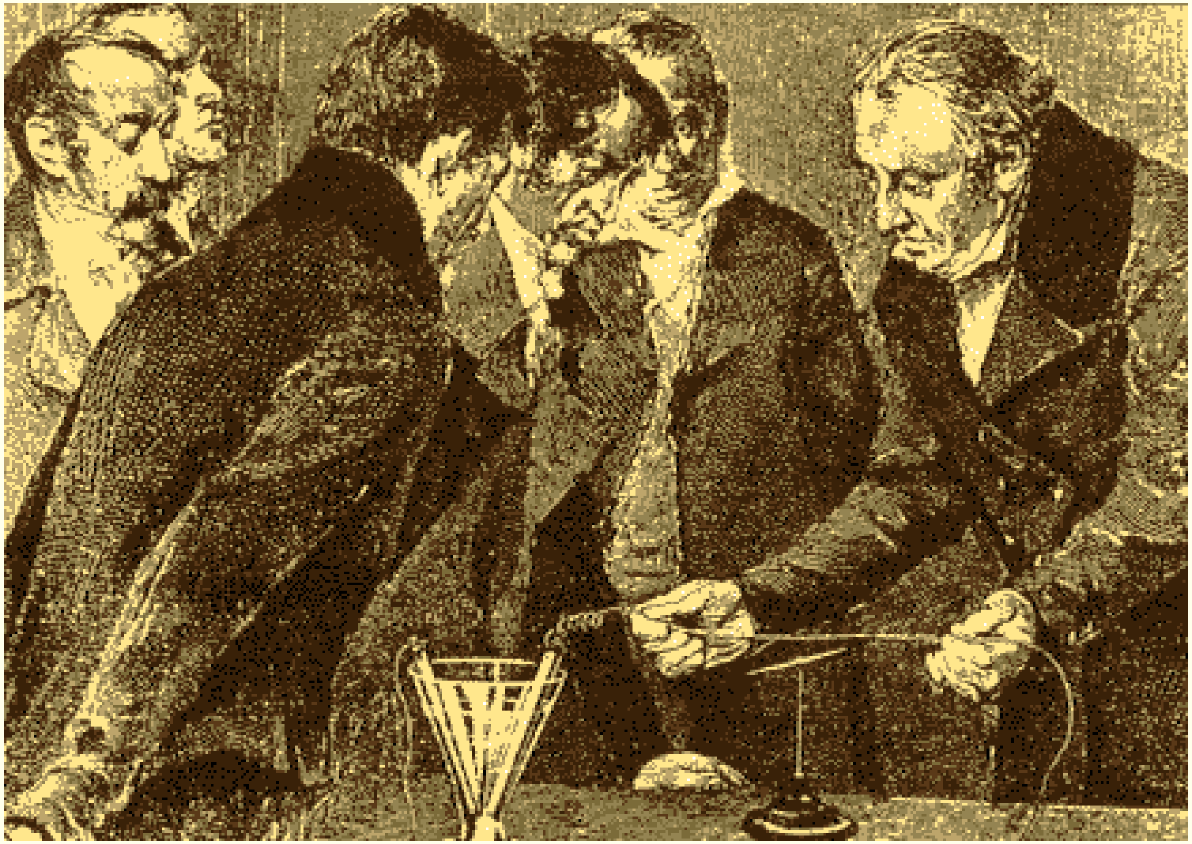
http://www.irf.se/Observatory/?link[All-skycamera]=Aurora_sp_statistics

Birkeland currents in the auroral oval

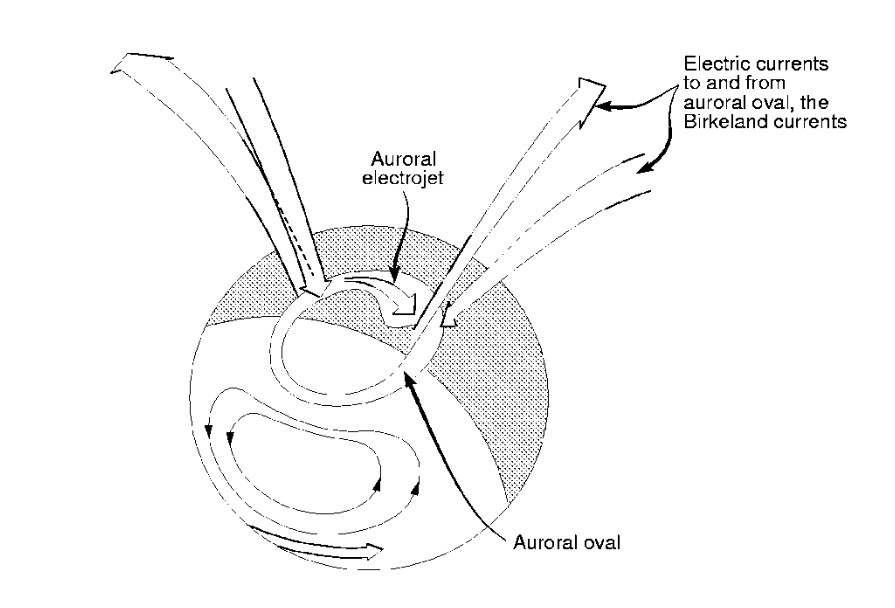




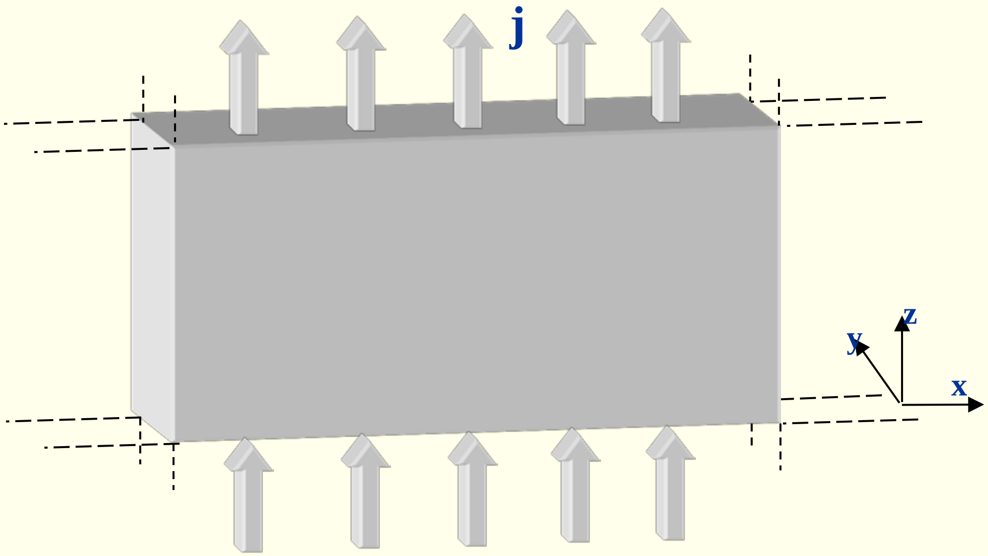
How can you measure currents in space?



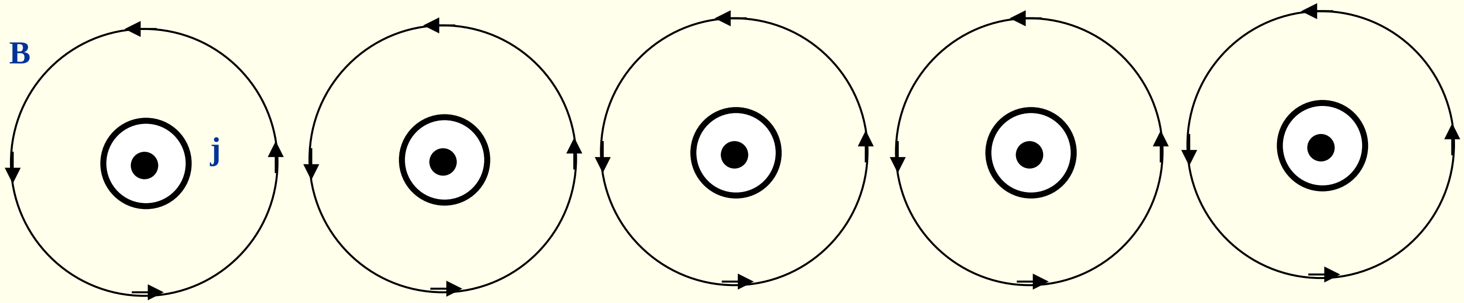
Current sheet approximation



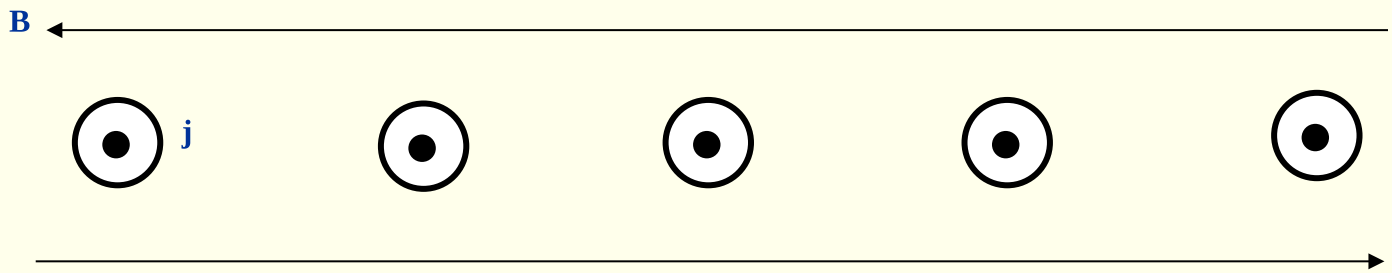
Approximate currents by thin current sheets with infinite size in the x- och z-directions.



Current sheet approximation

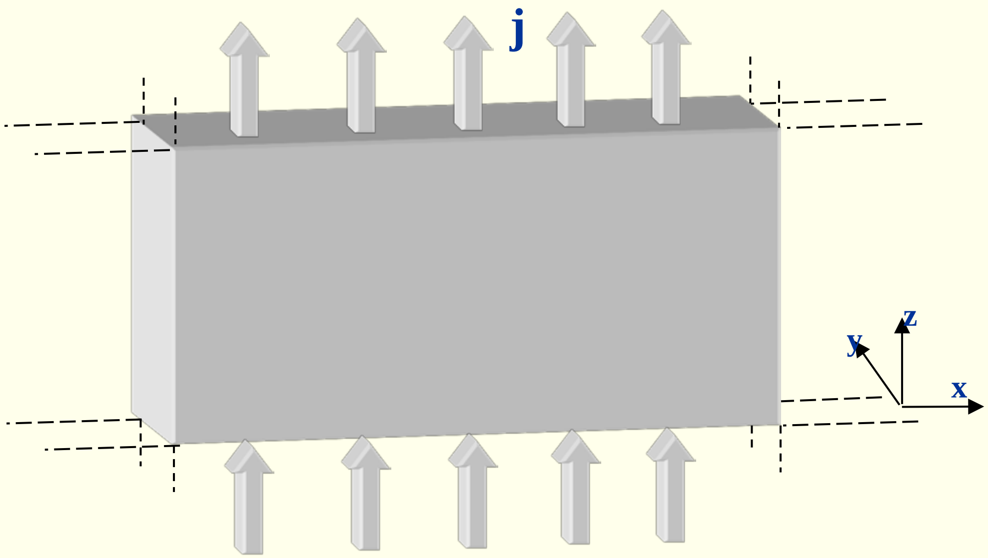


What will the magnetic field around such a current configuration be? Start by approximating with line currents to get a qualitative picture.



The closer you place the line currents, the more the magnetic fields between the line currents will cancel

Current sheet approximation and Ampére's law



$$\left(\frac{\partial B_z}{\partial y}-\frac{\partial B_y}{\partial z},\ \frac{\partial B_x}{\partial z}-\frac{\partial B_z}{\partial x},\ \frac{\partial B_y}{\partial x}-\frac{\partial B_x}{\partial y}\right)=\mu_0\left(j_x,\ j_y,\ j_z\right)$$

But $\frac{\partial}{\partial x}=0$ and $\frac{\partial}{\partial z}=0$

$$\left(\frac{\partial B_z}{\partial y},\ 0,\ -\frac{\partial B_x}{\partial y}\right)=\mu_0\left(0,\ 0,\ j_z\right)$$

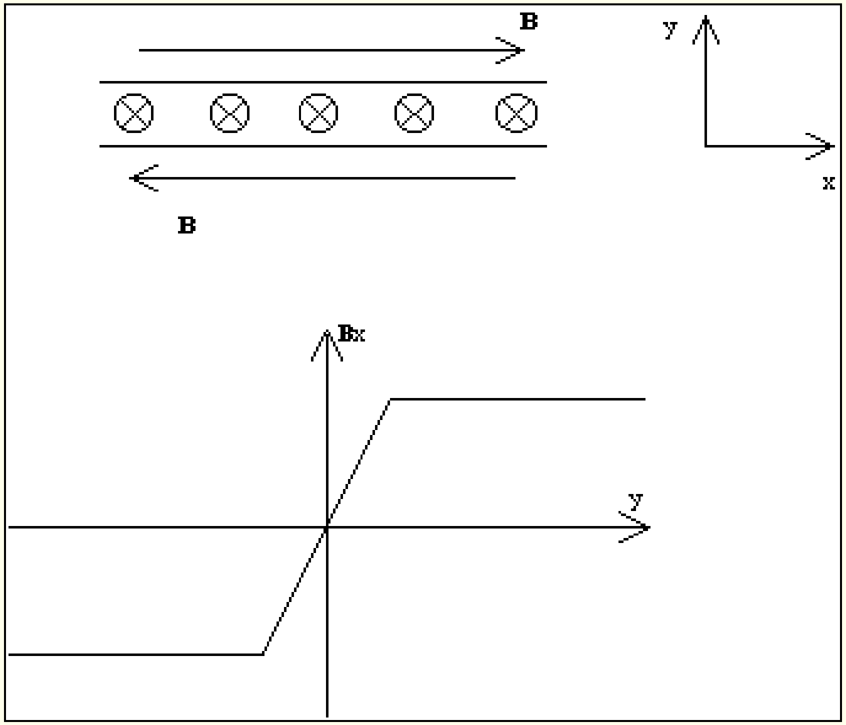
eller

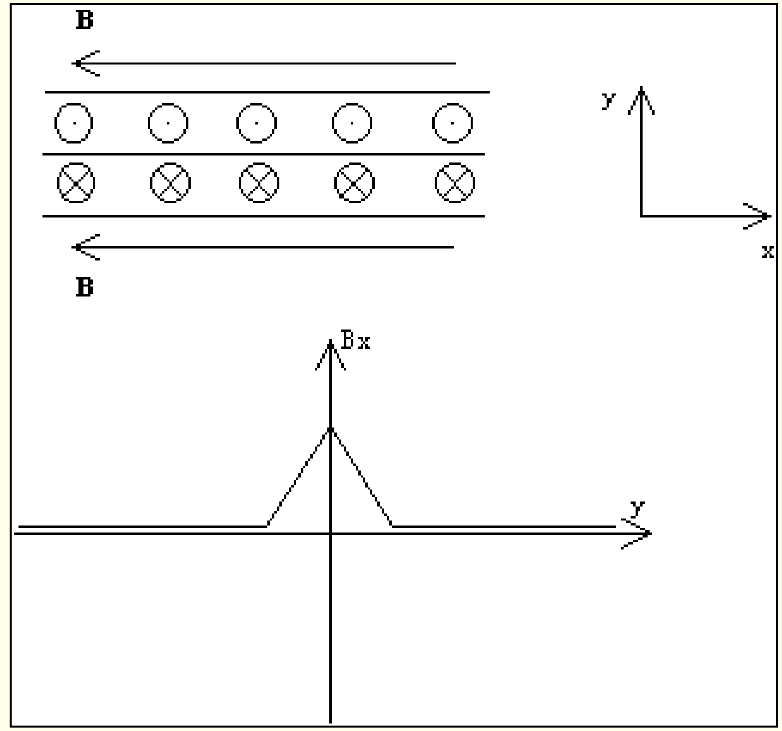
Ampére's law (no time dependence):

$$\nabla\times\mathbf{B}=\mu_0\mathbf{j}$$

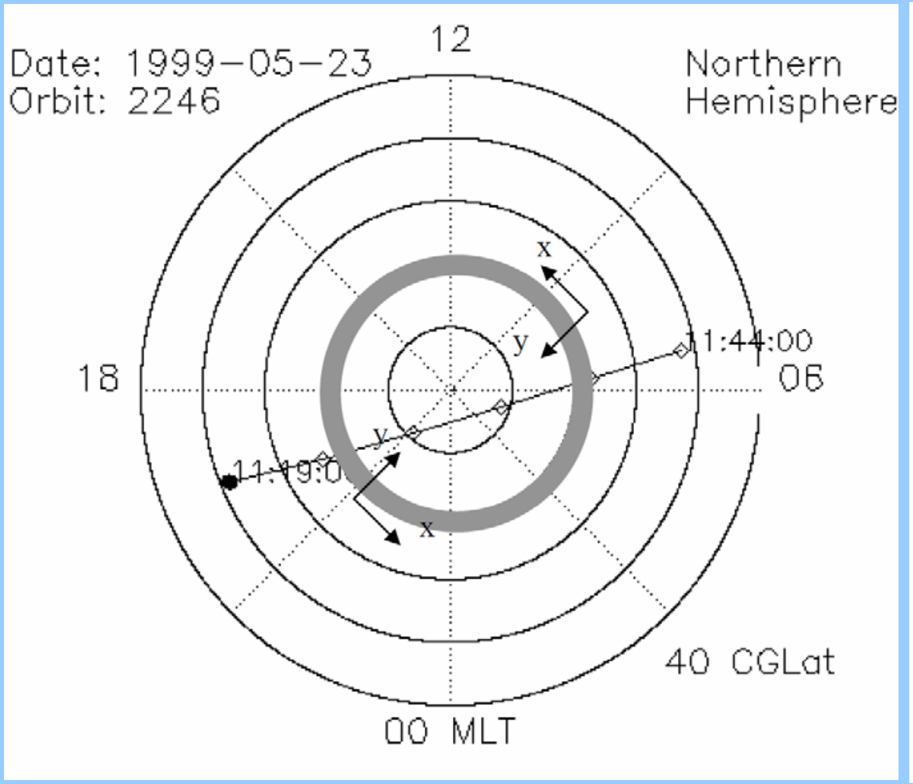
$$j_z=-\frac{1}{\mu_0}\frac{\partial B_x}{\partial y}$$

