Space physics EF2240

Tomas Karlsson

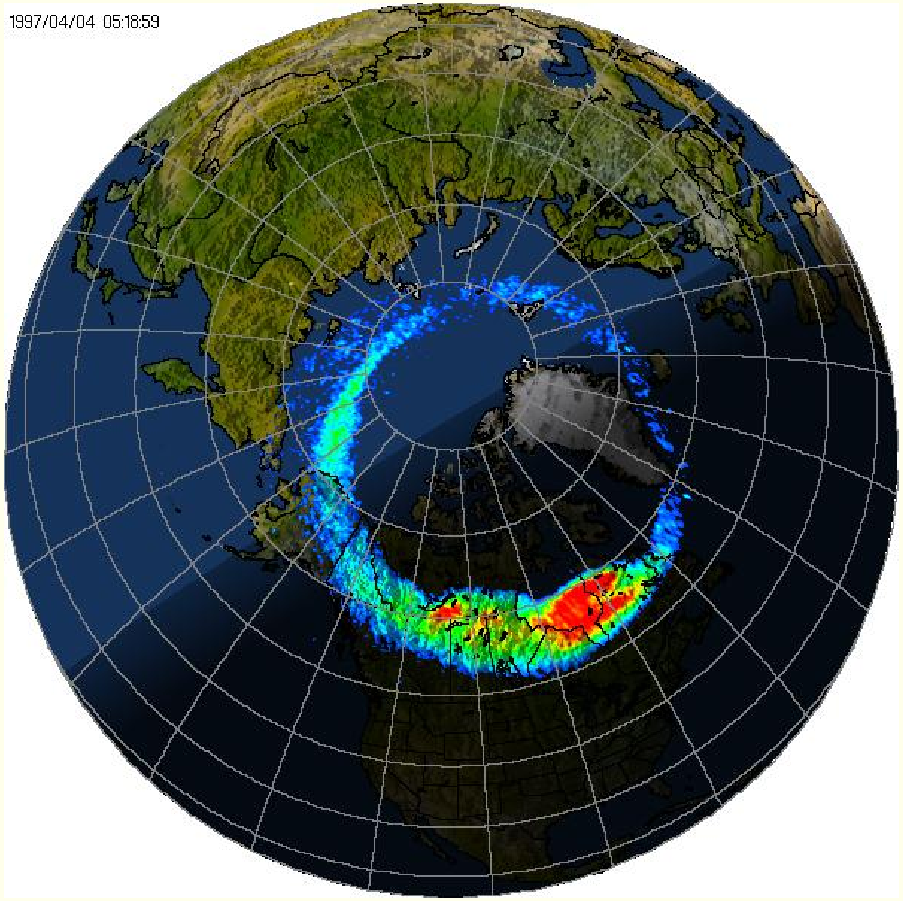
Space and Plasma Physics

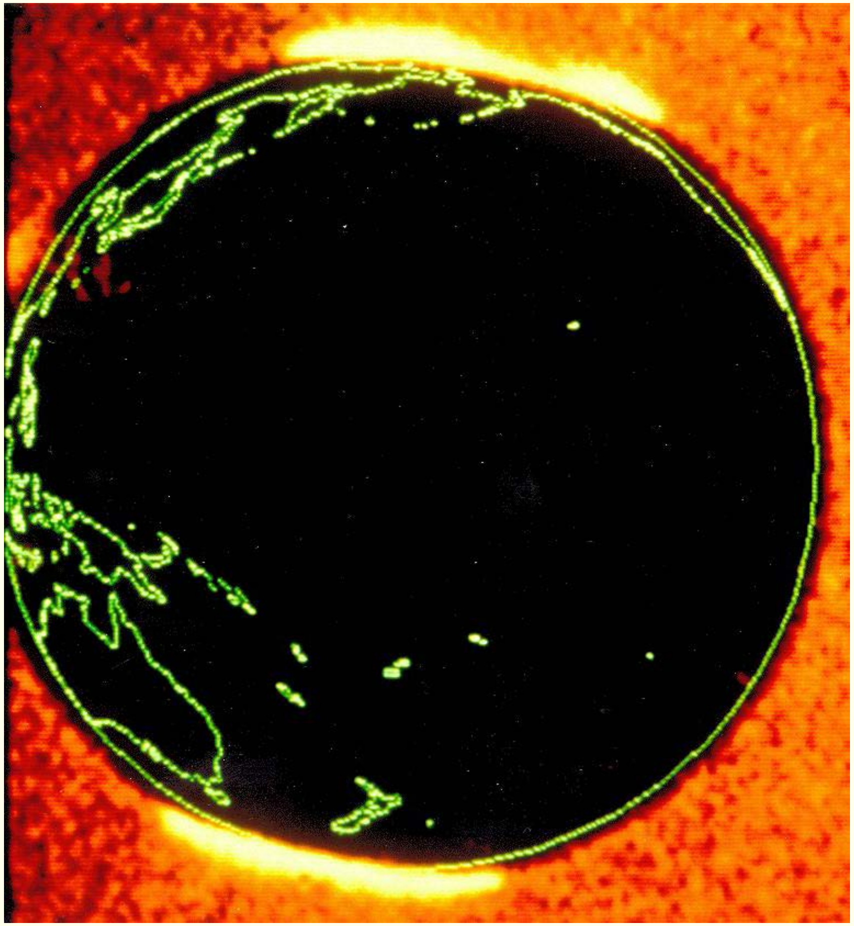
School of Electrical Engineering

What is this?

Clue...

The auroral ovals





Today's lecture

- Course info
- Overview of course content
- Definition of plasma
- Solar interior and atmosphere

Steps to take to take the course

1) Make sure you have signed up for the course.

   If you haven't: contact your Masters coordinator or studievägledare

2) Register for the course! (My Pages)

   You have to do this yourself!

Definition of Space Physics

- **Studies of space in Earth's vicinity with the help of *in situ* measurements (unique for this area, cf. astronomy and astrophysics).**

- **More than 99% of matter in space is in the *plasma* state.**

- **Alternative names:**
  - ***Space plasma physics***
  - ***Solar-terrestrial physics (incl. space weather)***

Schedule

10×2 h Lectures

6×2 h Tutorials

Activity	Date	Time	Room	Subject	Litterature
L1	2/9	10-12	Q33	Course description, Introduction, The Sun 1, Plasma physics 1	**CGF** Ch 1, 5, (p 110-113)
L2	4/9	10-12	Q21	The Sun 2, Plasma physics 2	**CGF** Ch 5 (p 114-121), 6.3
L3	8/9	13-15	Q36	Solar wind, The ionosphere and atmosphere 1, Plasma physics 3	**CGF** Ch 6.1, 2.1-2.6, 3.1-3.2, 3.5, **LL** Ch III, Extra material
T1	10/9	10-12	Q33	Mini-group work 1	
L4	15/9	13-15	Q31	The ionosphere 2, Plasma physics 4	**CGF** Ch 3.4, 3.7, 3.8
T2	17/9	10-12	Q33	Mini-group work 2	
L5	19/9	15-17	Q31	The Earth's magnetosphere 1, Plasma physics 5	**CGF** 4.1-4.3, **LL** Ch I, II, IV.A
L6	23/9	8-10	Q31	The Earth's magnetosphere 2, Other magnetospheres	**CGF** Ch 4.6-4.9, **LL** Ch V.
T3	24/9	14-16	Q21	Mini-group work 3	
L7	29/9	11-13	Q36	Aurora, Measurement methods in space plasmas and data analysis 1	**CGF** Ch 4.5, 10**, LL** Ch VI, Extra material
T4	1/10	15-17	Q31	Mini-group work 4	
L8	2/10	15-17	Q34	Space weather and geomagnetic storms	**CGF** Ch 4.4, **LL** Ch IV.B-C, VII.A-C
L9	8/10	13-15	Q36	Interstellar and intergalactic plasma, Cosmic radiation, Swedish and international space physics research.	**CGF** Ch 7-9
T5	9/10	15-17	Q31	Mini-group work 5	
L10	13/10	15-17	Q33	Guest lecture (preliminary): Swedish astronaut Christer Fuglesang	
T6	16/10	10-12	Q36	Round-up	
Written exami-nation	30/10	8-13	M33, M37, M38		

L = Lecture, T = Tutorial

Guest lecturer

Swedish astronaut Christer Fuglesang
Lecture 10

Course goals

At the end of the course you should be able to:

• define what a plasma is, and classify various types of plasma.

• describe the plasma physical properties of various regions of space, with emphasis of the near-earth region.

• explain how some important plasma populations in the solar system (e.g. Earth's ionosphere and magnetosphere) get their basic properties and how these properties can vary between the planets.

• make order of magnitude estimates of some properties of space plasmas and space physics phenomena, for example the power dissipated in the aurora or the magnitude of electric currents floating from the magnetosphere into the ionosphere.

• make simple analyses of measurement data from satellites and ground-based instruments. (E.g. calculate currents in space from magnetometer data.)

• make simple models of some space physics phenomena by applying basic physical laws expressed with simple mathematics. (An example would be to model the basic shape of the magnetosphere or estimate the temperature of a sunspot.)

• describe to interested laymen or "the man in the street" what we can learn from space physics and how it affects our everyday life (for example by various space weather phenomena.)

Examination

1. Written examination (open book*), 30/10

100 p

2. Continous examination (mini-group works)

25 p

Grades:

A:	111-125 p
B:	96-110 p
C:	81-95 p
D:	66-80 p
E:	50-65 p

(Fx)

Written examination,
30/10 2014, 08.00-13.00, M33, M37, M38

You may bring:

- all the course material
- any notes you have made
- pocket calculator
- mathematics and physics formula books or your favourite physics book
- formula sheet

(No computers are allowed, due to the possibility to communicate with the outside world.)

Approx. 5 different problems (which may contain sub-problems).

The character of the problems is such that to get a high score you will have to show that you have obtained a certain course goal, e.g. to make a reasonable order of magnitude estimate or figure out a simple model for some space physics phenomenon.

Continous examination
Mini-group works

5 mini-group works
(5×5 p = 25 p)

Approx. 1 h during Tutorials 1-5

- *A problem similar to those on the written examination is given*
- *Groups of 3 (randomized).*
- *Elect a secretary!*
- *Write down a solution!*

Litterature

- C-G. Fälthammar, ”Space Physics” (compendium), 2nd Ed, Third Printing, 2001.
- Larry Lyons, ”Space Plasma Physics”, from *Encyclopedia of Physical Science and Technology, 3rd edition, 2002.*
- Lecture notes and extra material handed out during lectures.

Course home page

KTH Social:

https://www.kth.se/social/course/EF2240/

At the home page I will post new information continuously. Here you can also find lecture notes, exercises (and some solutions), etc.

Last Minute!

What was the most important thing of today's lecture? Why?

What was the most unclear or difficult thing of today's lecture, and why?

Other comments:

Study the Course Description carefully!