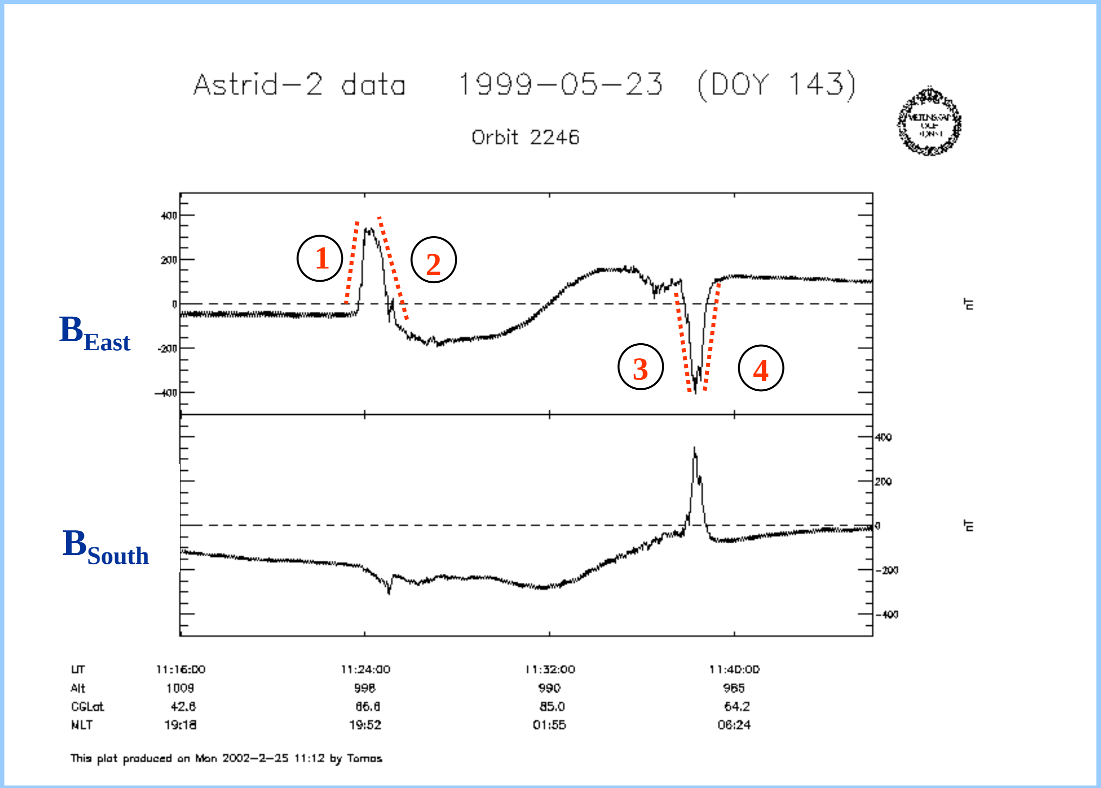
Current sheet - example





$$j_z = -\frac{1}{\mu_0}\frac{\partial B_x}{\partial y}$$

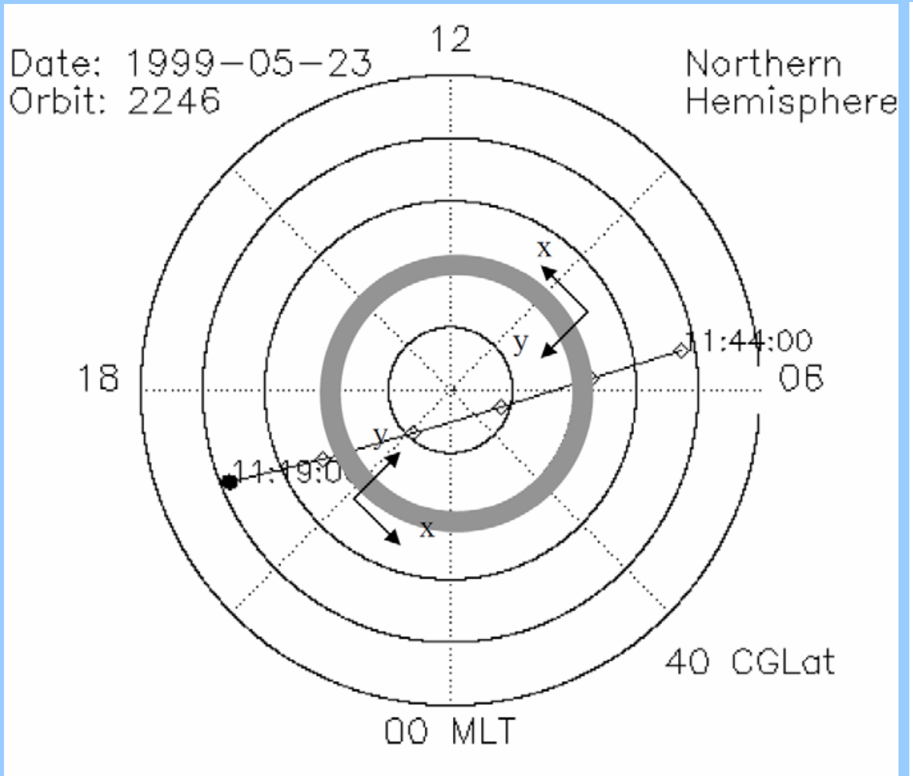


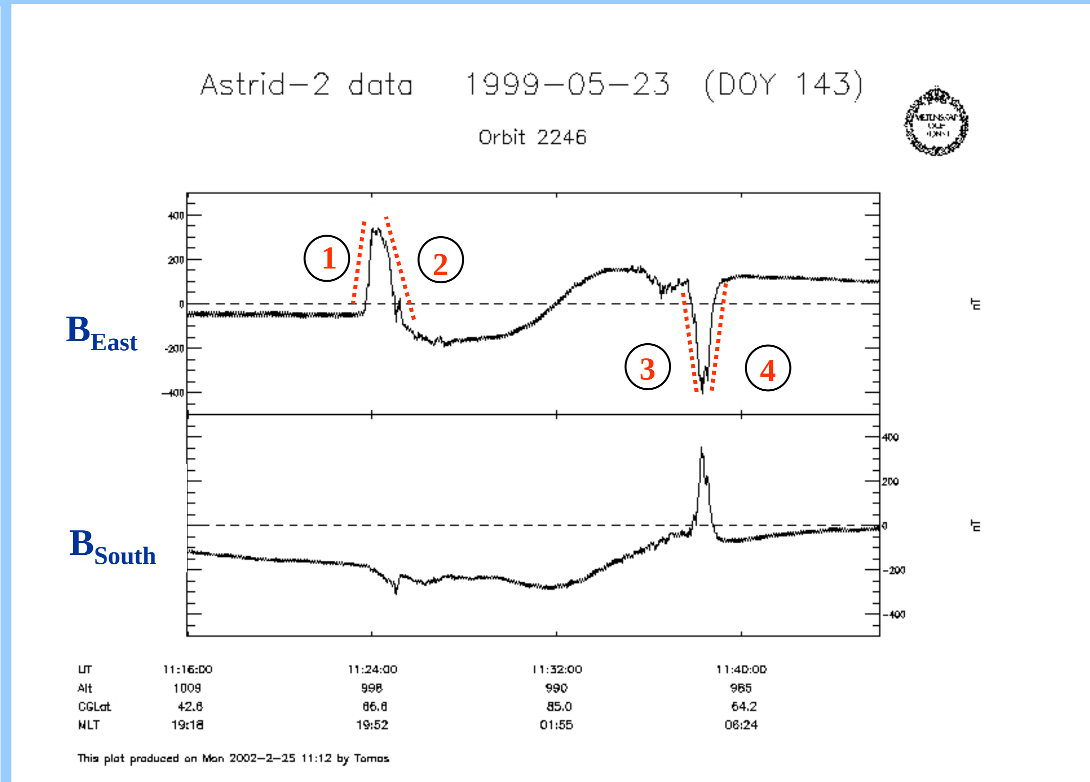


$$j_z = -\frac{1}{\mu_0}\frac{\partial B_x}{\partial y}$$

What is the direction of the current in current sheet 1?

Blue	Into the ionosphere
Red	Out of the ionosphere





What is the direction of the current in current sheet 1?

$$j_z = -\frac{1}{\mu_0}\frac{\partial B_x}{\partial y}$$

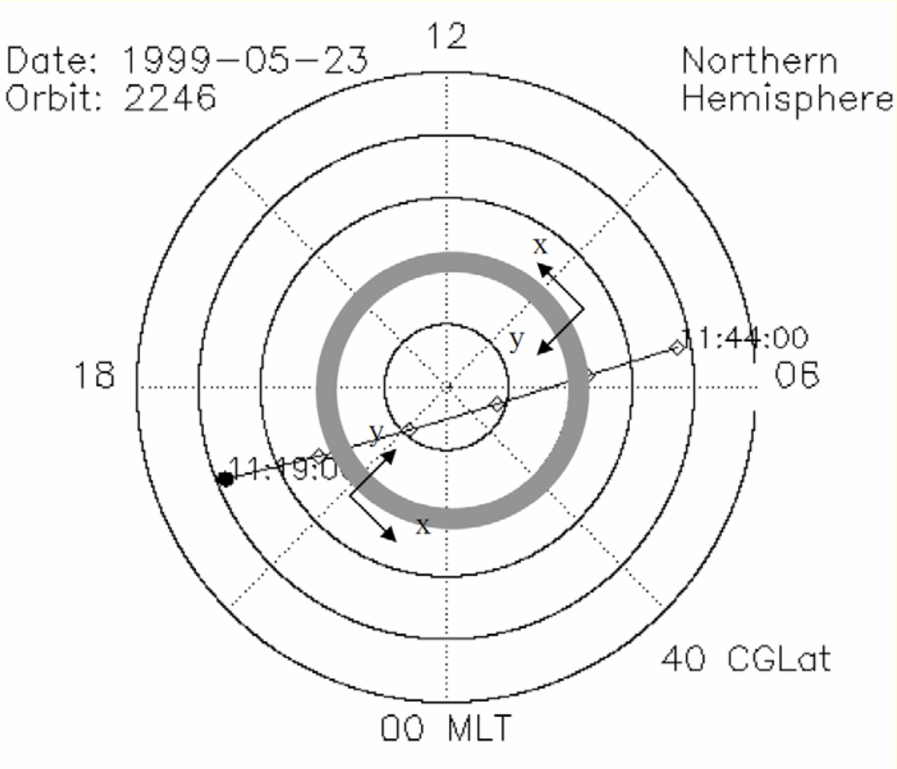
$$\frac{\partial B_x}{\partial y} = \frac{\partial B_{East}}{\partial y} > 0$$

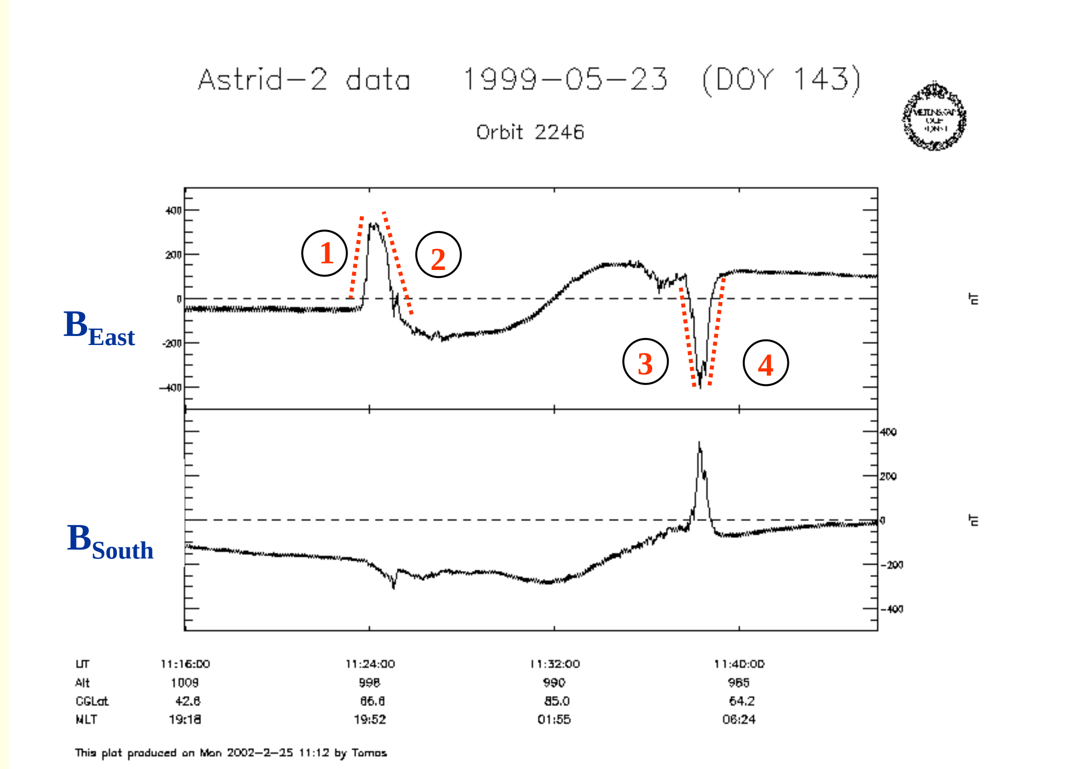
Blue

Into the ionosphere

$$\Rightarrow$$

$$j_z < 0$$

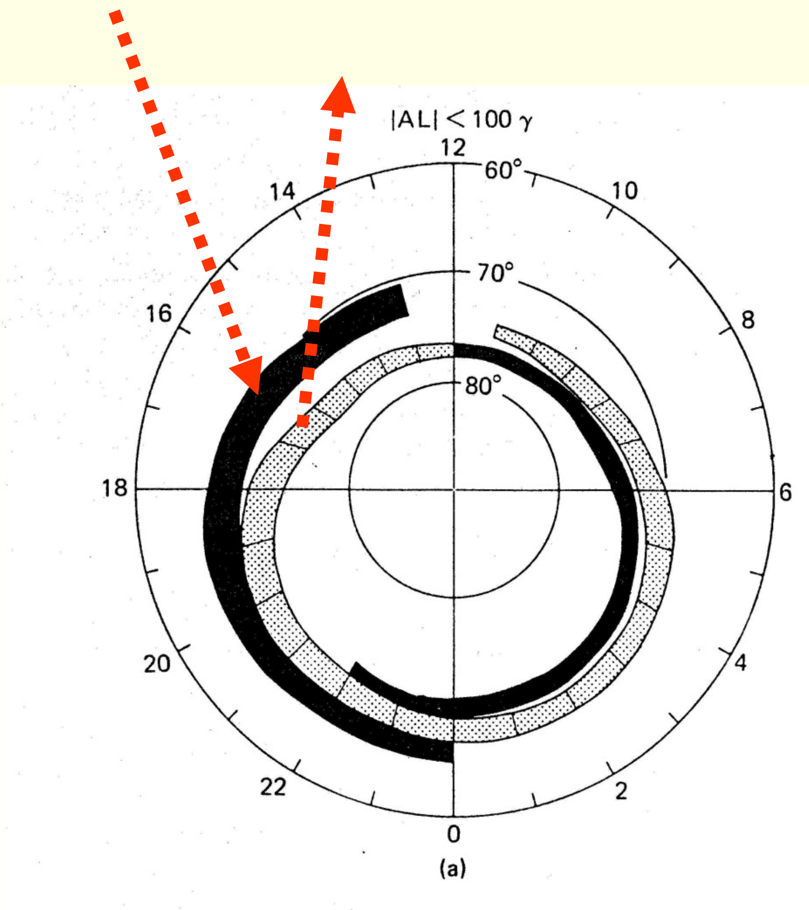


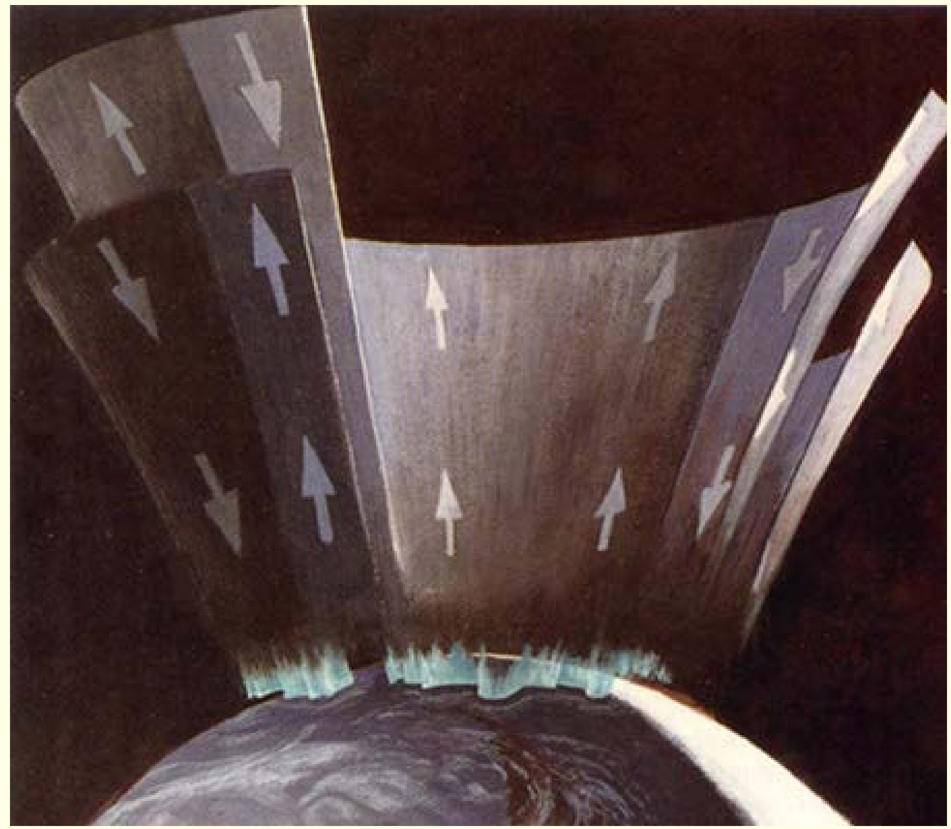


$$j_z = -\frac{1}{\mu_0}\frac{\partial B_x}{\partial y}$$

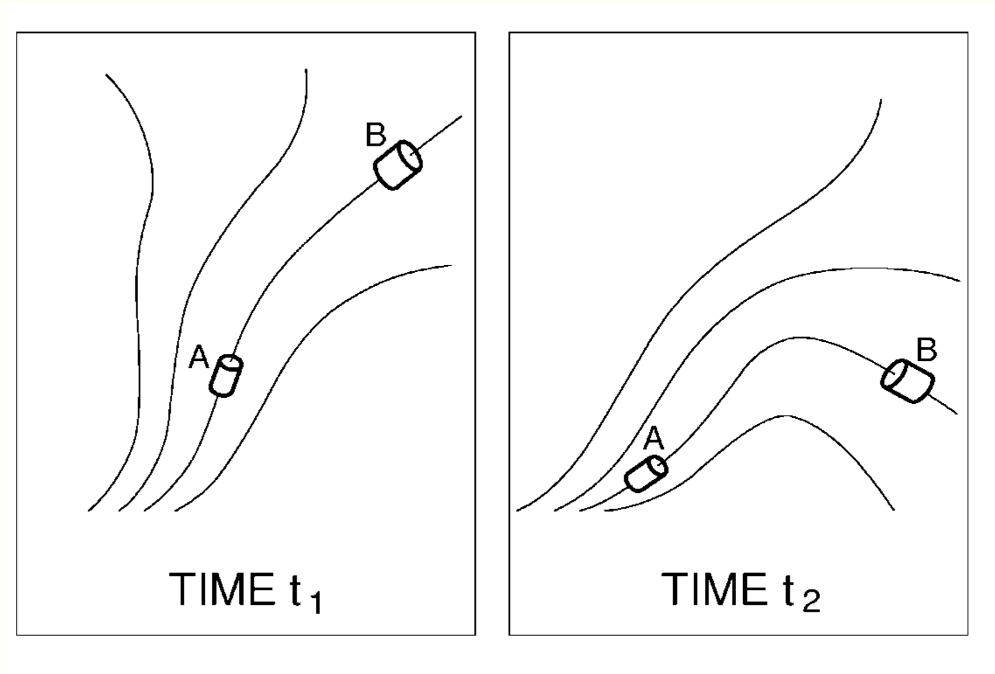
1)	$\frac{\partial B_x}{\partial y} > 0$	$\Rightarrow$	$j_z < 0$	Into the ionosphere
2)	$\frac{\partial B_x}{\partial y} < 0$	$\Rightarrow$	$j_z > 0$	Out of the ionosphere
3)	$\frac{\partial B_x}{\partial y} > 0$	$\Rightarrow$	$j_z < 0$	Into the ionosphere
4)	$\frac{\partial B_x}{\partial y} < 0$	$\Rightarrow$	$j_z > 0$	Out of the ionosphere

Birkeland currents in the auroral oval





Frozen in magnetic field lines



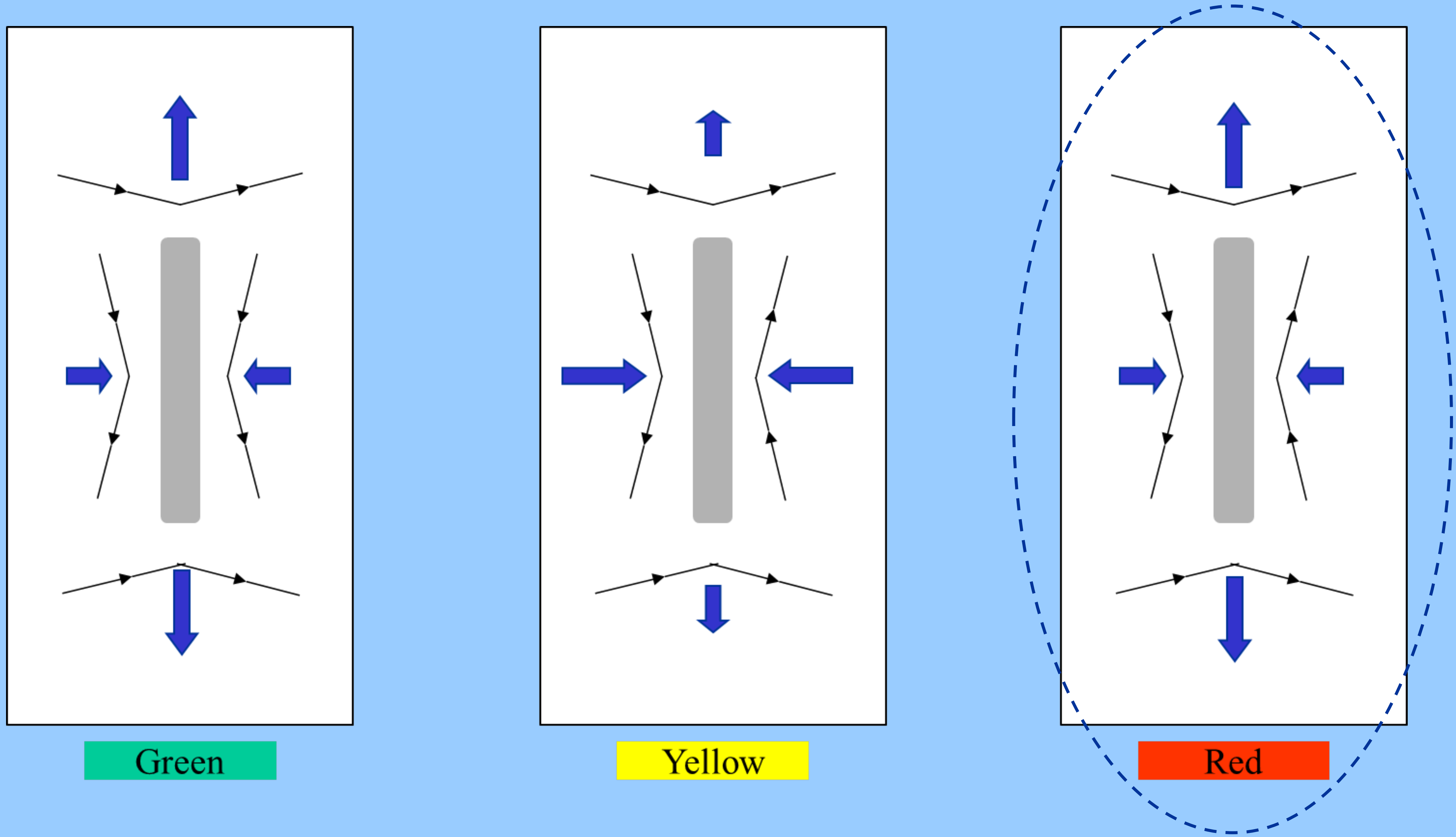
In fluid description of plasma two plasma elements that are connected by a common magnetic field line at time t_1 will be so at any other time t_2.

This applies if the magnetic Reynolds number is large:

$$R_m = \mu_0 \sigma l_c v_c >> 1$$

An example of the collective behaviour of plasmas.

Magnetic reconnection

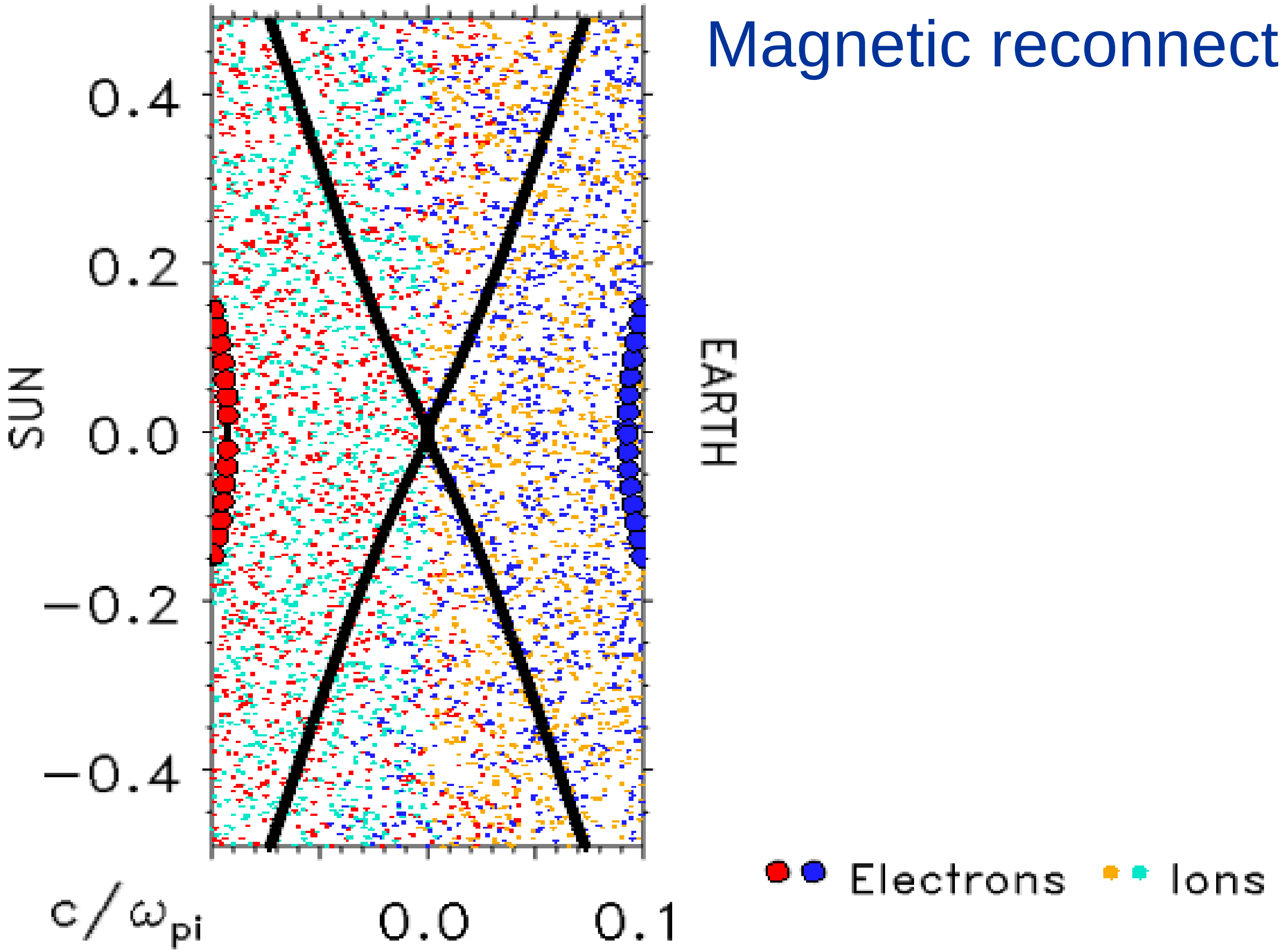


Green

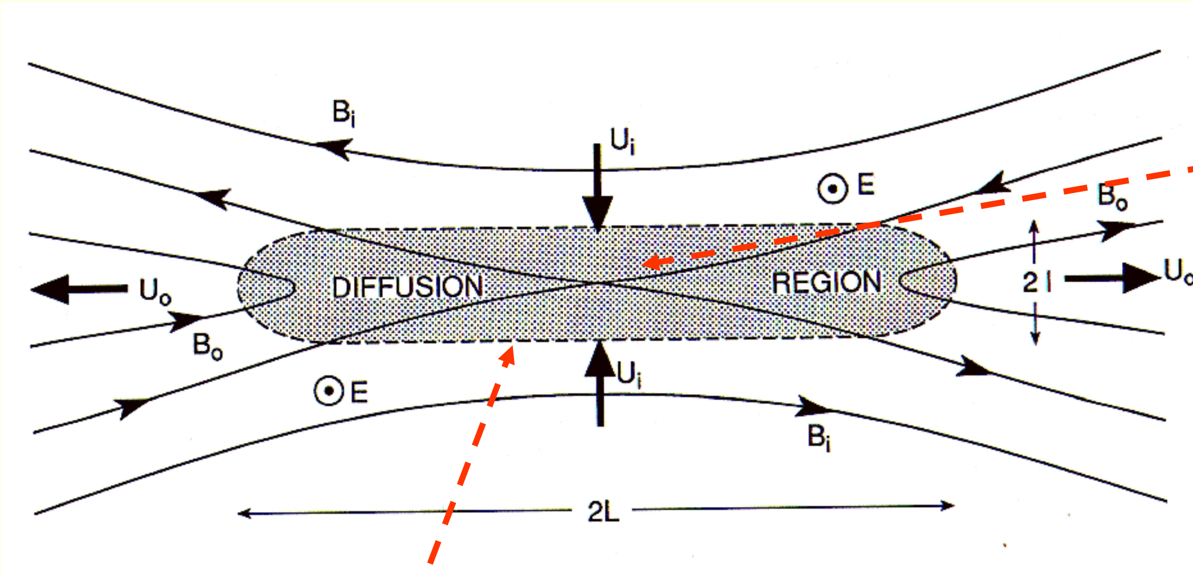
Yellow

Red

Magnetic reconnection



Reconnection



- Field lines are "cut" and can be re-connected to other field lines
- **Magnetic energy is transformed into kinetic energy ($U_o >> U_i$)**

In 'diffusion region':

$R_m = \mu_0 \sigma l v \sim 1$

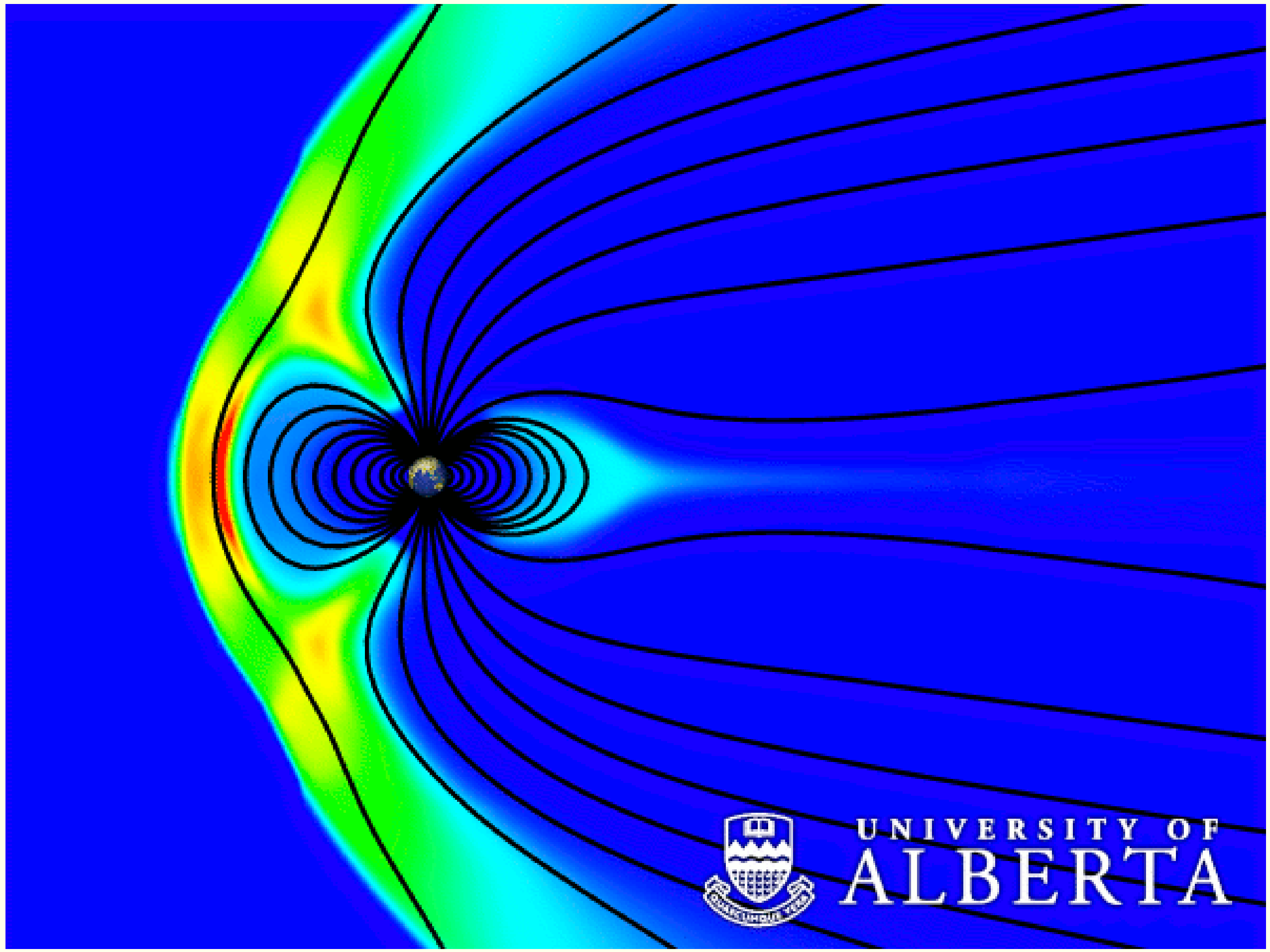
Thus: condition for frozen-in magnetic field breaks down.

A second condition is that there are two regions of magnetic field pointing in *opposite* direction:

- **Plasma from different field lines can mix**

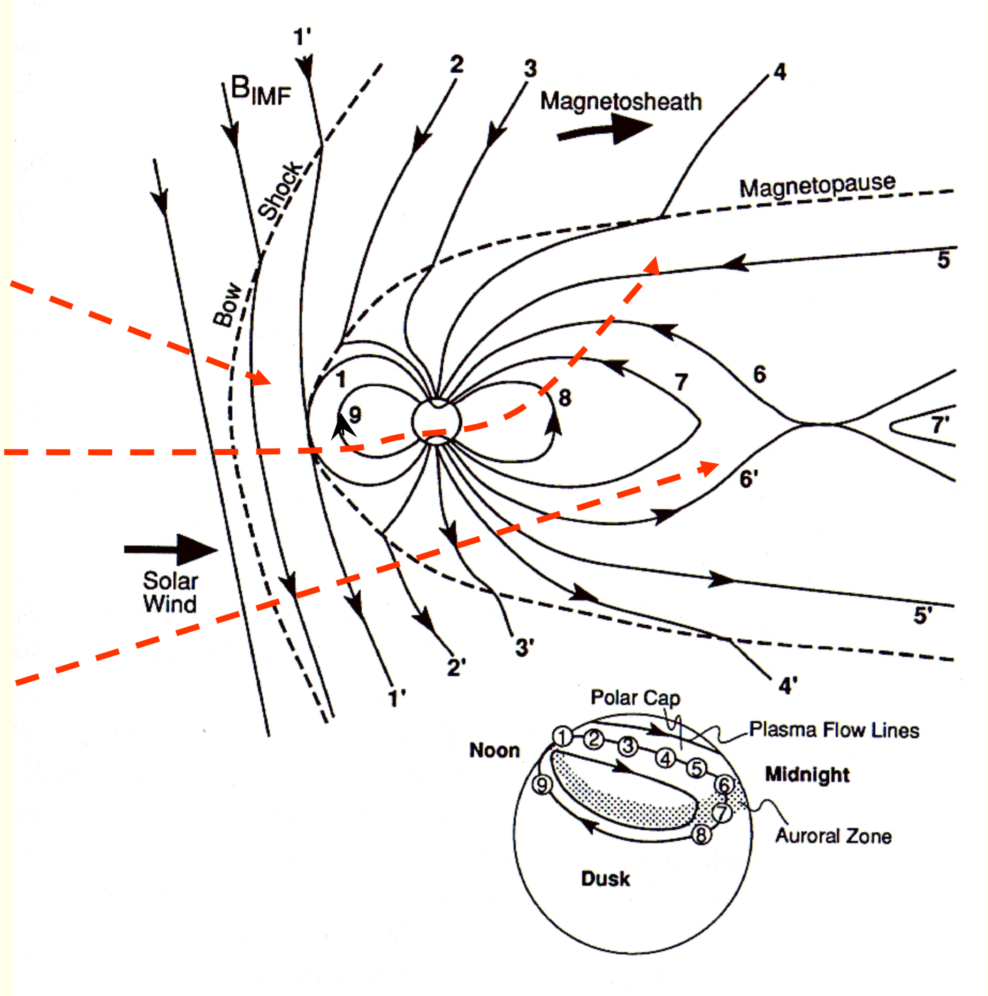
Reconnection and plasma convection

Solar wind

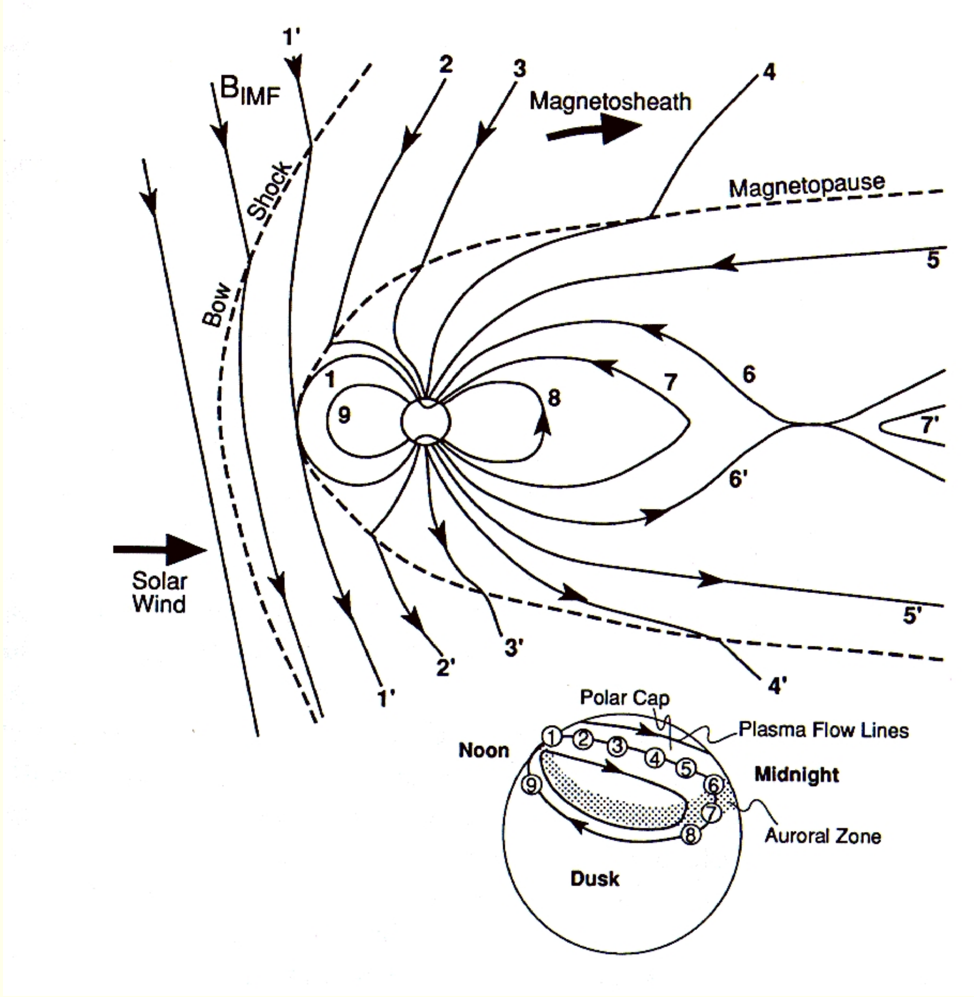


Reconnection och plasma convection

- Reconnection on the dayside "re-connects" the solar wind magnetic field and the geomagnetic field
- In this way the plasma convection in the outer magnetosphere is driven
- In the night side a second reconnection region drives the convection in the inner magnetosphere.
  The reconnection also heats the plasmasheet plasma.

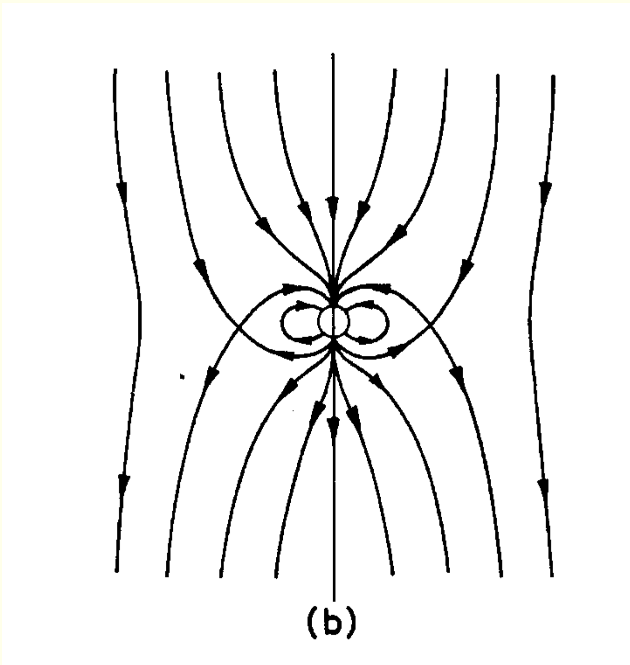


What happens if IMF is northward instead?

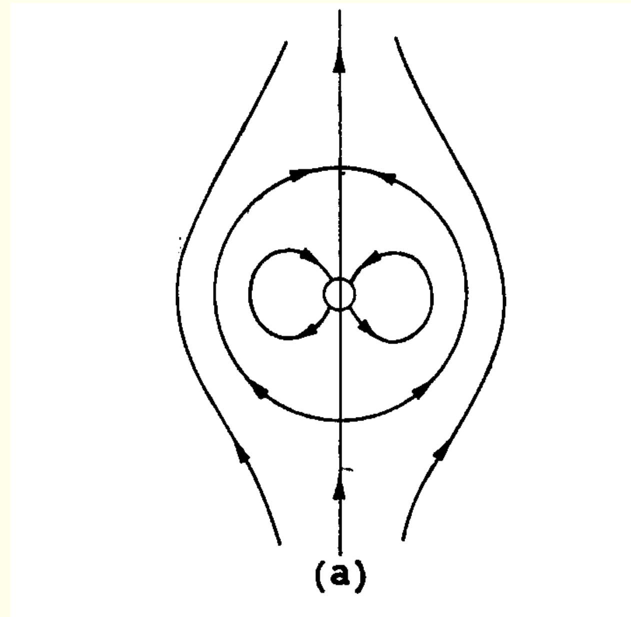


Magnetospheric dynamics

open magnetosphere



closed magnetosphere



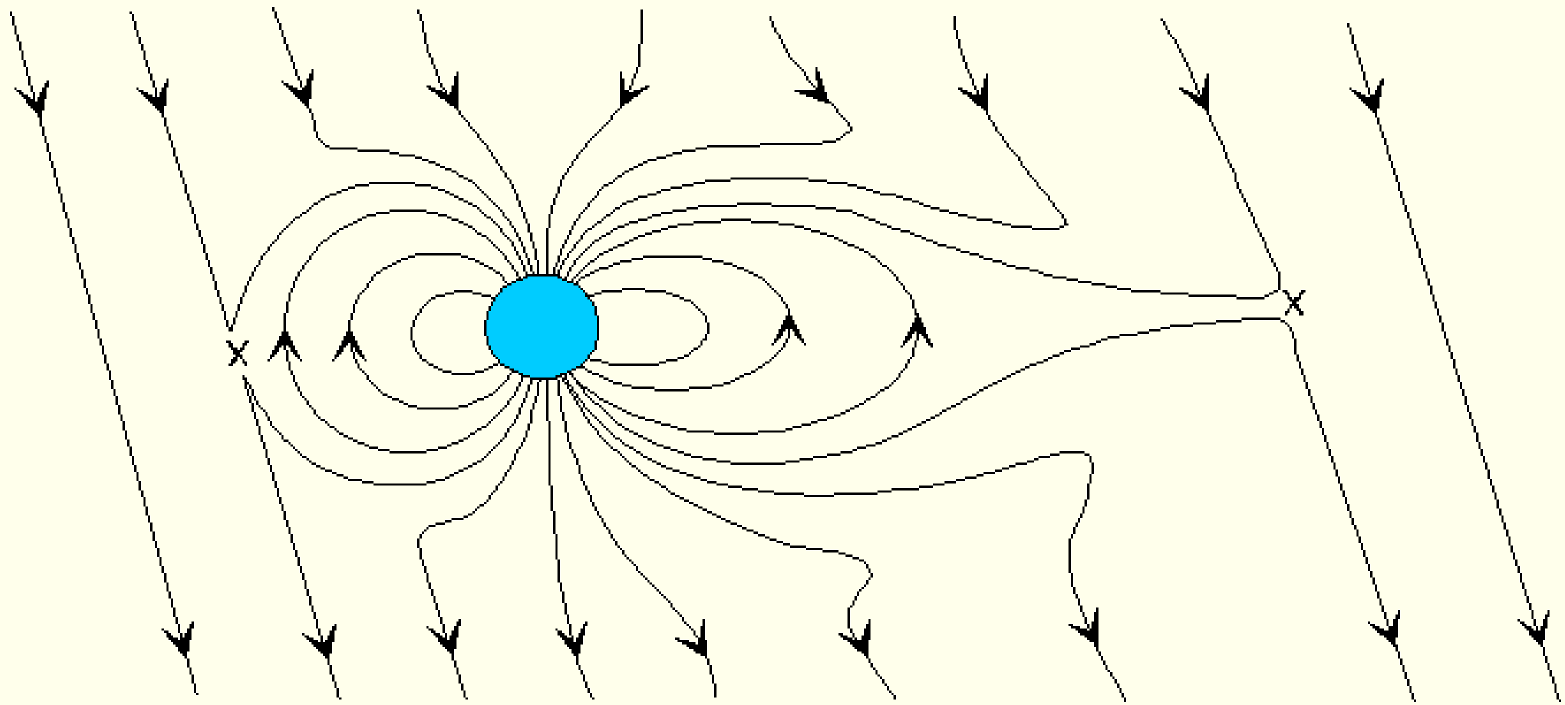
southward

Interplanetary
magnetic field (IMF)

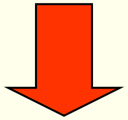
northward

Magnetospheric dynamics

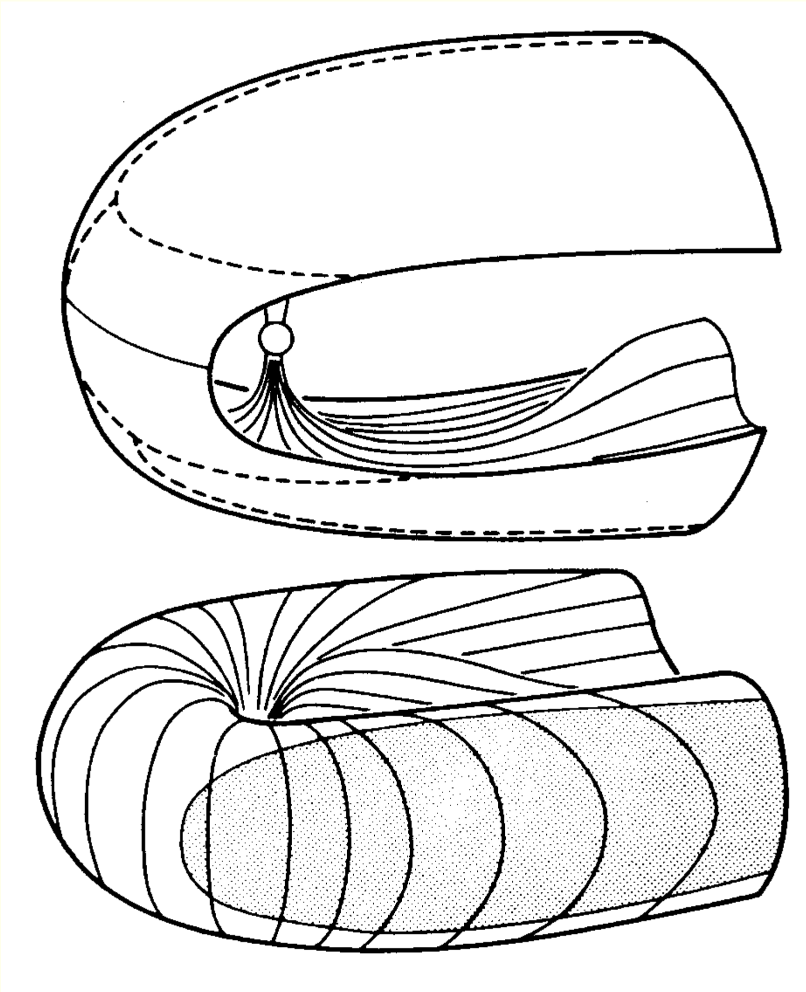
open magnetosphere



Southward IMF



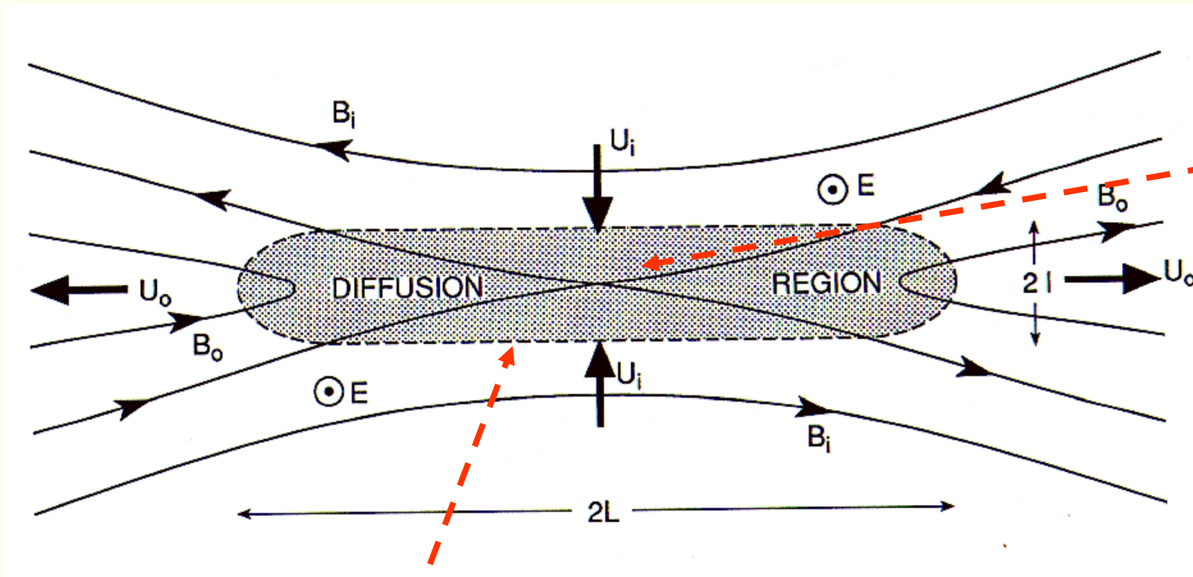
Magnetospheric topology



Open field lines

Closed field lines

Reconnection



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$R_m = \mu_0 \sigma l v \sim 1$

Thus: condition for frozen-in magnetic field breaks down.