EF22445 Space Physics II
7.5 ECTS credits, P2

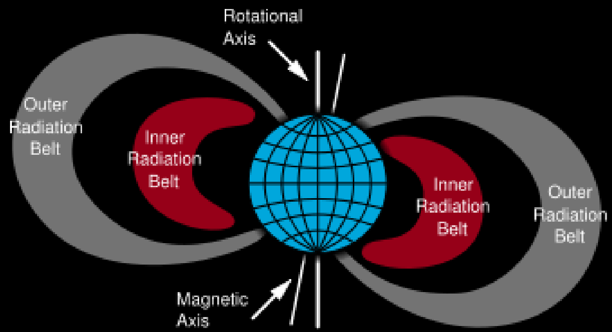
- shocks and boundaries in space
- solar wind interaction with magnetized and unmagnetized bodies
- reconnection
- sources of magnetospheric plasma
- magnetospheric and ionospheric convection
- auroral physics
- storms and substorms
- global oscillations of the magnetosphere

Courses at the Alfvén Laboratory

EF2260 SPACE ENVIRONMENT AND SPACECRAFT ENGINEERING , 6 ECTS credits, period 2

- environments spacecraft may encounter in various orbits around the Earth, and the constraints this places on spacecraft design
- basic operation principles underlying the thermal control system and the power systems in spacecraft
- measurements principles in space

The Astrid-2 satellite



Radiation environment in near-earth space

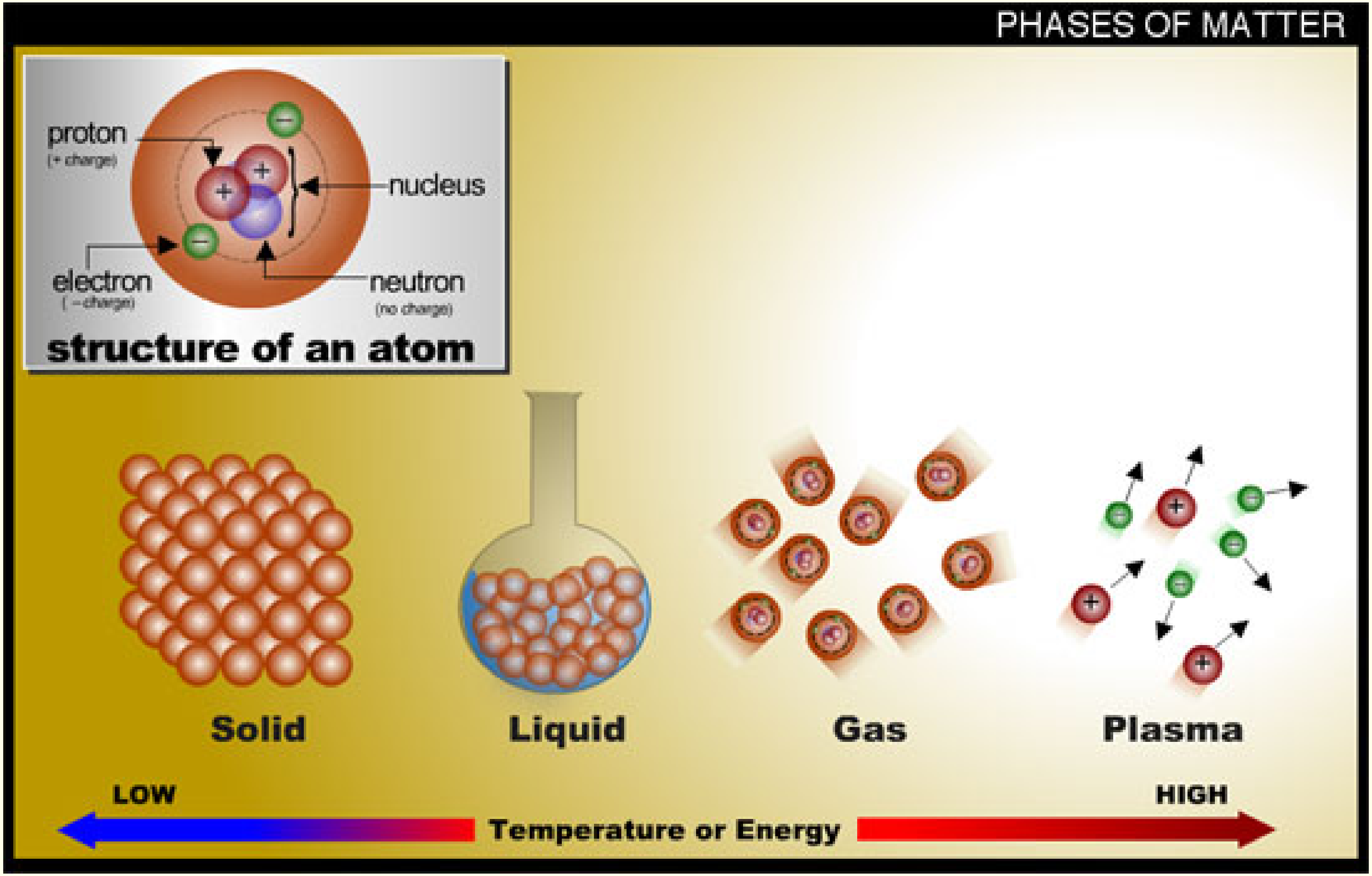
Projects:

- Design power supply for spacecraft
- Study of radiation effects on electronics

What is a plasma?

Where in the universe can you find it?

Plasma



PHASES OF MATTER
proton
(+ charge)
nucleus
electron
(– charge)
neutron
(no charge)
structure of an atom
Solid
Liquid
Gas
Plasma
LOW
HIGH
Temperature or Energy

Plasma

E.g. Water

1 eV ~ 11600 K

T

Plasma — ~10^5 K *(H: 13.60 eV, ionization*
O: 13.61 eV, ionization)

~10^4 K *(5 eV, dissociation)*

Gas

373 K

Liquid

273 K

Solid

Definition: A plasma is an ionized gas, showing collective behaviour.