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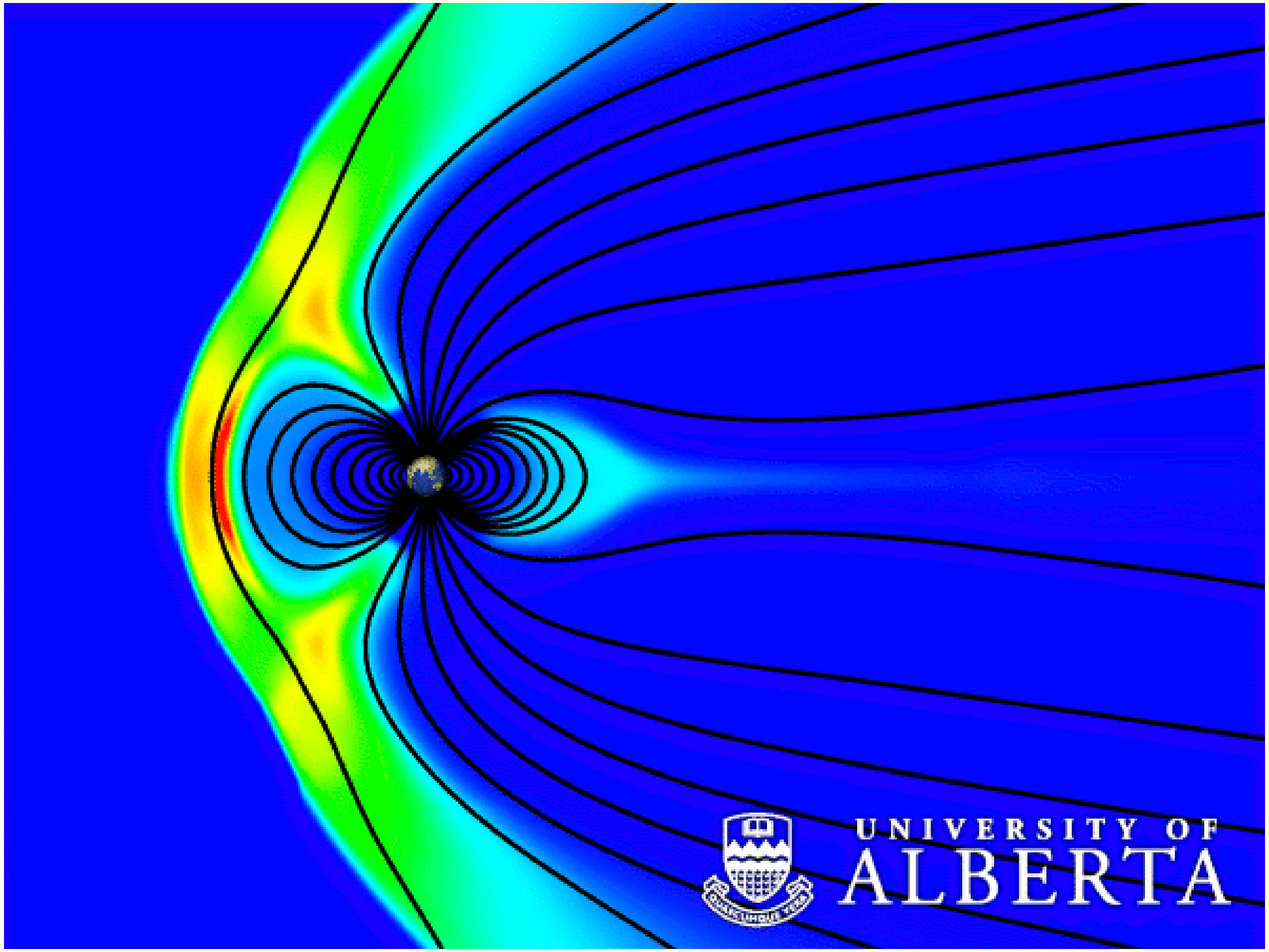
- **Plasma from different field lines can mix**

Reconnection and plasma convection

Solar wind



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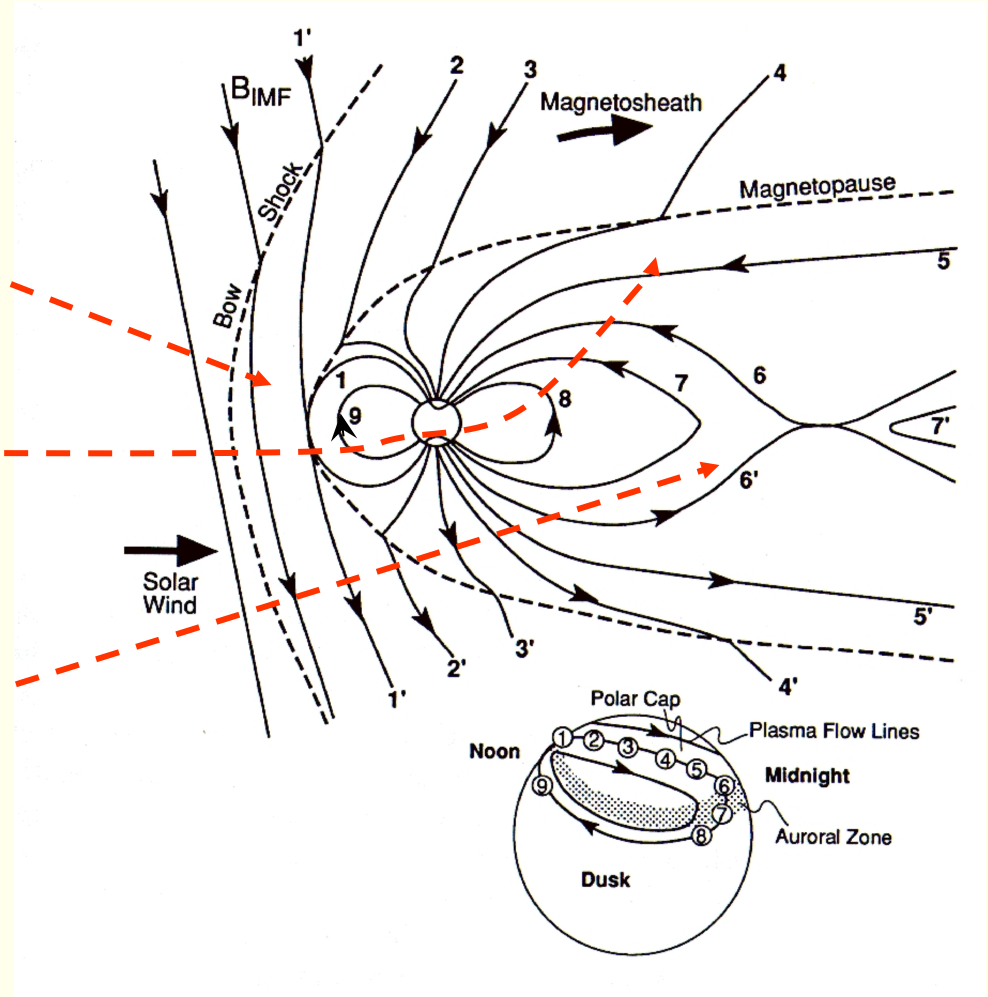
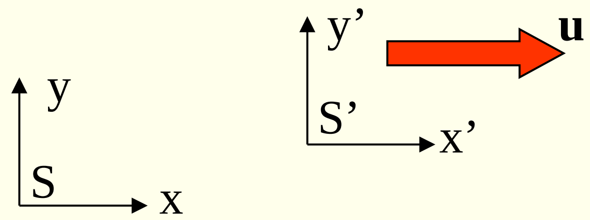


Field transformations (relativistic)



Relativistic transformations
(perpendicular to the velocity $\mathbf{u}$*):*

$$\mathbf{E}' = \frac{\mathbf{E} + \mathbf{u}\times\mathbf{B}}{\sqrt{1-u^2/c^2}}$$

$$\mathbf{B}' = \frac{\mathbf{B} - \left(\mathbf{u}/c^2\right)\times\mathbf{E}}{\sqrt{1-u^2/c^2}}$$

For $u << c$*:*

$$\mathbf{E}' = \mathbf{E} + \mathbf{u}\times\mathbf{B}$$

induced electric field

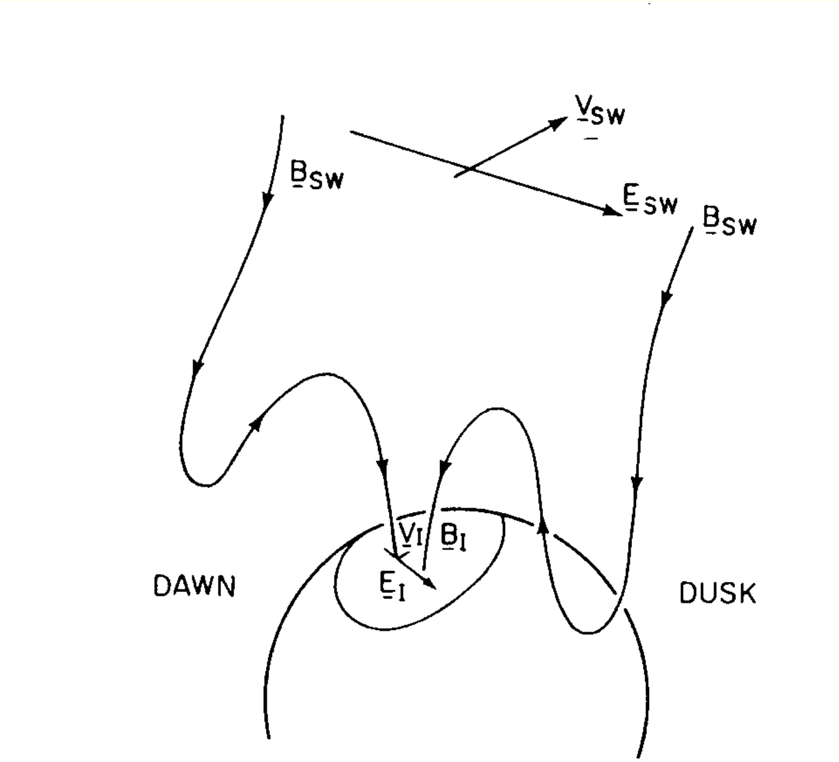
$$\mathbf{E} = \mathbf{E}' - \mathbf{u}\times\mathbf{B}$$

$$\mathbf{B}' = \mathbf{B}$$

Magnetospheric dynamics

open magnetosphere

Viewpoint 1



The solar wind generates an electric field

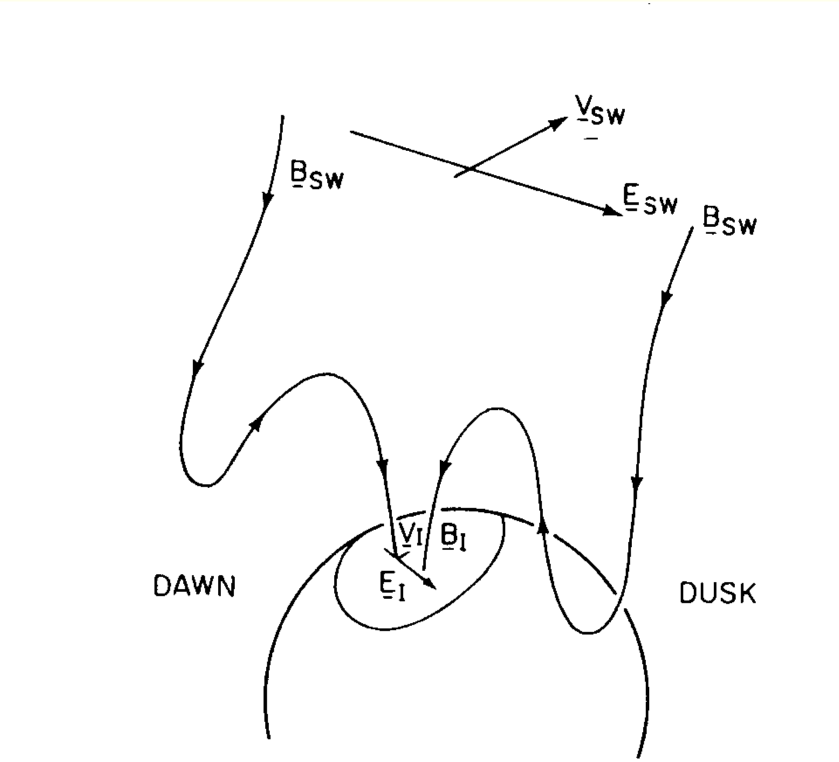
$$\mathbf{E}_{SW} = - \mathbf{v}_{SW} \times \mathbf{B}_{SW}$$

which maps down to the ionosphere, since the field lines are very good conductors

Magnetospheric dynamics

open magnetosphere

Viewpoint 2



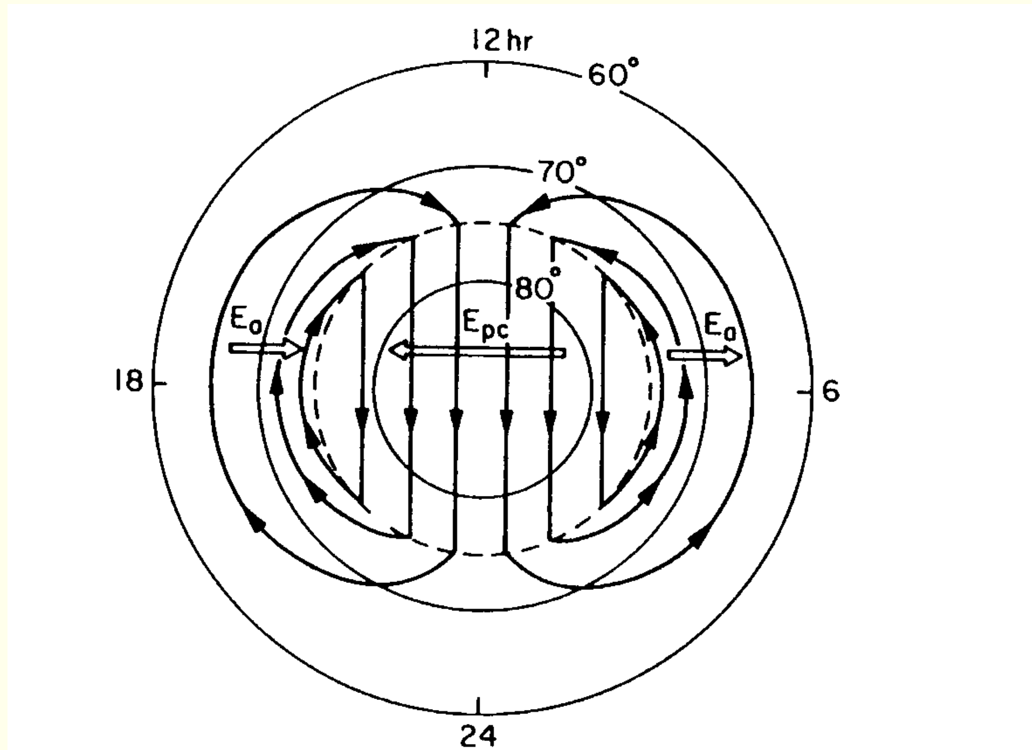
The solar wind magnetic field draws the ionospheric plasma with it, since the field is frozen into the plasma. This motion induces an ionospheric electric field

$$\mathbf{E}_I = -\,\mathbf{v}_I \times \mathbf{B}_I$$

Magnetospheric dynamics

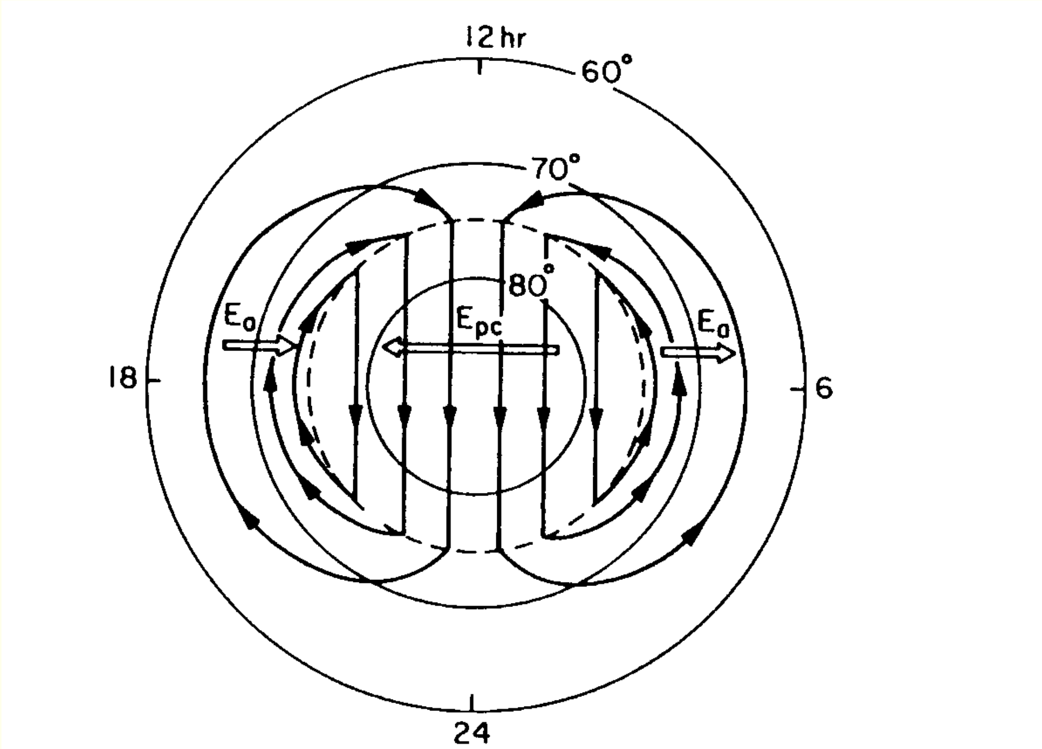
Plasma convection in the ionosphere

The electric field ”propagates” to the ionosphere, since the field lines are good conductors, and thus equipotentials



Do you recognize this pattern?

Plasma convection in the ionosphere



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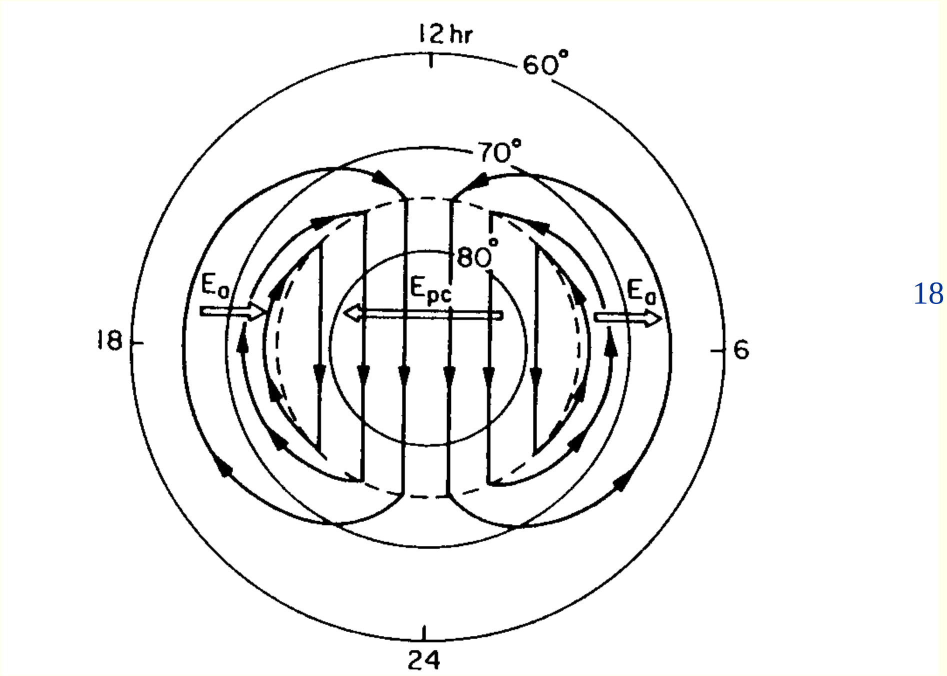
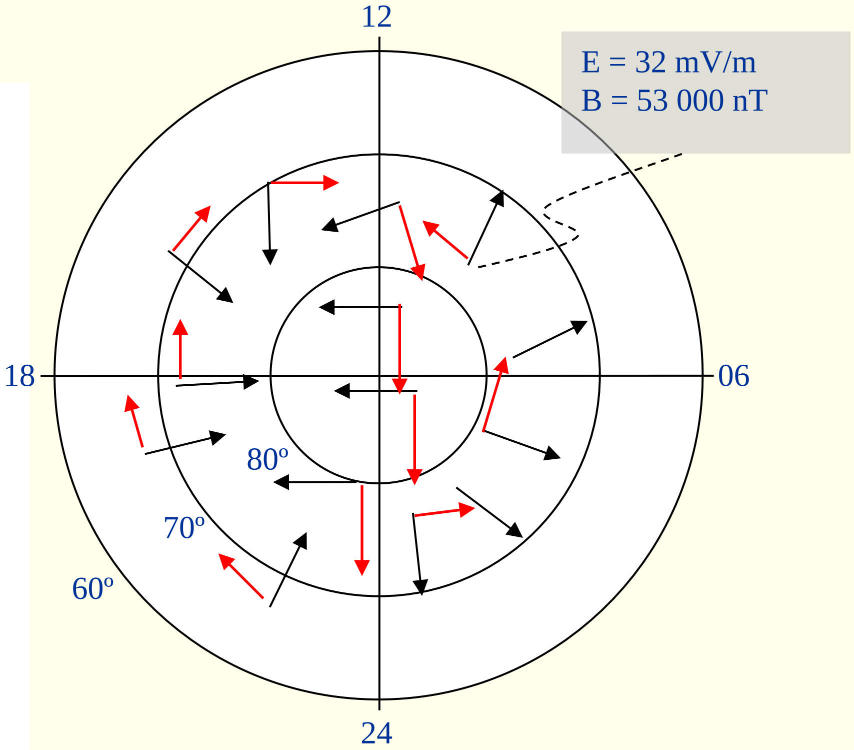
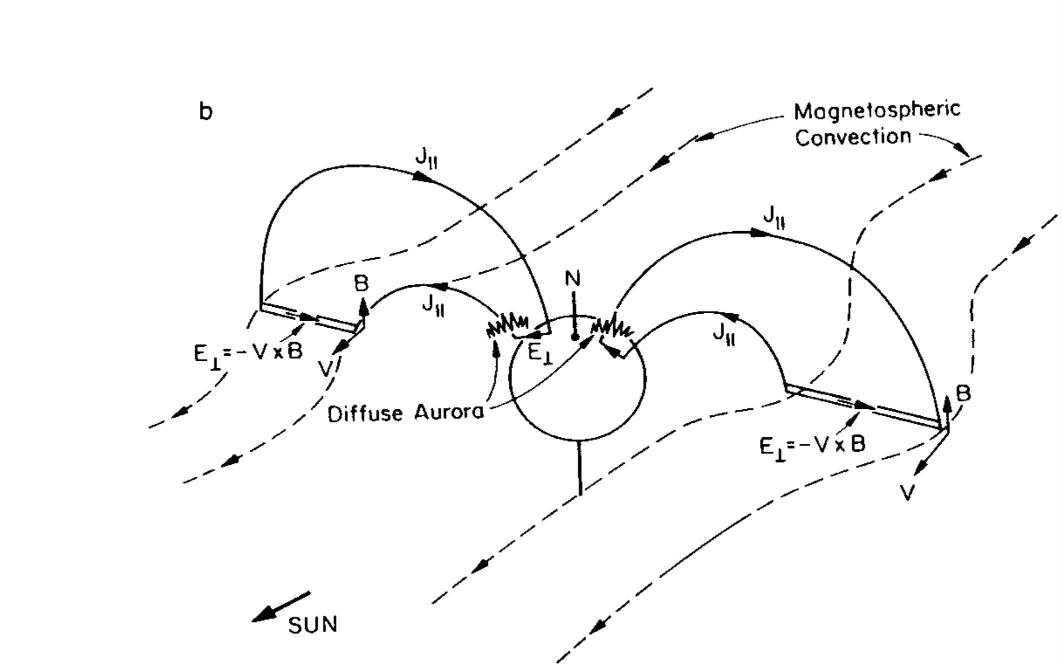
Plasma convection in the ionosphere





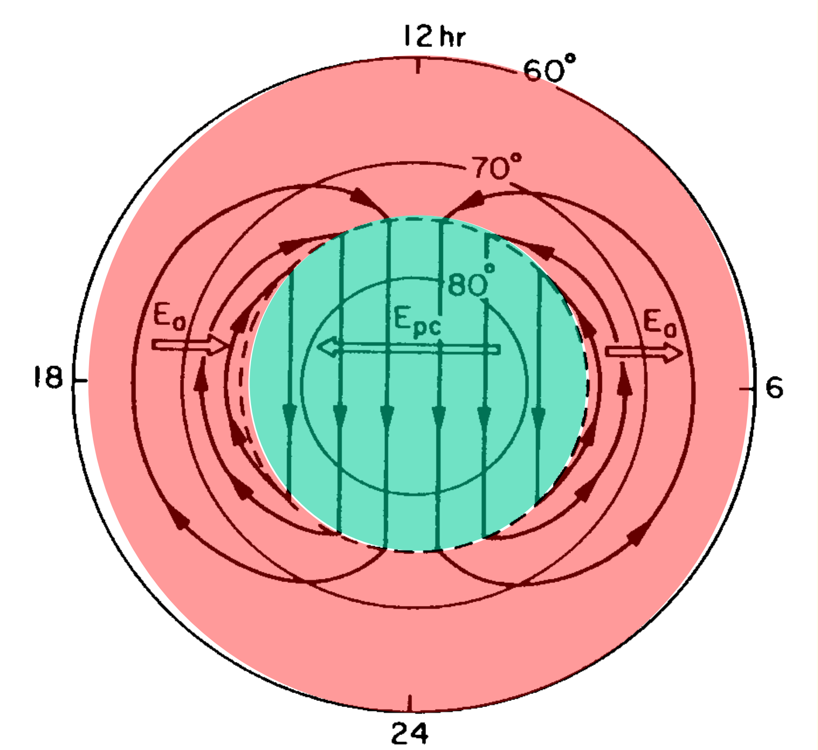
Static, large-scale MI-coupling

Magnetospheric and ionospheric convection

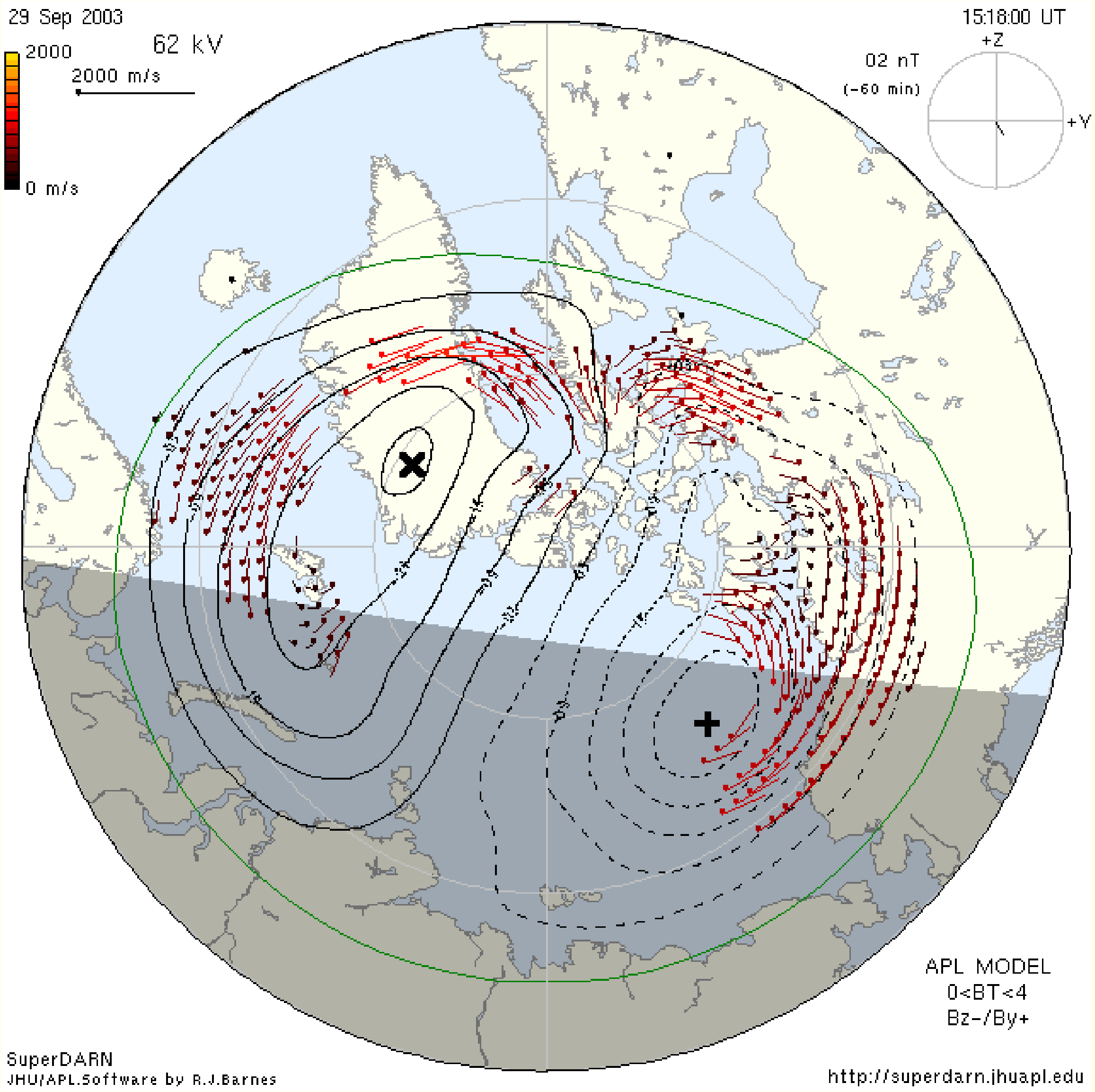


Kelley, 1989

Ionospheric convection

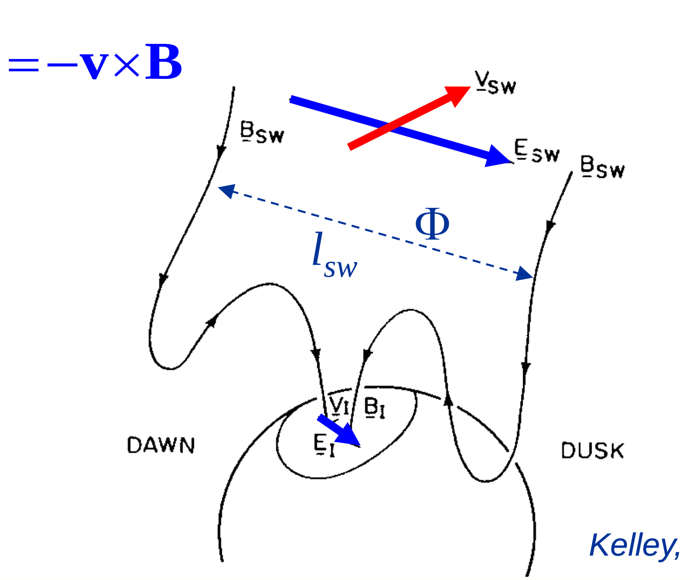


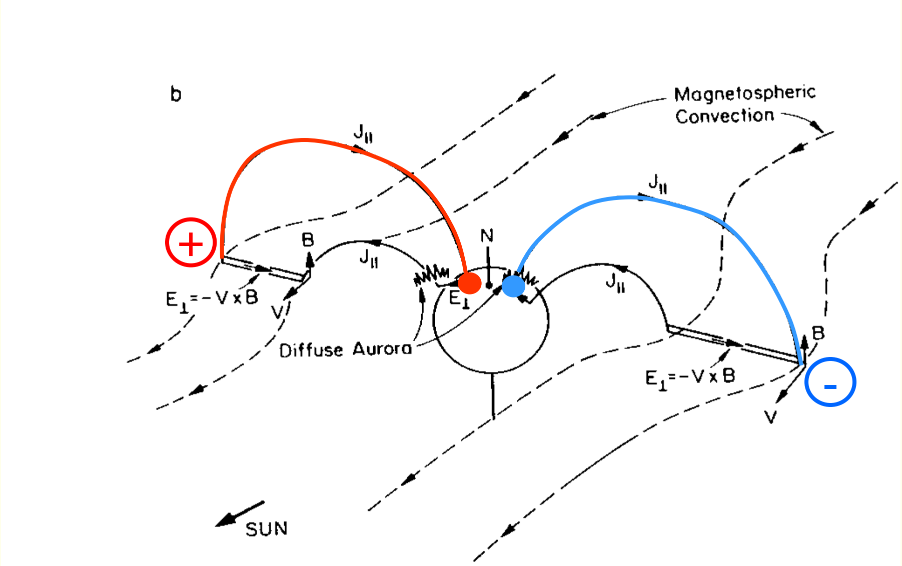
Measurements of plasma convection in the magnetosphere



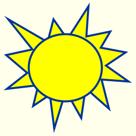
Magnetospheric plasma convection

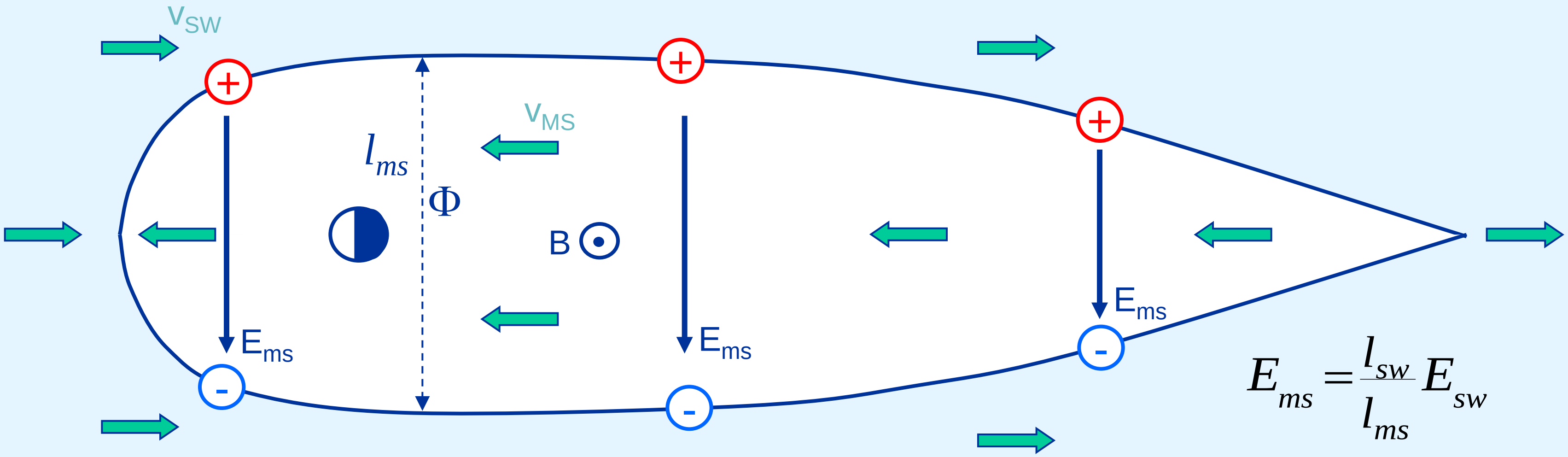
$$\mathbf{E}_{sw} = -\mathbf{v} \times \mathbf{B}$$





Kelley, 1989

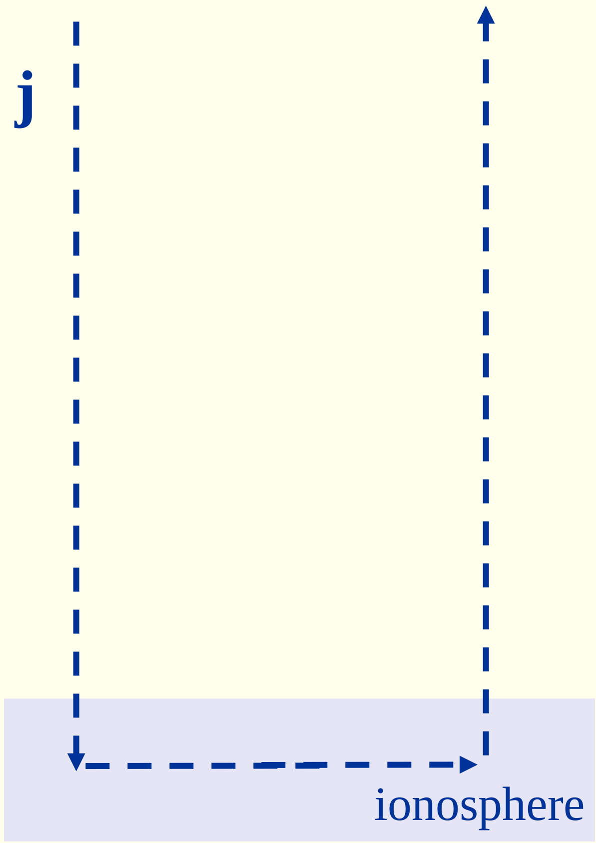




$$E_{ms} = \frac{l_{sw}}{l_{ms}} E_{sw}$$

Geomagnetic activity, definition

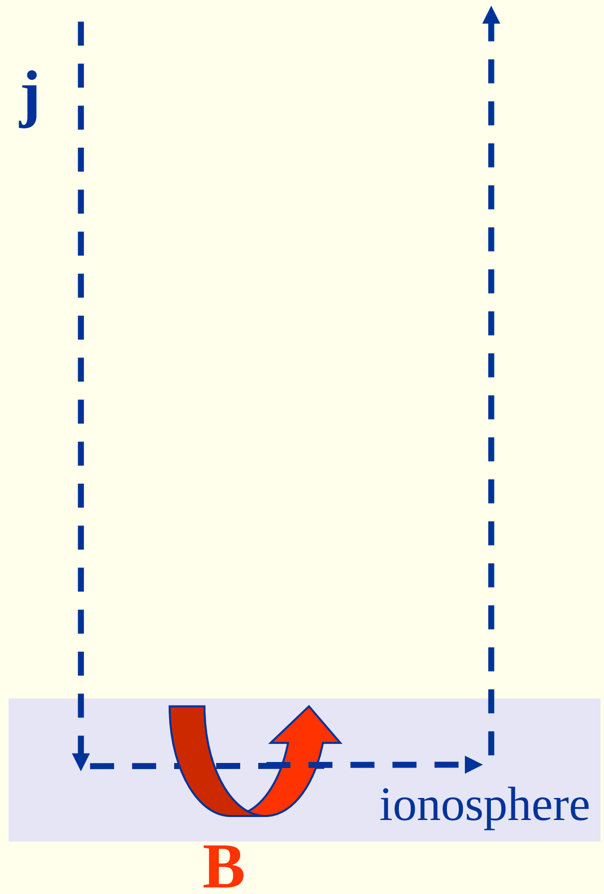
- Geomagnetic activity = temporal variations in the geomagnetic field.
- These variations are caused by temporal variations in the currents in the magnetosphere and ionosphere.



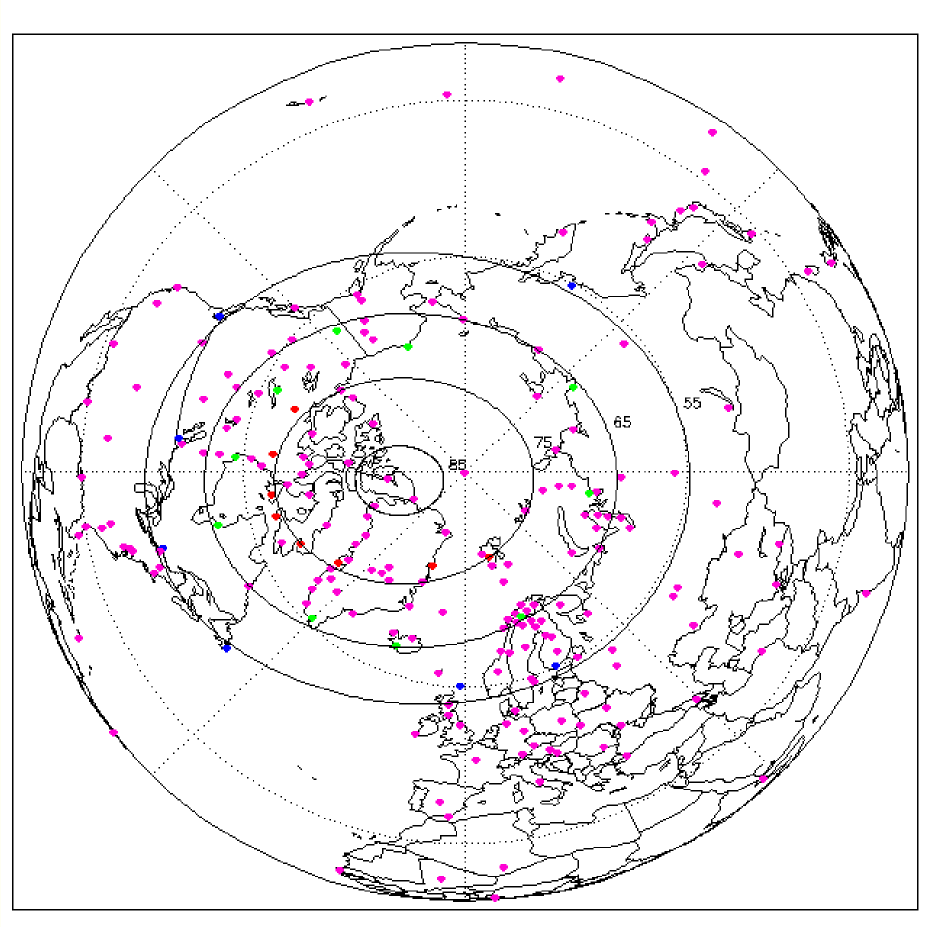
How can you observe these changing currents on Earth?