"Fourth state of matter"

Somewhat misleading:

- *No phase transition*
- *Ionization can be caused by other mechanisms than heating, e.g. UV radiation.*

Energy - temperature

Average energy of molecule/atom:

$$E = \frac{3}{2} k_B T \quad \Rightarrow$$

$$T = \frac{2E}{3k_B}$$

$$1\ \text{eV} = 1.6 \cdot 10^{-19}\ \text{J} \Rightarrow$$

$$T = \frac{2E}{3k_B} = \frac{2 \cdot 1.6 \cdot 10^{-19}\ \text{J}}{3 \cdot 1.38 \cdot 10^{-23}\ \frac{\text{J}}{\text{K}}} = 7729\,\text{K}$$

But beware!

In plasma physics, usually:

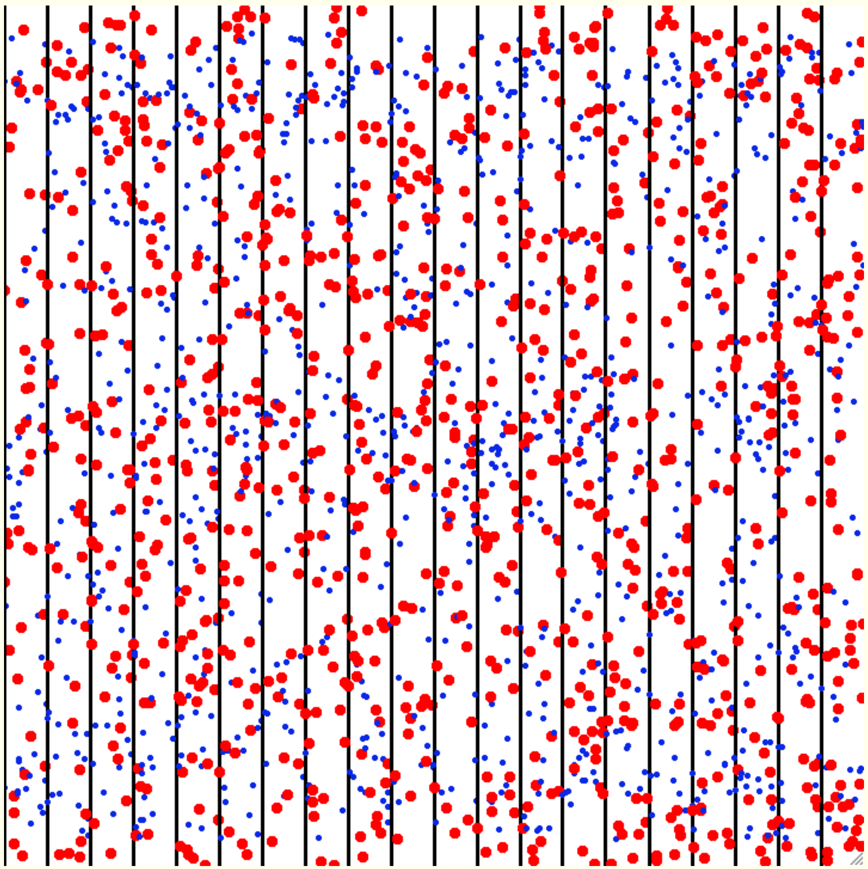
$$E = \cancel{\frac{3}{2}} k_B T \quad \Rightarrow$$

$$T = \frac{E}{k_B}$$

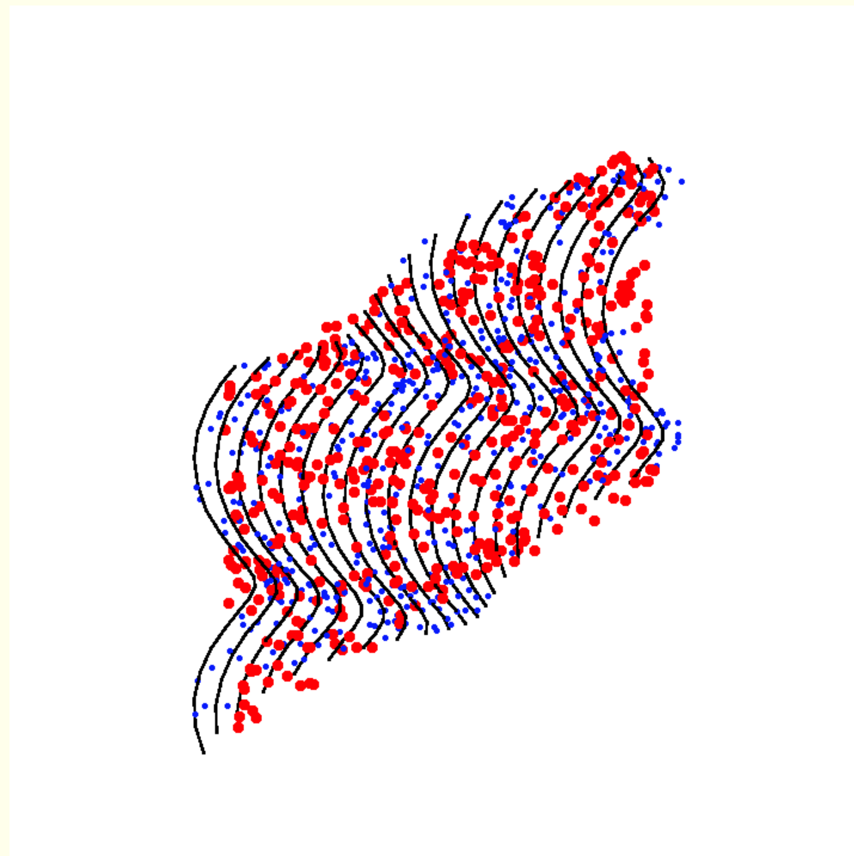
$$1\ \mathrm{eV} = 1.6 \cdot 10^{-19}\ \mathrm{J} \Rightarrow$$

$$E = k_B T = \frac{1.6 \cdot 10^{-19}\,\mathrm{J}}{1.38 \cdot 10^{-23}\,\frac{\mathrm{J}}{\mathrm{K}}} = 11594\,\mathrm{K}$$