Geomagnetic activity, definition

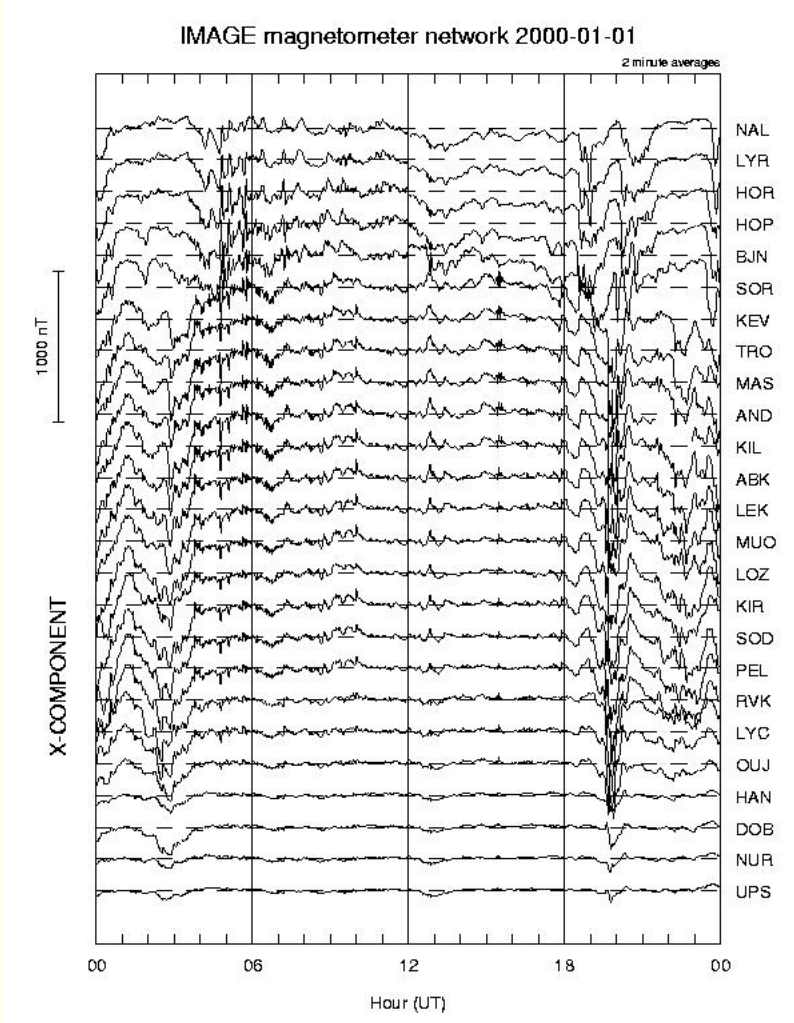
- Geomagnetic activity = temporal variations in the geomagnetic field.
- These variations are caused by temporal variations in the currents in the magnetosphere and ionosphere.
- The variations are observed by geomagnetic observatories



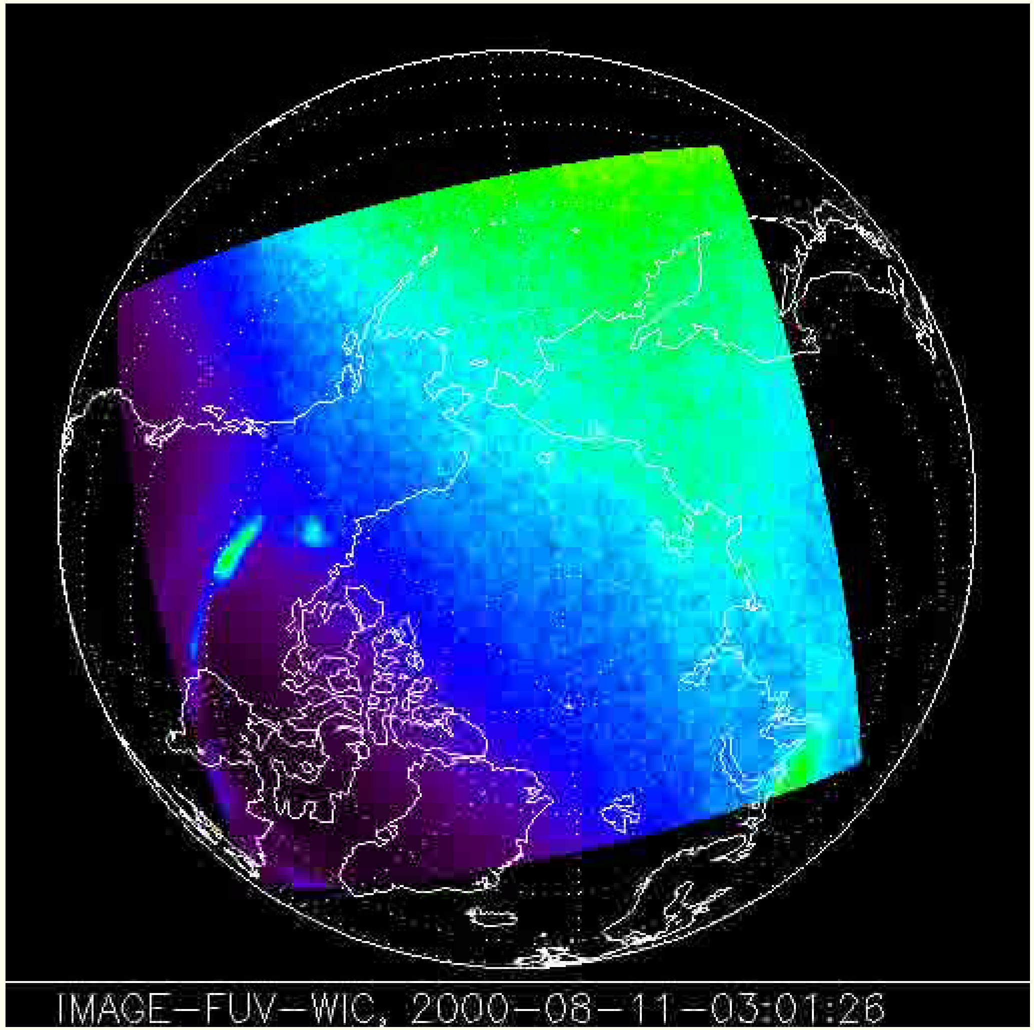
Magnetic observatories



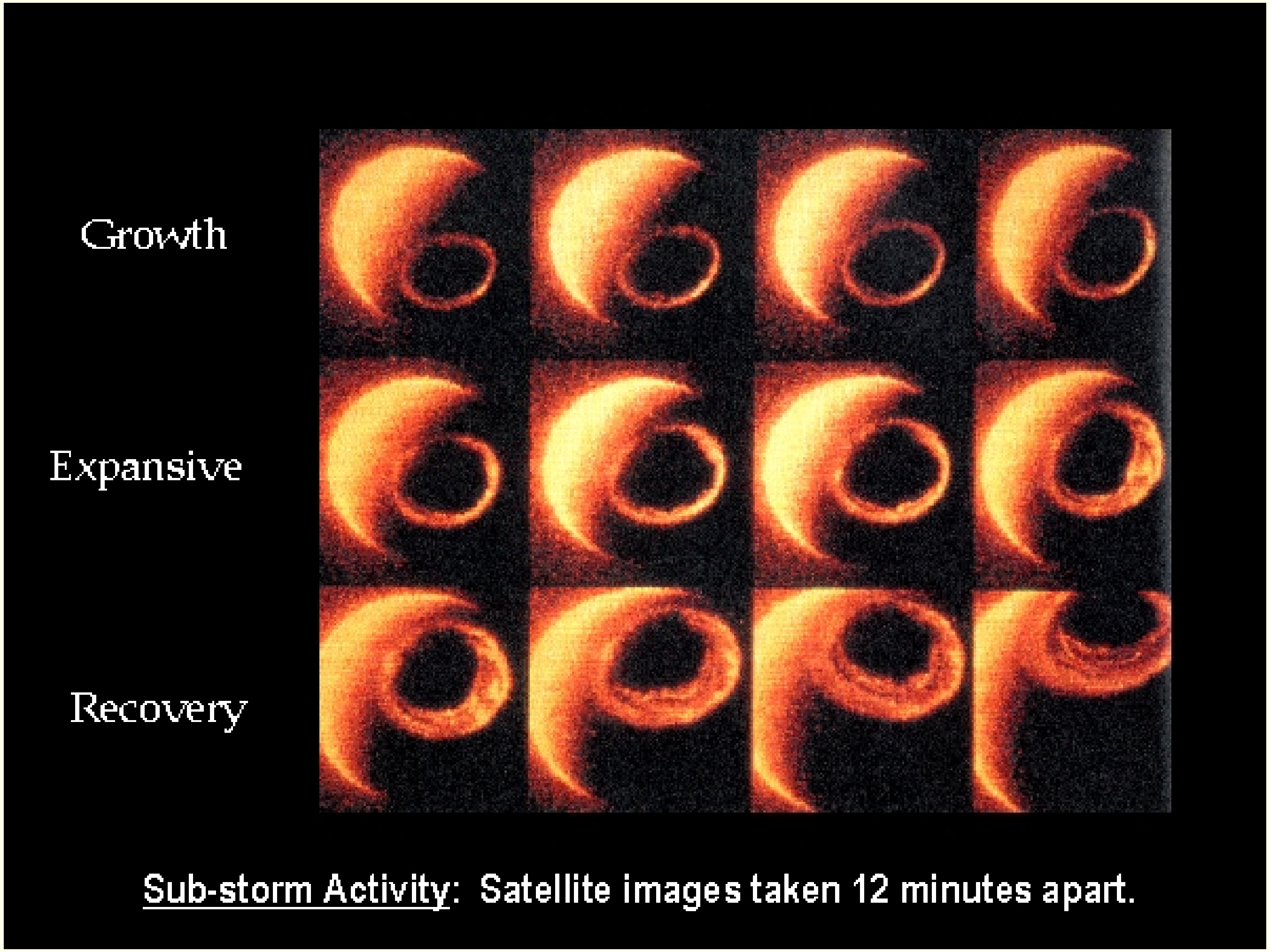
Magnetogram



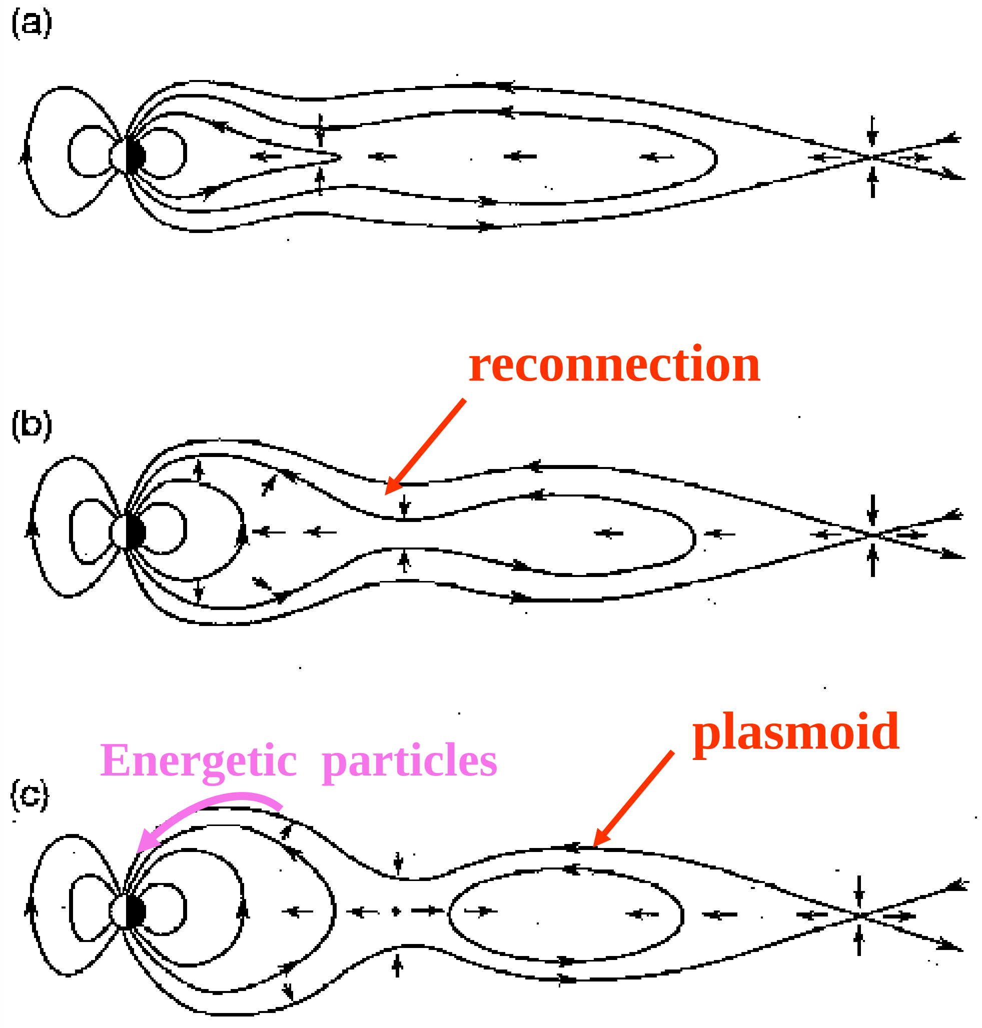
Aurora during substorm



Aurora during substorm

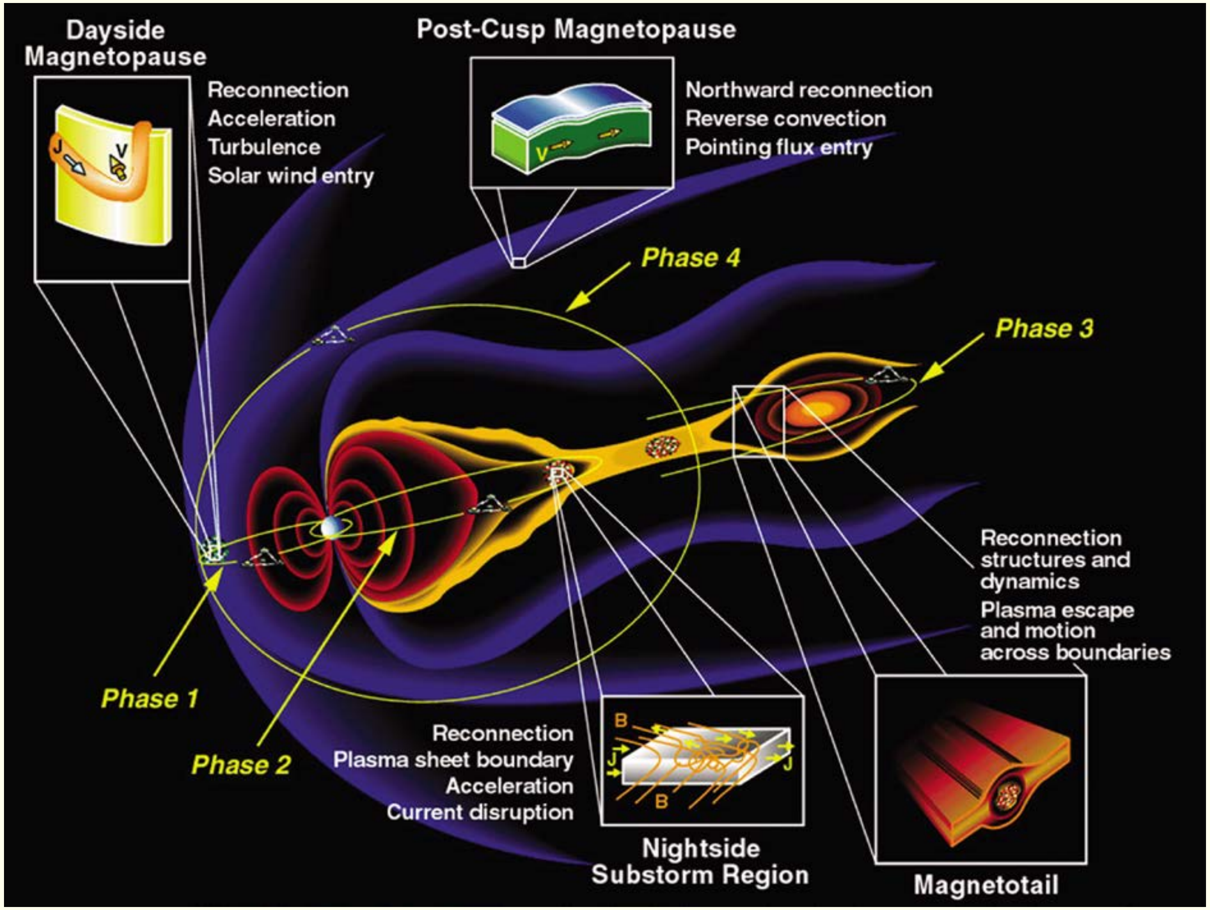


Substorms - magnetosphere



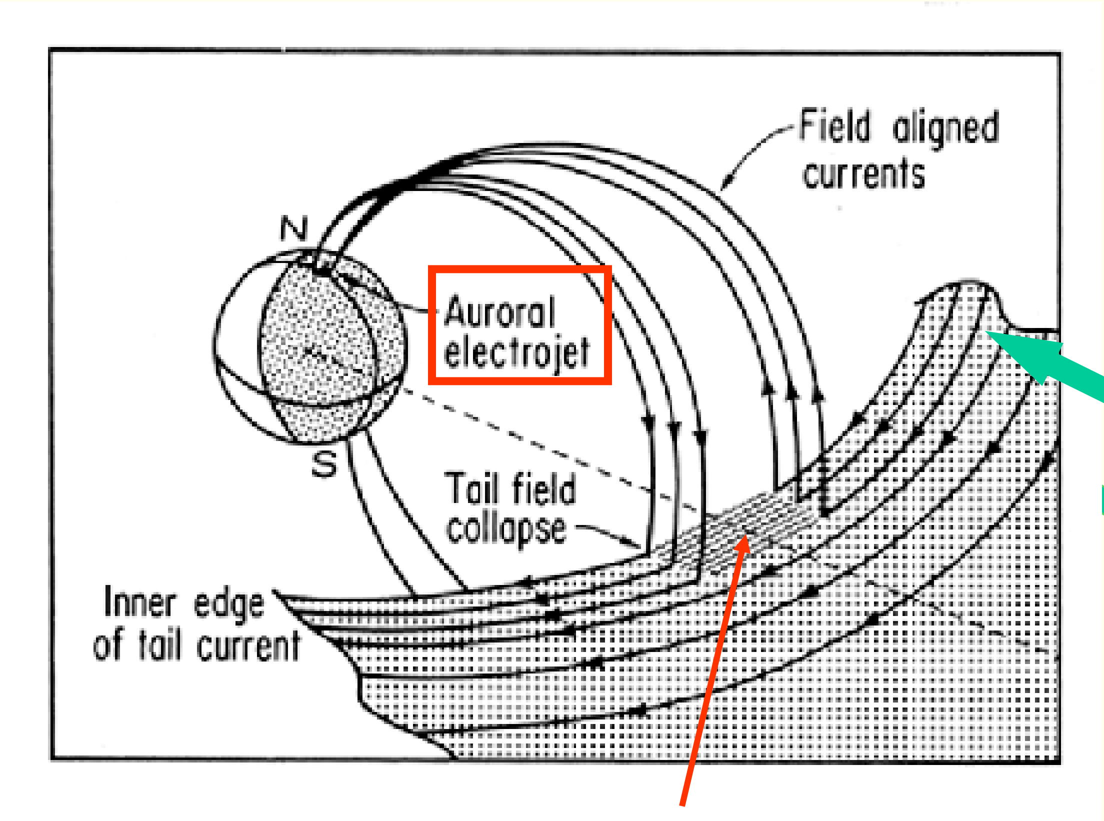
- **GROWTH PHASE**: When IMF southward, energy is pumped into magnetostail and is stored as megnetic energy
- **ONSET:** After a certain time (~1 h) the magnetostail goes unstable and "snaps" due to fast reconnection.
- **EXPANSION/MAIN PHASE**: Close to Earth the magnetosphere returns to dipole-like cinfiguration. Plasma is energized and injected into the inner parts of the magnetosphere.
- **RECOVERY PHASE**: In the outer parts of the magnetotail a *plasmoid* is ejected. The magnetosphere returns to its ground state.

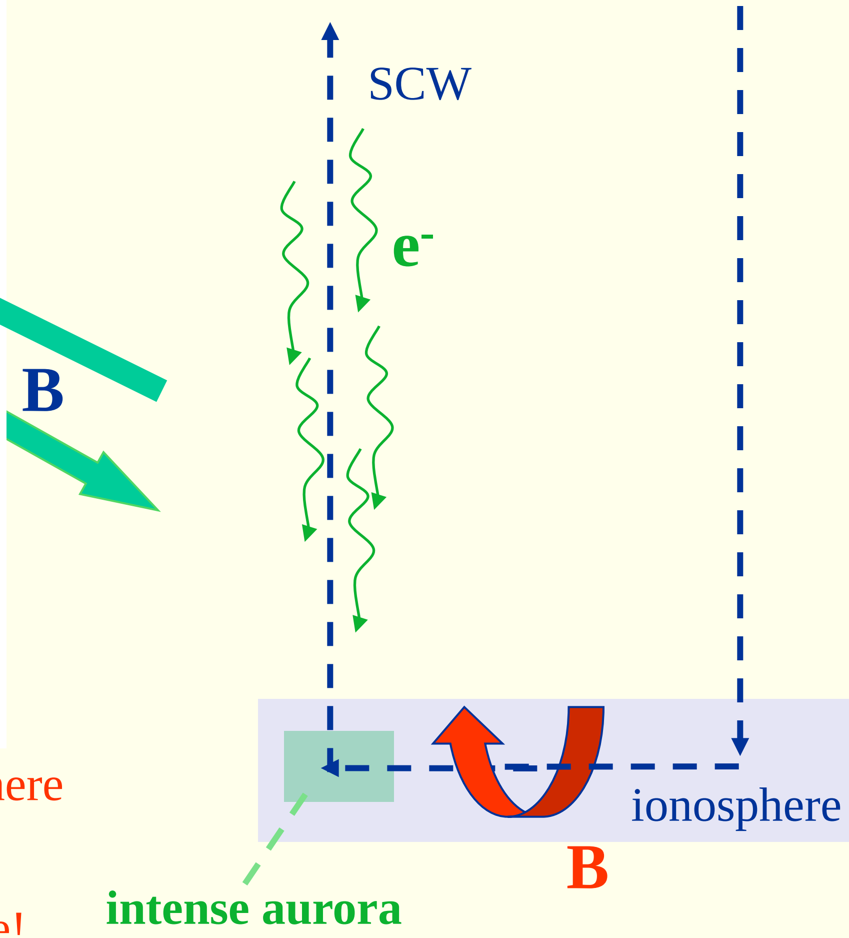
Substorms - magnetosphere



- **GROWTH PHASE**: When IMF southward, energy is pumped into magnetostail and is stored as megnetic energy
- **ONSET:** After a certain time (~1 h) the magnetostail goes unstable and "snaps" due to fast reconnection.
- **EXPANSION/MAIN PHASE**: Close to Earth the magnetosphere returns to dipole-like cinfiguration. Plasma is energized and injected into the inner parts of the magnetosphere.
- **RECOVERY PHASE**: In the outer parts of the magnetotail a *plasmoid* is ejected. The magnetosphere returns to its ground state.