Example of collective behaviour: Plasma waves

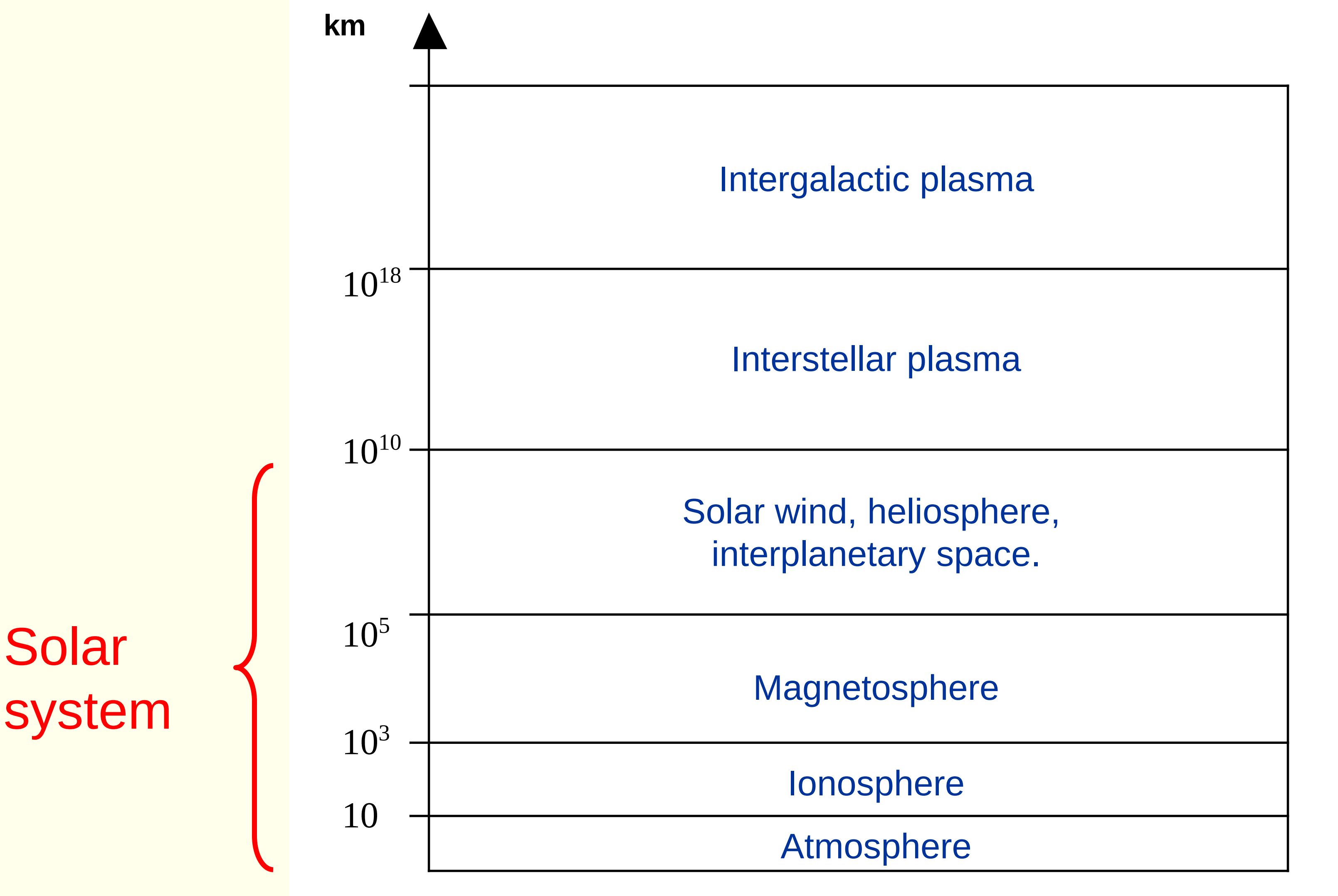


Electron plasma waves

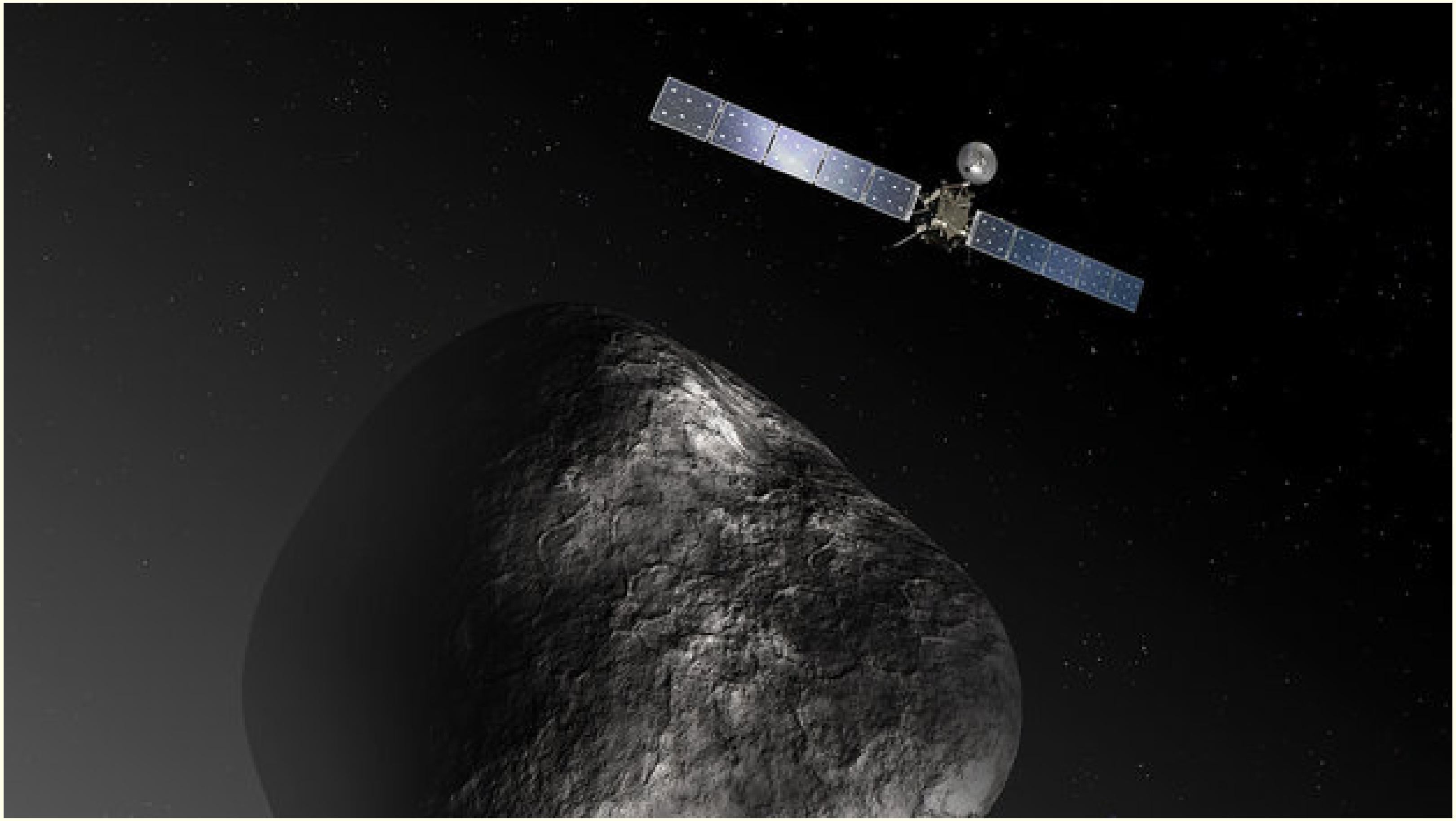


Whistler waves

From atmosphere to intergalactic plasma!



The Rosetta mission to comet 67P/Churiumov-Gerasimenko



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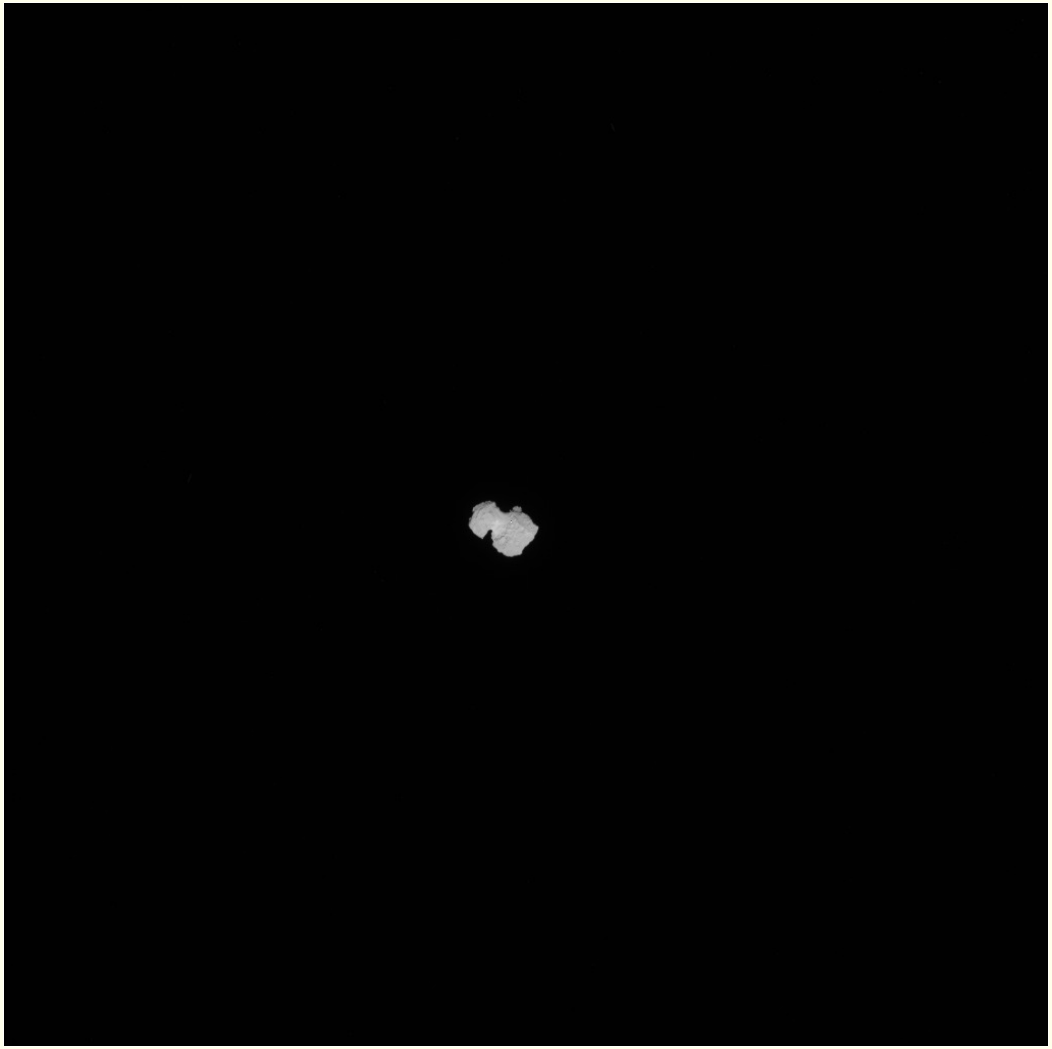
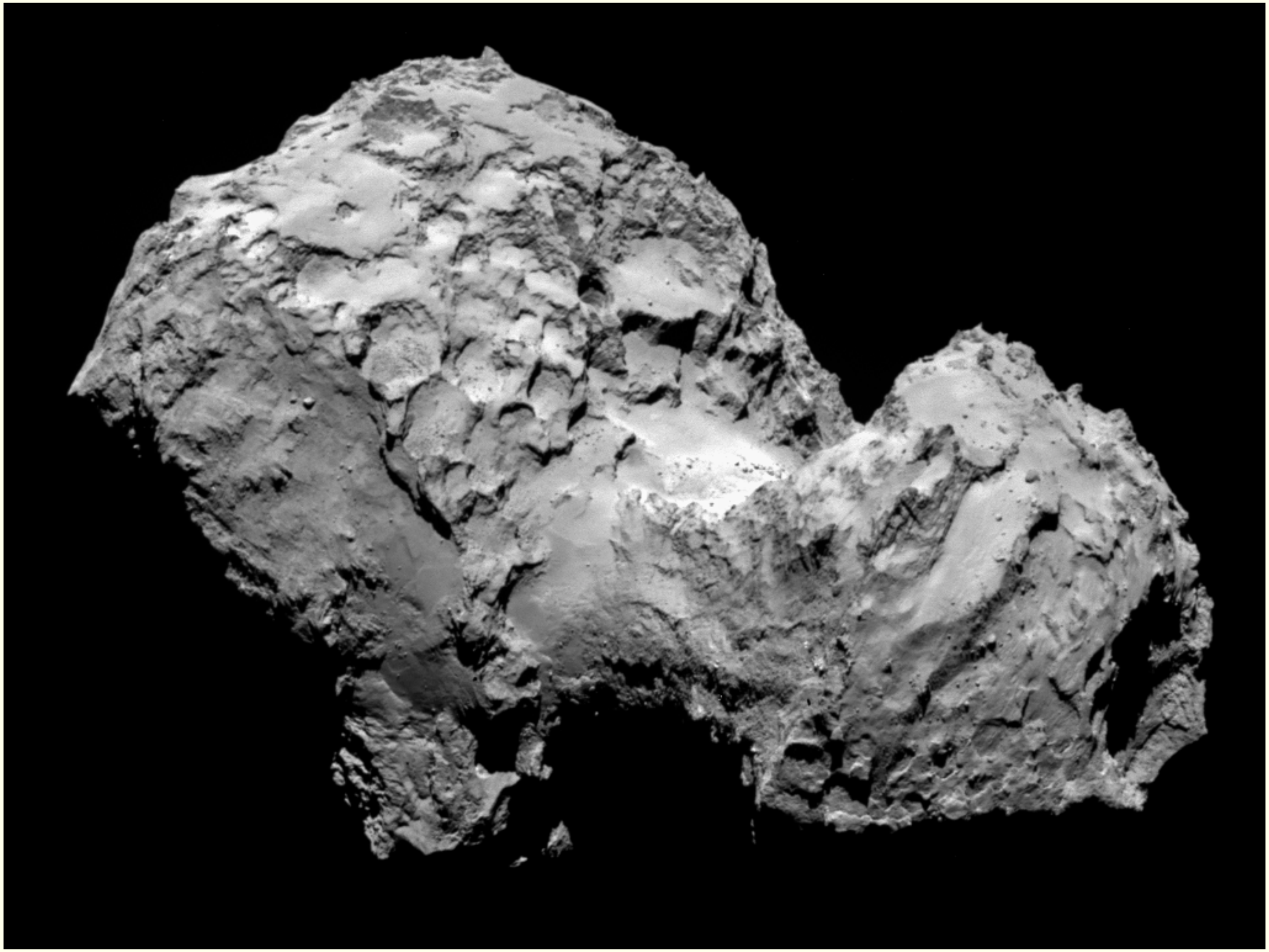
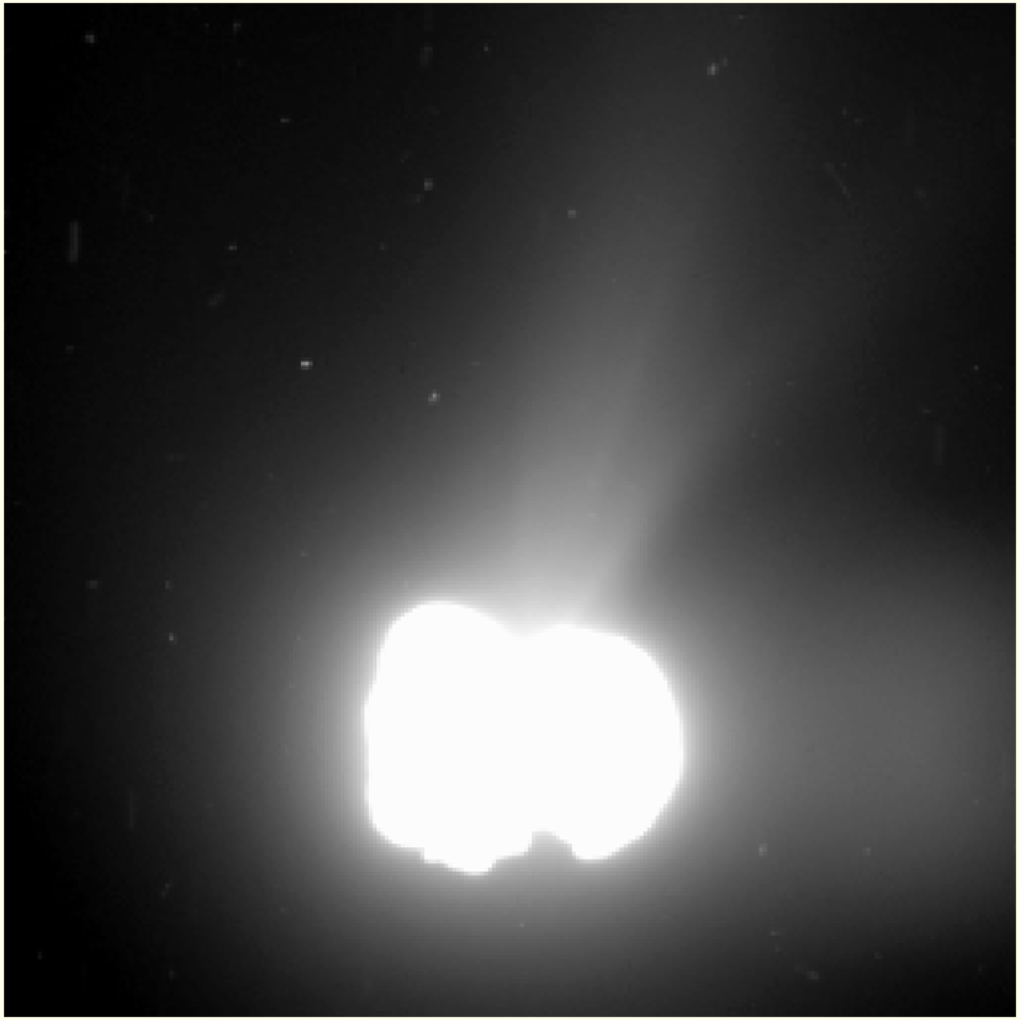