Substorm Current Wedge (SCW)



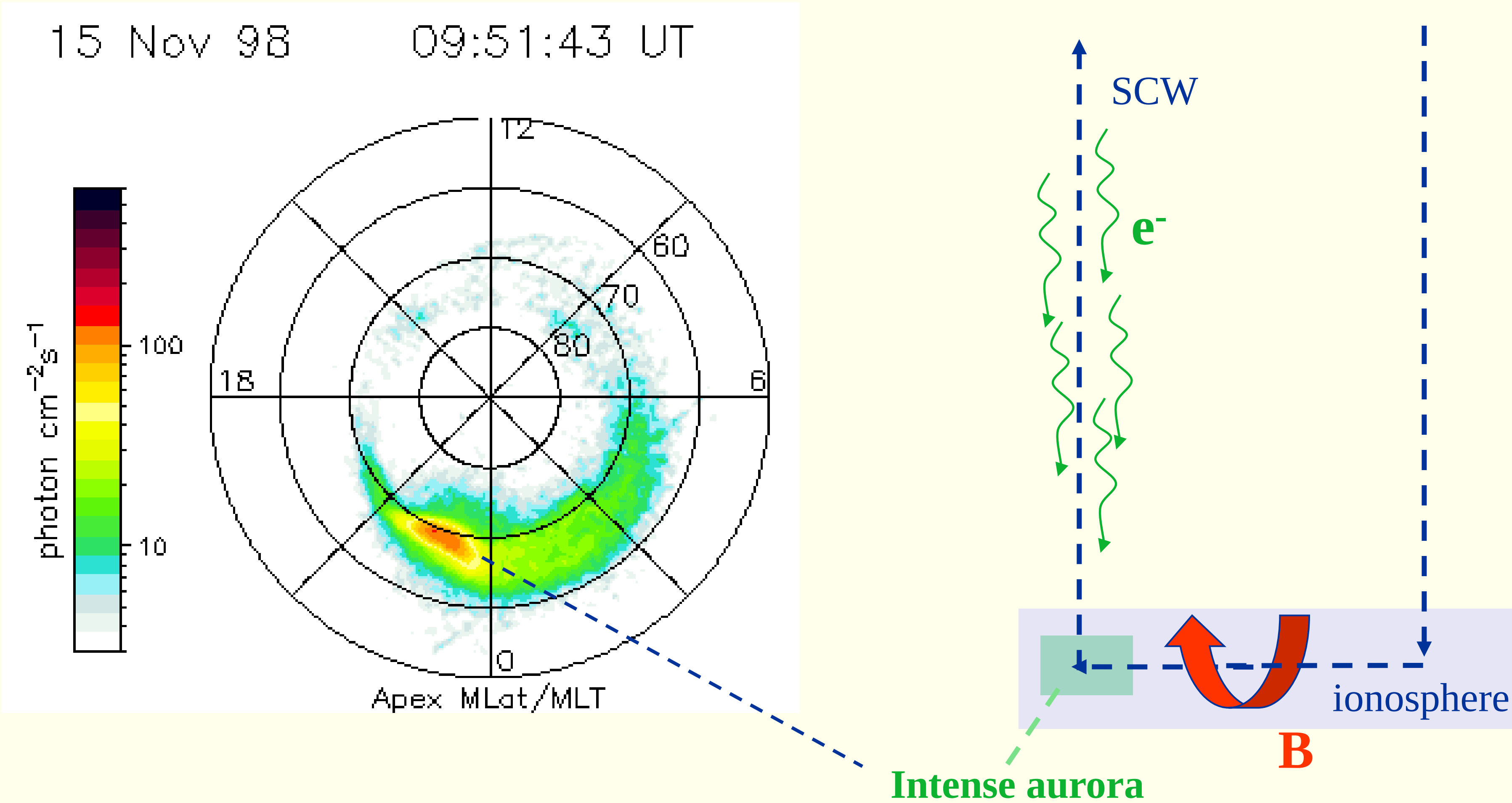


Due to reconnection processes the resistivity increases here

⇒

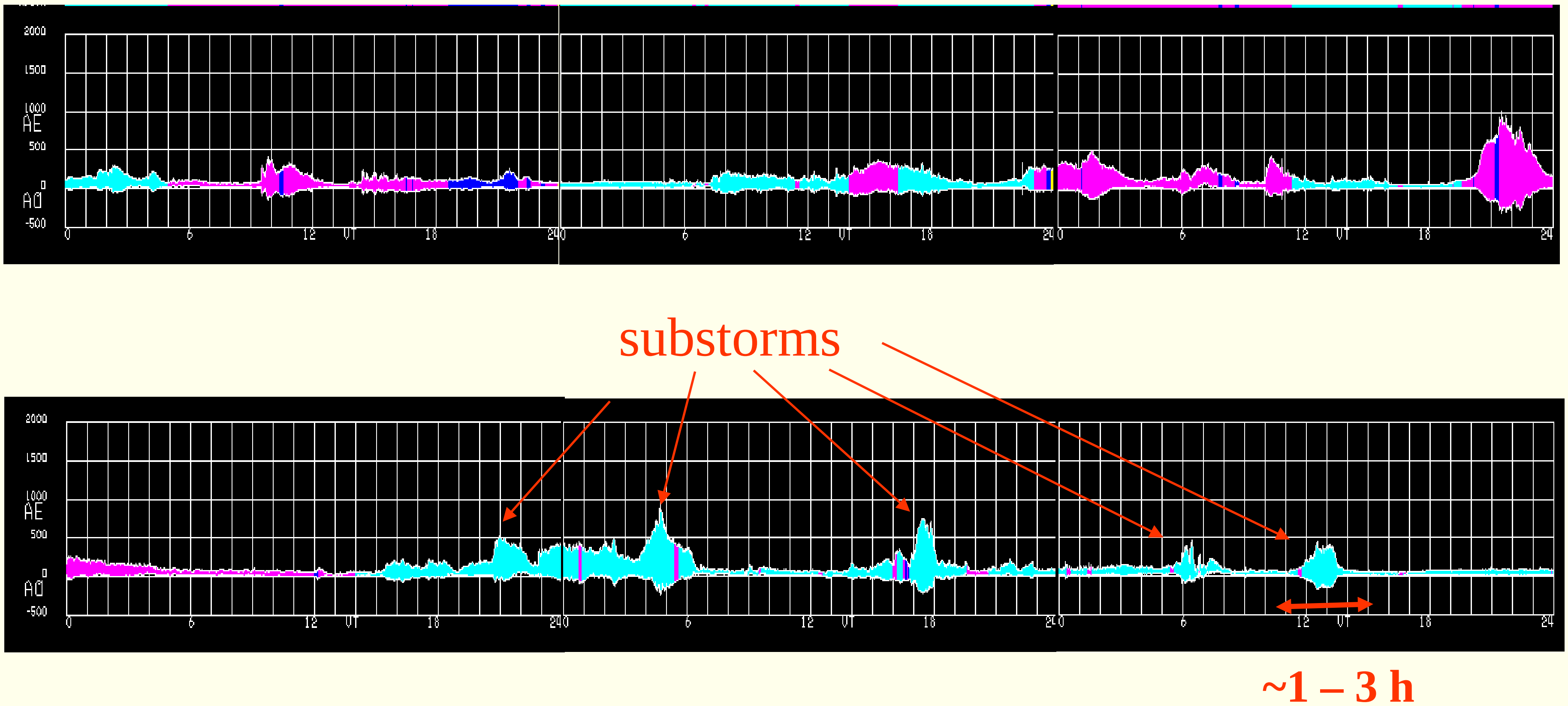
Current takes another direction – through the ionosphere!

Substorm Current Wedge (SCW)



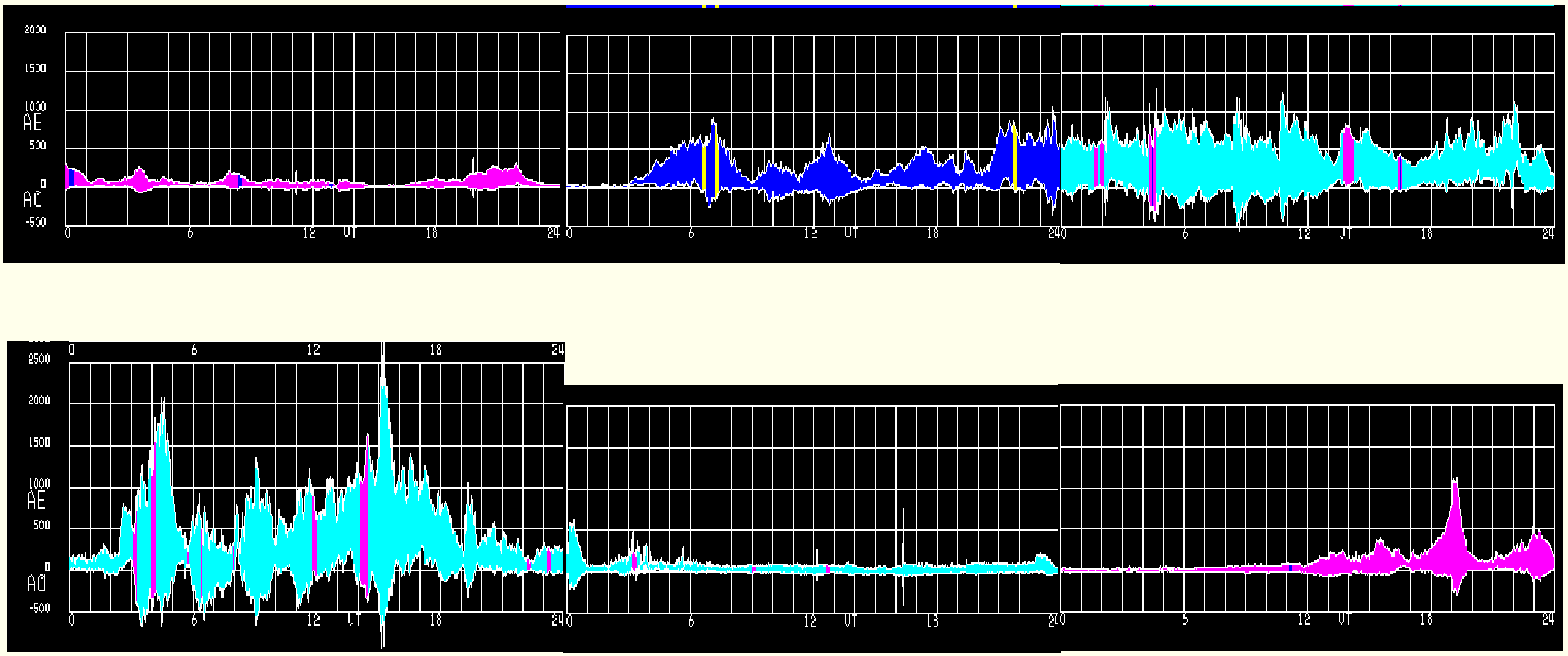
Auroral Electrojet (AE) index

The AE index Measures the strength of the substorm current wedge (SCW), by using the information from several magnetic observatories.



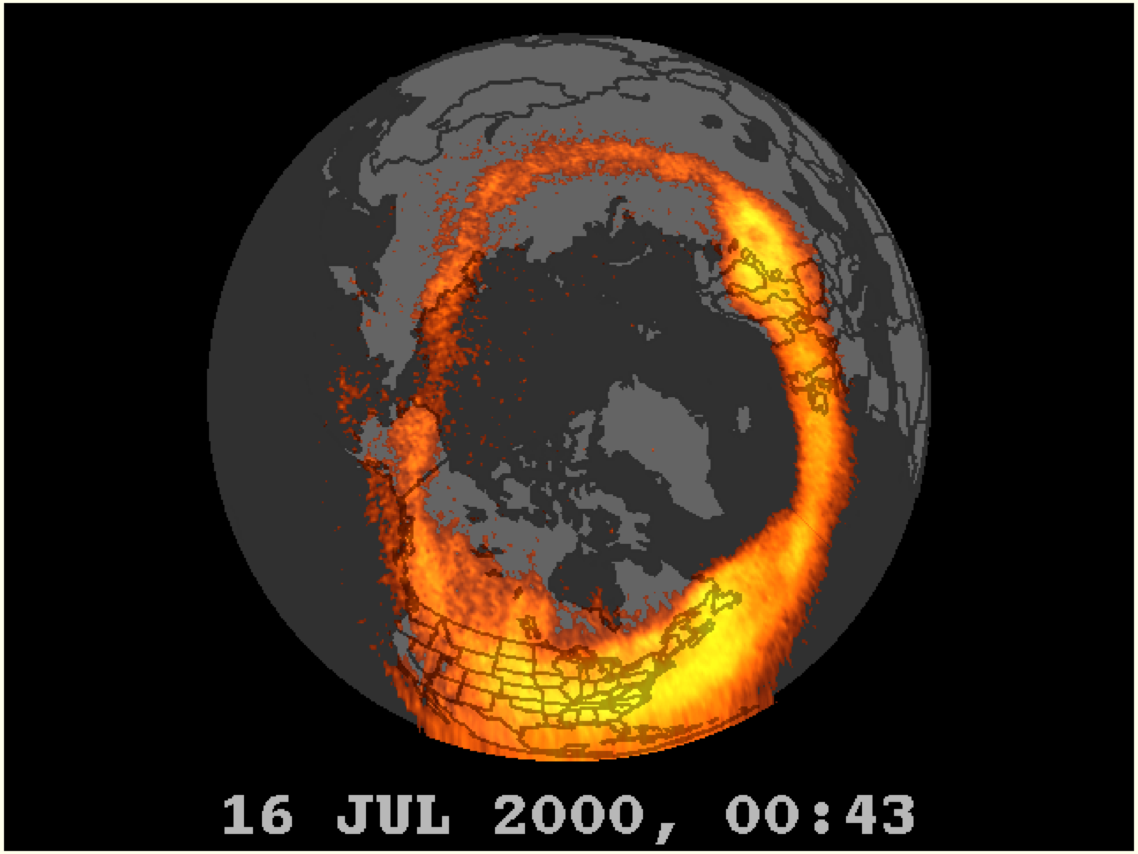
Geomagnetic storms

Geomagnetic storms are extended periods with southward interplanetary magnetic field (IMF) and a large energy input into the magnetosphere.

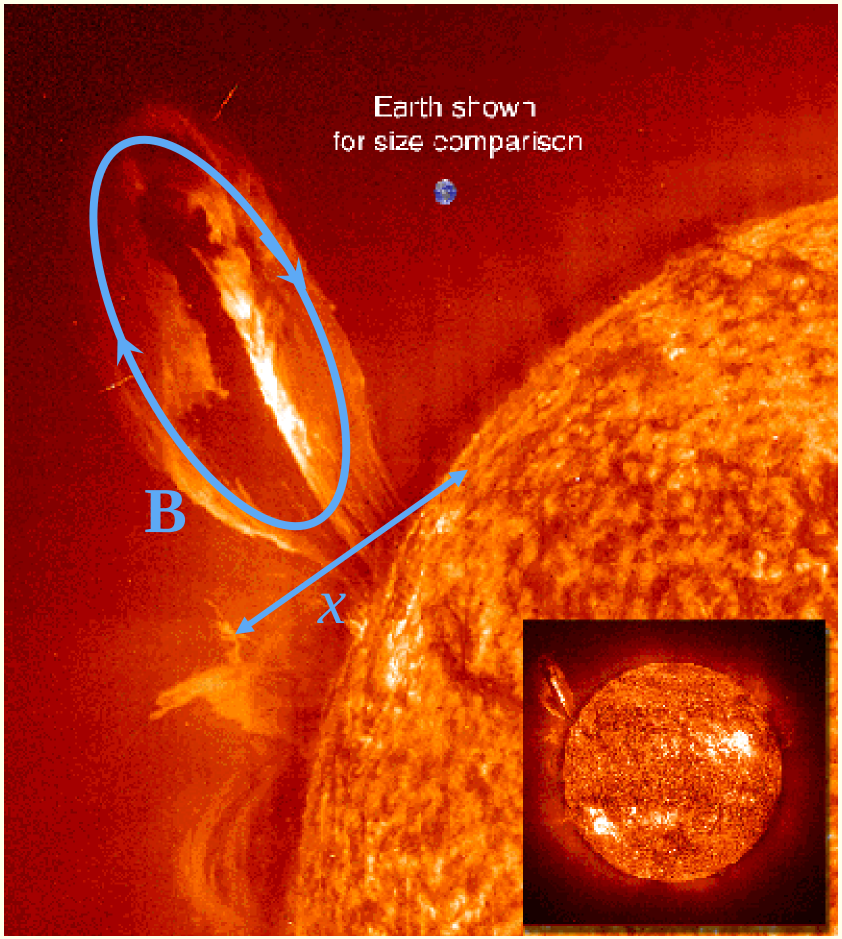


Geomagnetic storms

Auroral oval very extended

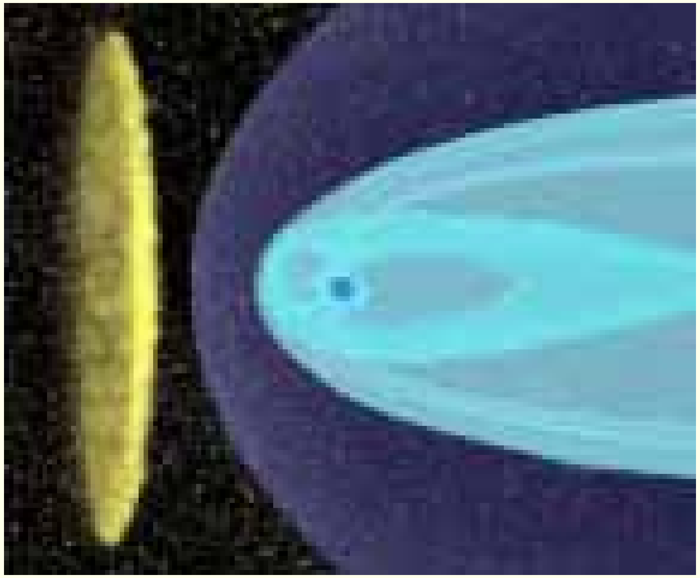


Geomagnetic storms and coronal mass ejections



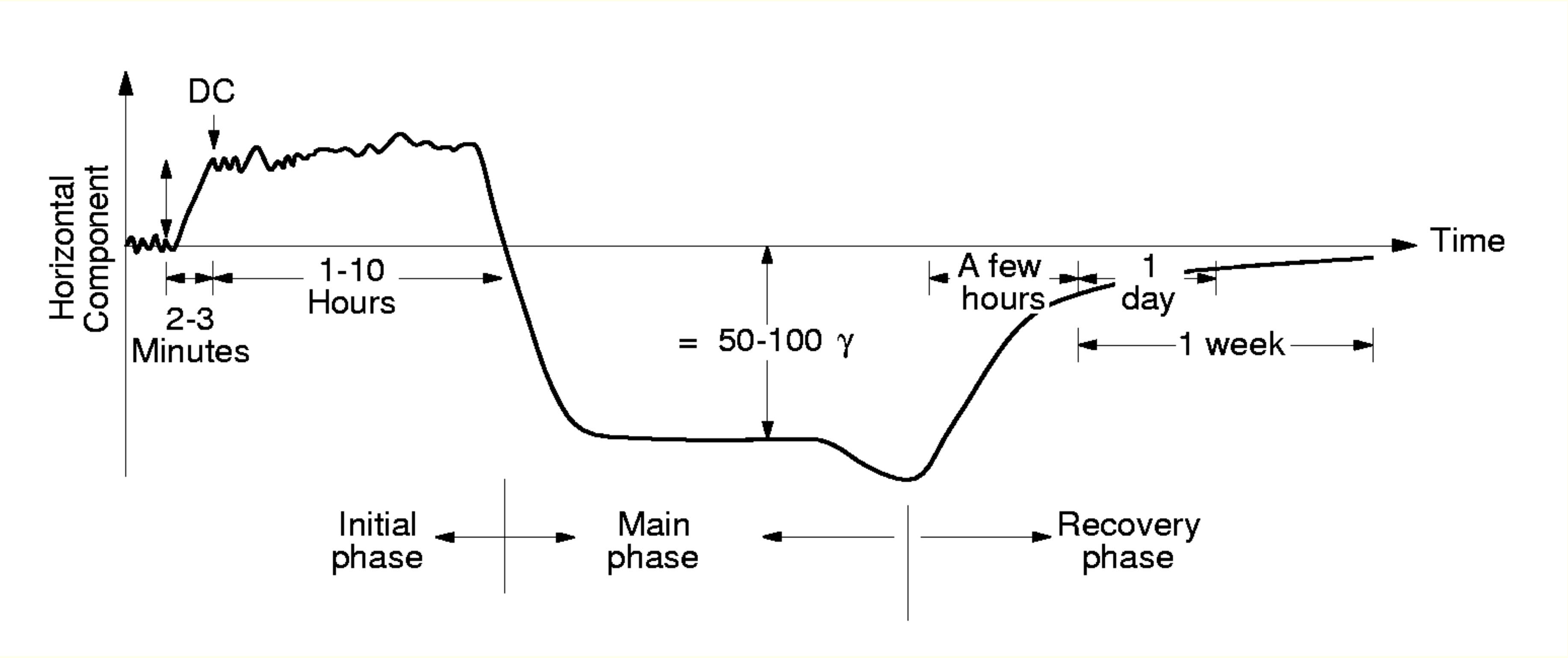
- Large geomagnetic storms are often associated with coronal mass ejections (CMEs)
- Because of their magnetic structure, they will give long periods with a constant IMF
- A typical time for a CME to pass Earth becomes $T = x/v \sim 10\ R_E/1000\ \mathrm{kms^{-1}} \sim 60$ h

What happens with the geomagnetic field when the CME hits the magnetosphere?

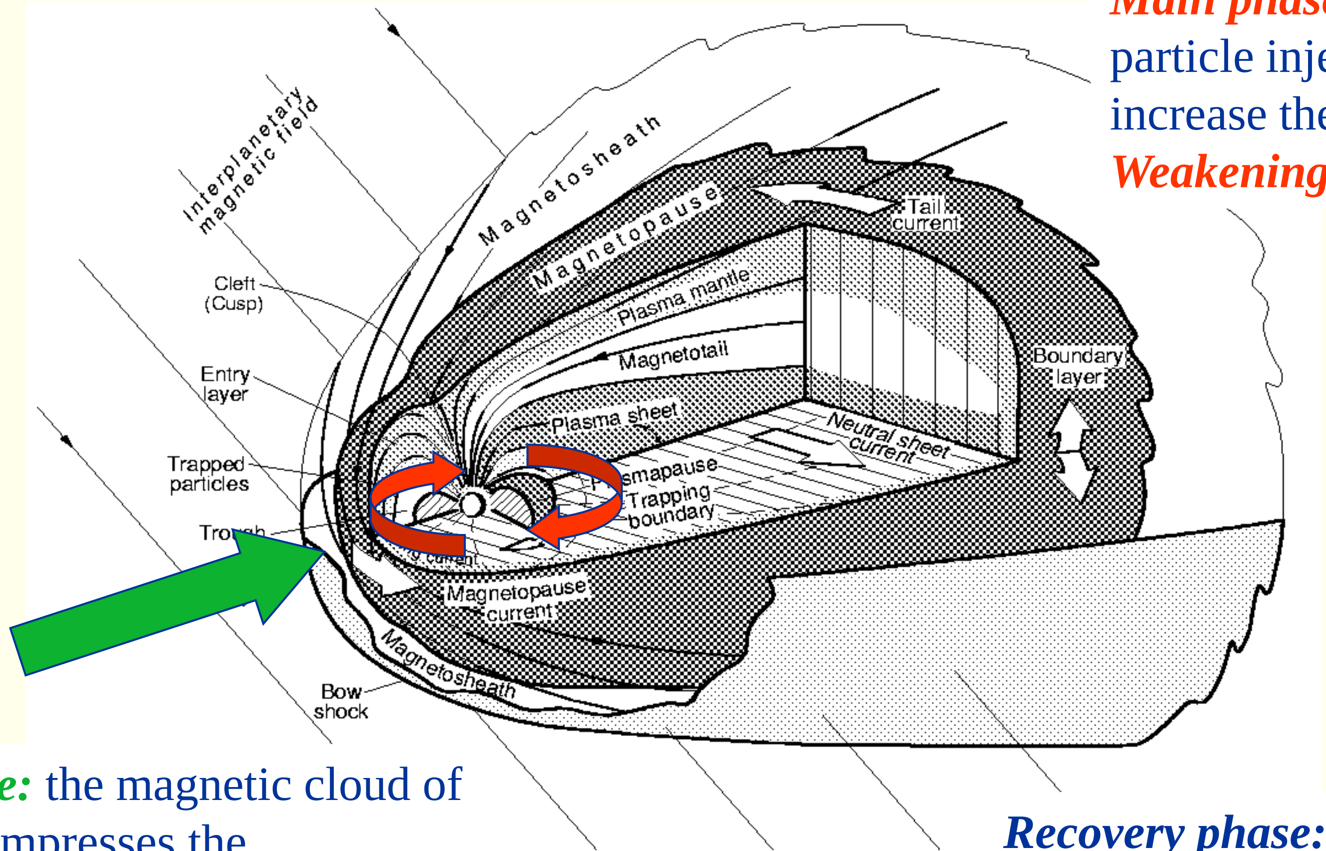


Geomagnetic storms - phases

Magnetogram



Geomagnetic storms - phases

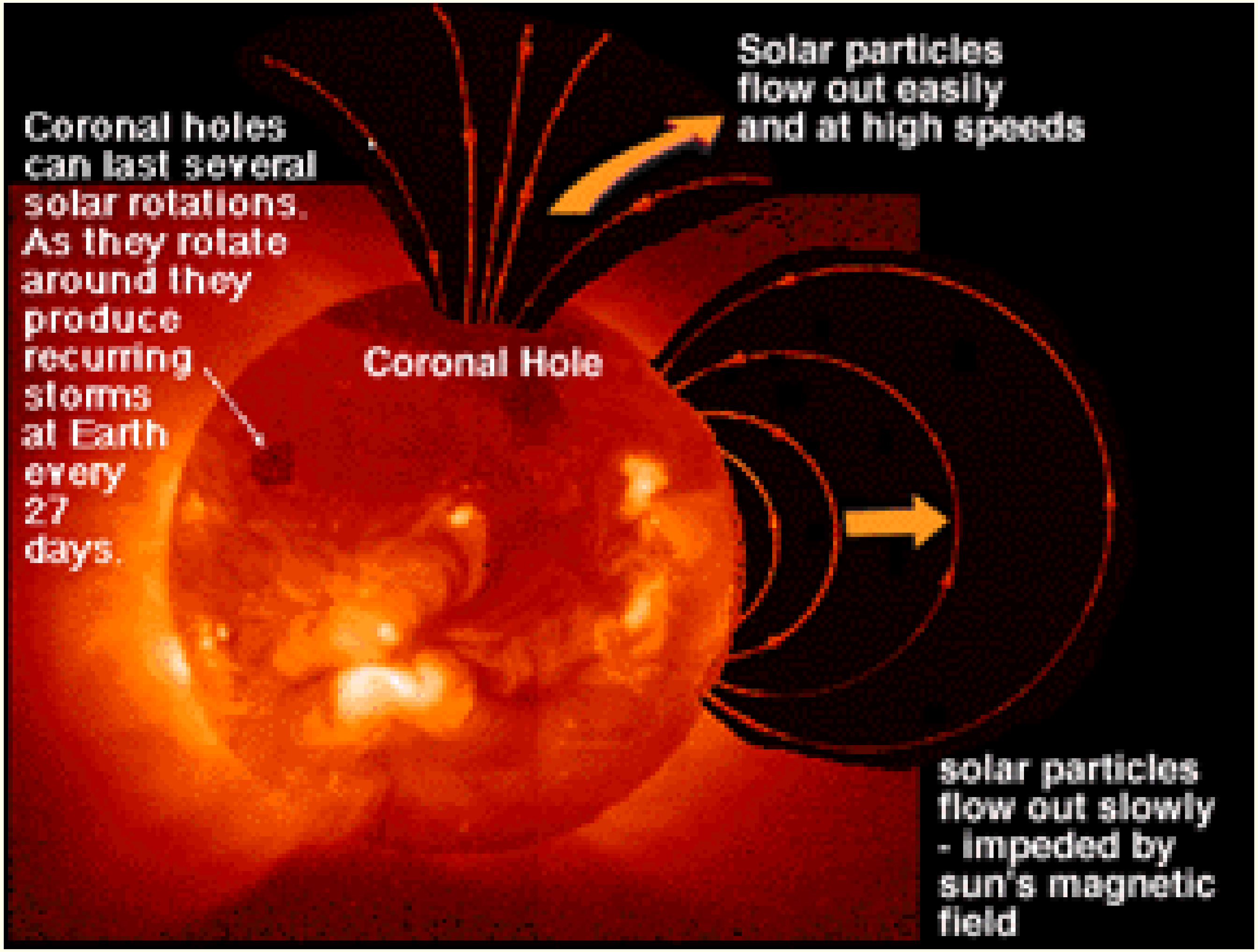


Main phase: Several particle injections increase the ring current.
Weakening of B

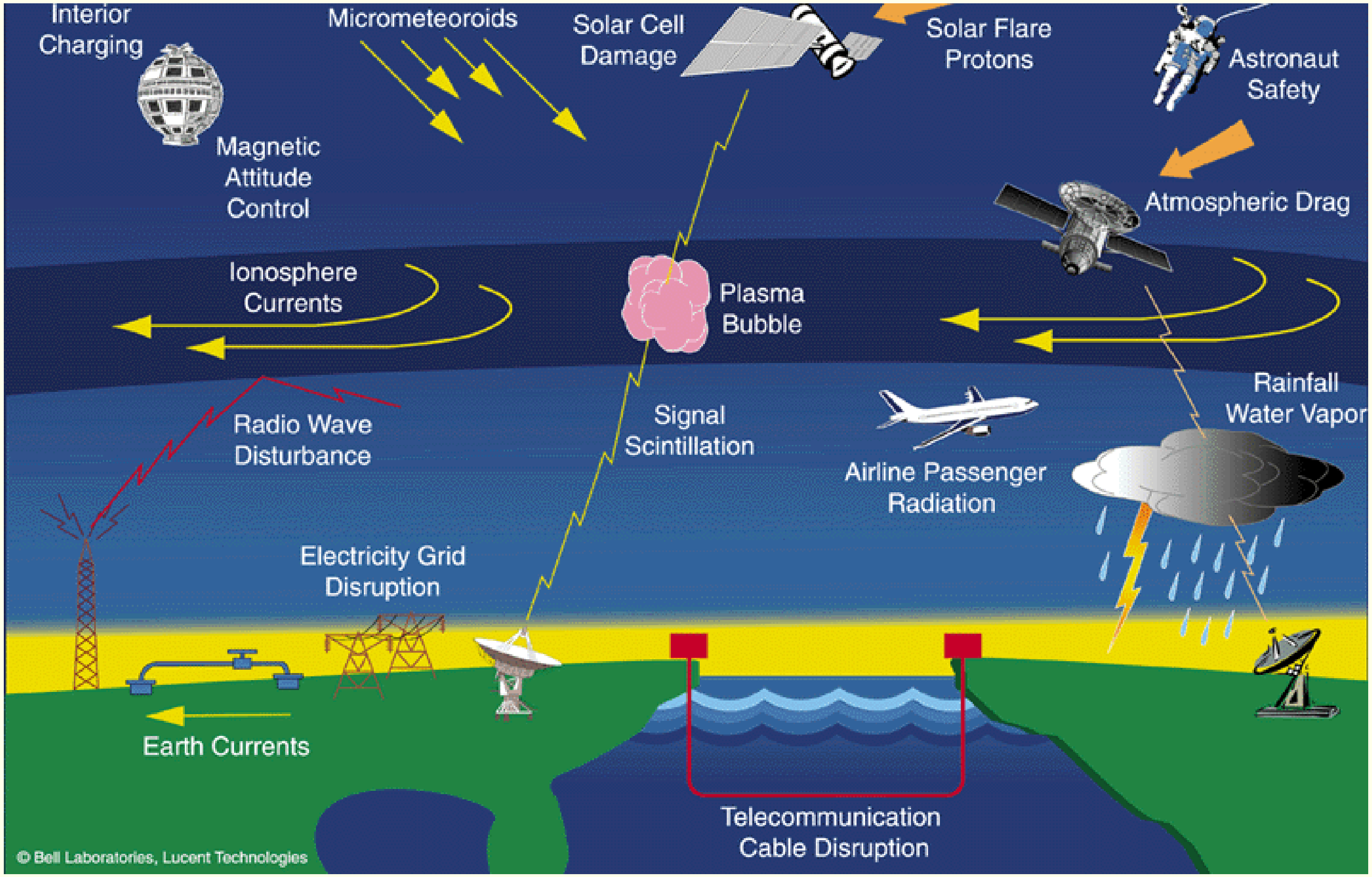
Initial phase: the magnetic cloud of the CME compresses the geomagnetic field.
Increase of B

Recovery phase: ring current returns to normal strength.
Recovery of B

Periodic geomagnetic activity

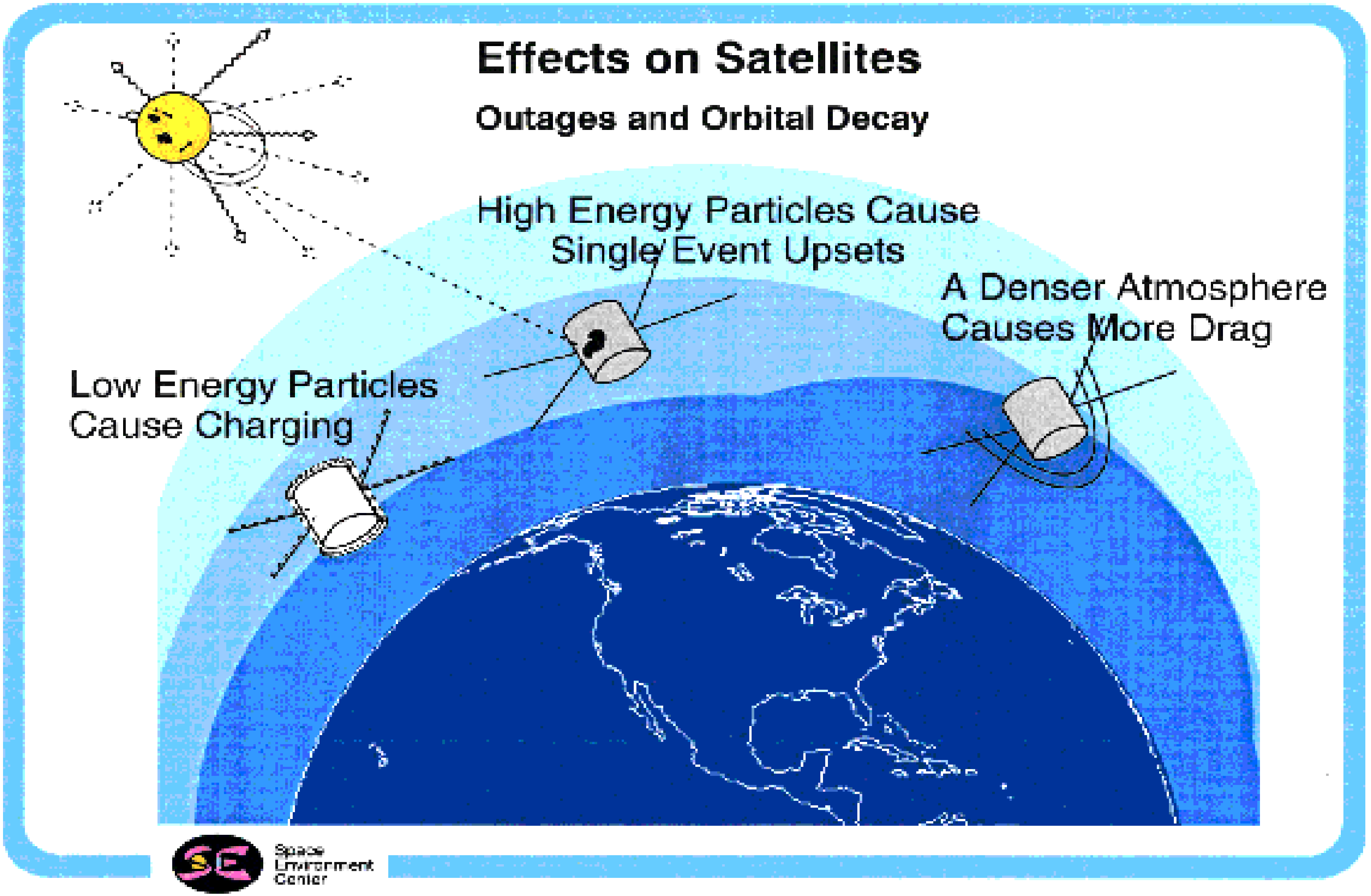


Space weather : consequences of solar and geomagnetic activity



"conditions on the Sun and in the solar wind, magnetosphere, ionosphere and thermosphere that can influence the performance and reliability of space-borne and ground-based technological systems and can endanger human life or health."

US National Space Weather Programme

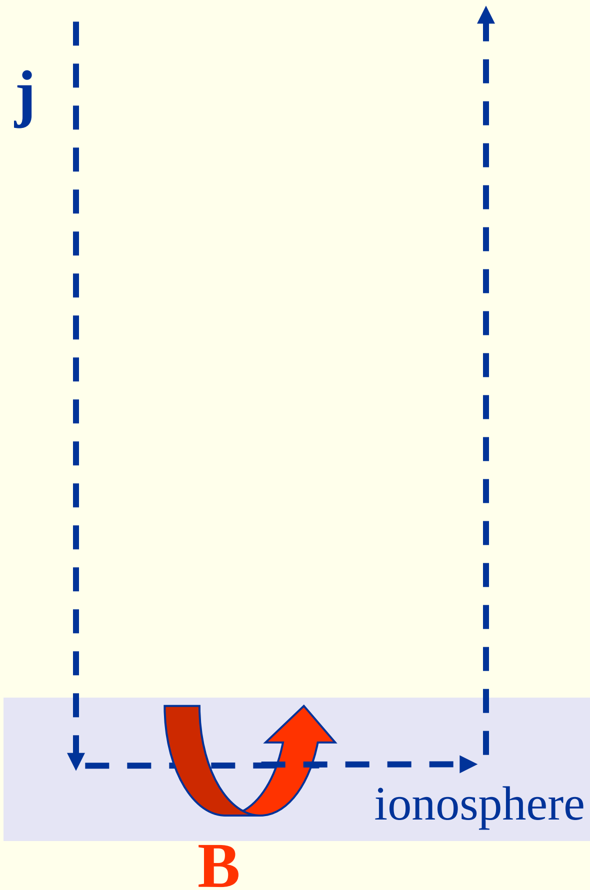

Effects on Satellites
Outages and Orbital Decay
High Energy Particles Cause
Single Event Upsets
A Denser Atmosphere
Causes More Drag
Low Energy Particles
Cause Charging
Space
Environment
Center

Damage To Solar Panels

Satellite power budgets can be very tight so degradation in solar panel performance is a serious issue.

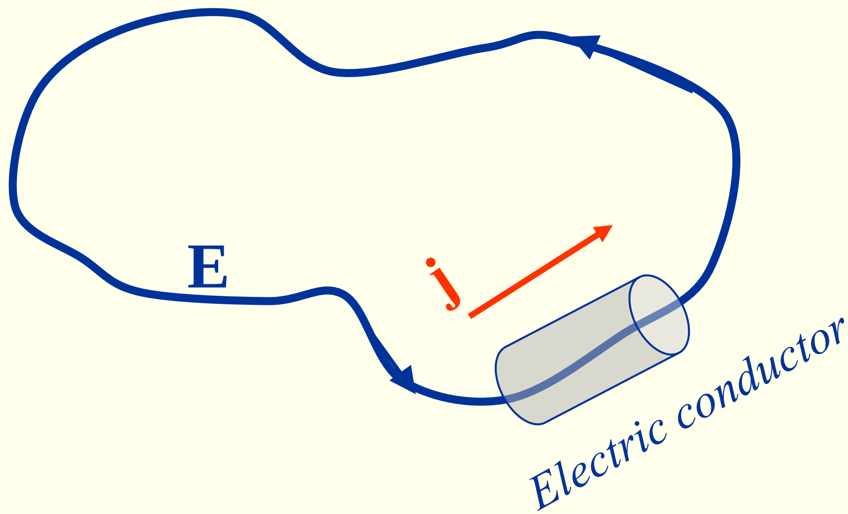
The damage is done by energetic particles which penetrate the surface of the panel and deposit a significant amount of energy inside the solar cells. This displaces the atoms within the cells and causes a loss in efficiency.

GIC – Geomagnetically Induced Currents



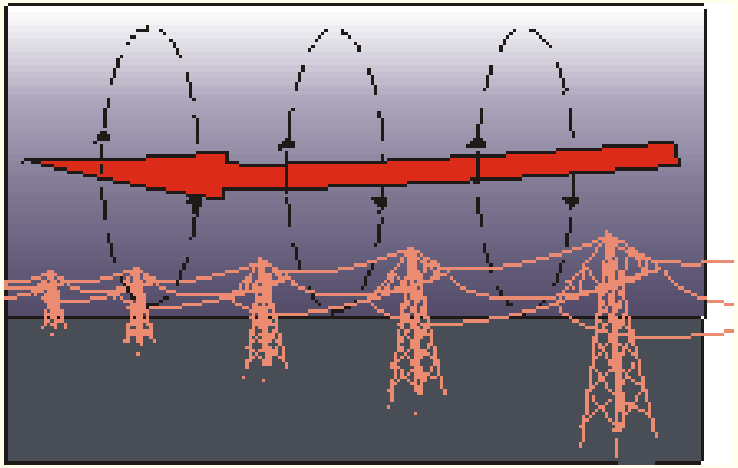
$$\frac{\partial B}{\partial t} = -\nabla \times \mathbf{E}$$

Faraday's law



GIC – Geomagnetically Induced Currents

Can damage electric power grids

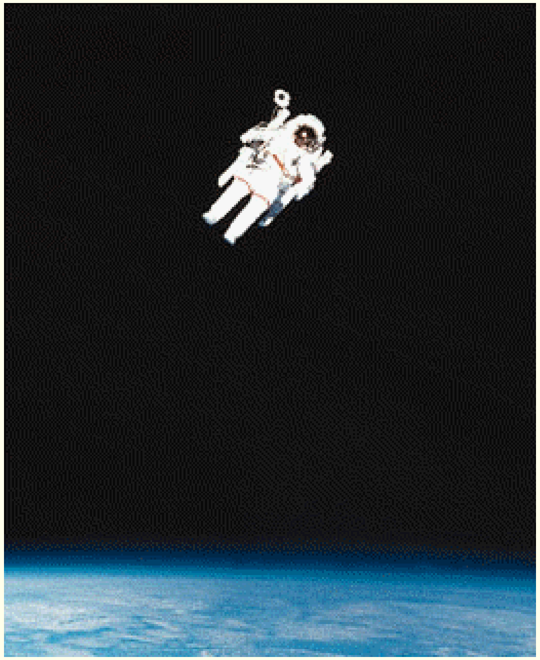


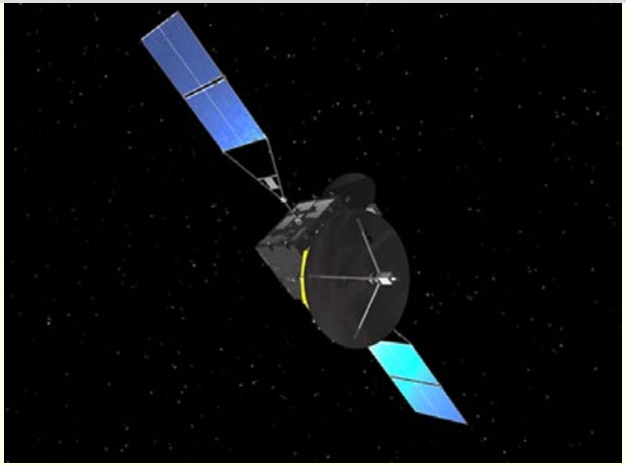
Induced currents is pipelines increase corrosion.

Highly energetic particles

- Particles in the radiation belts.
- Particles from solar activity (solar flares, CME)
- Cosmic radiation

Danger to astronauts





Disturb or damage electronics on satellites and aeoreplanes.

Increase the rate of ionization in lower D region and thus increases absorption of radio waves.

Space weather on the internet

www.spaceweather.com

www.swpc.noaa.gov/SWN (Space Weather Prediction Centre)

Last Minute!

Last Minute!

- What was the most important thing of today's lecture? Why?
- What was the most unclear or difficult thing of today's lecture, and why?
- Other comments