The Rosetta mission to comet 67P



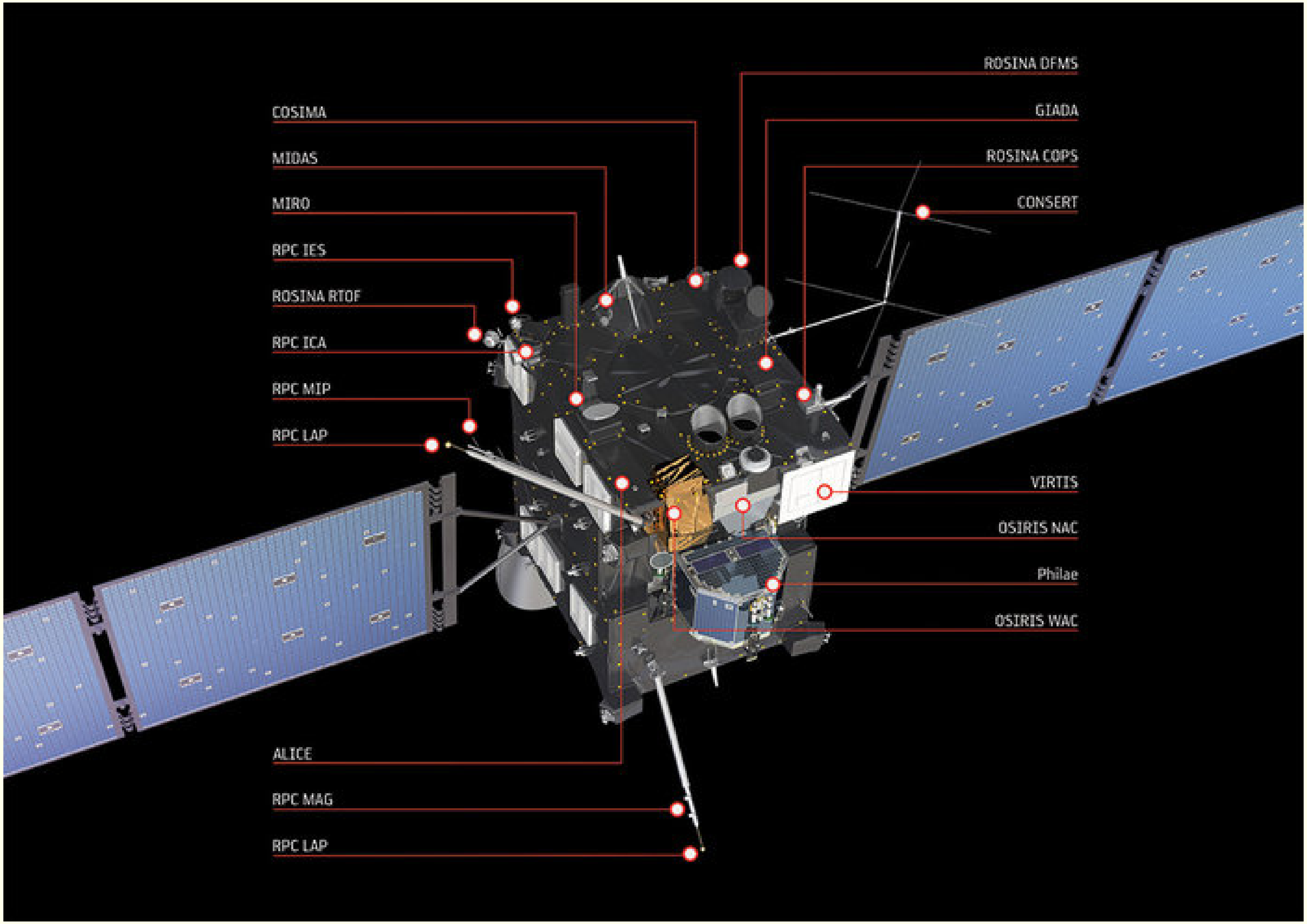
3 August 2014

The Rosetta mission to comet 67P

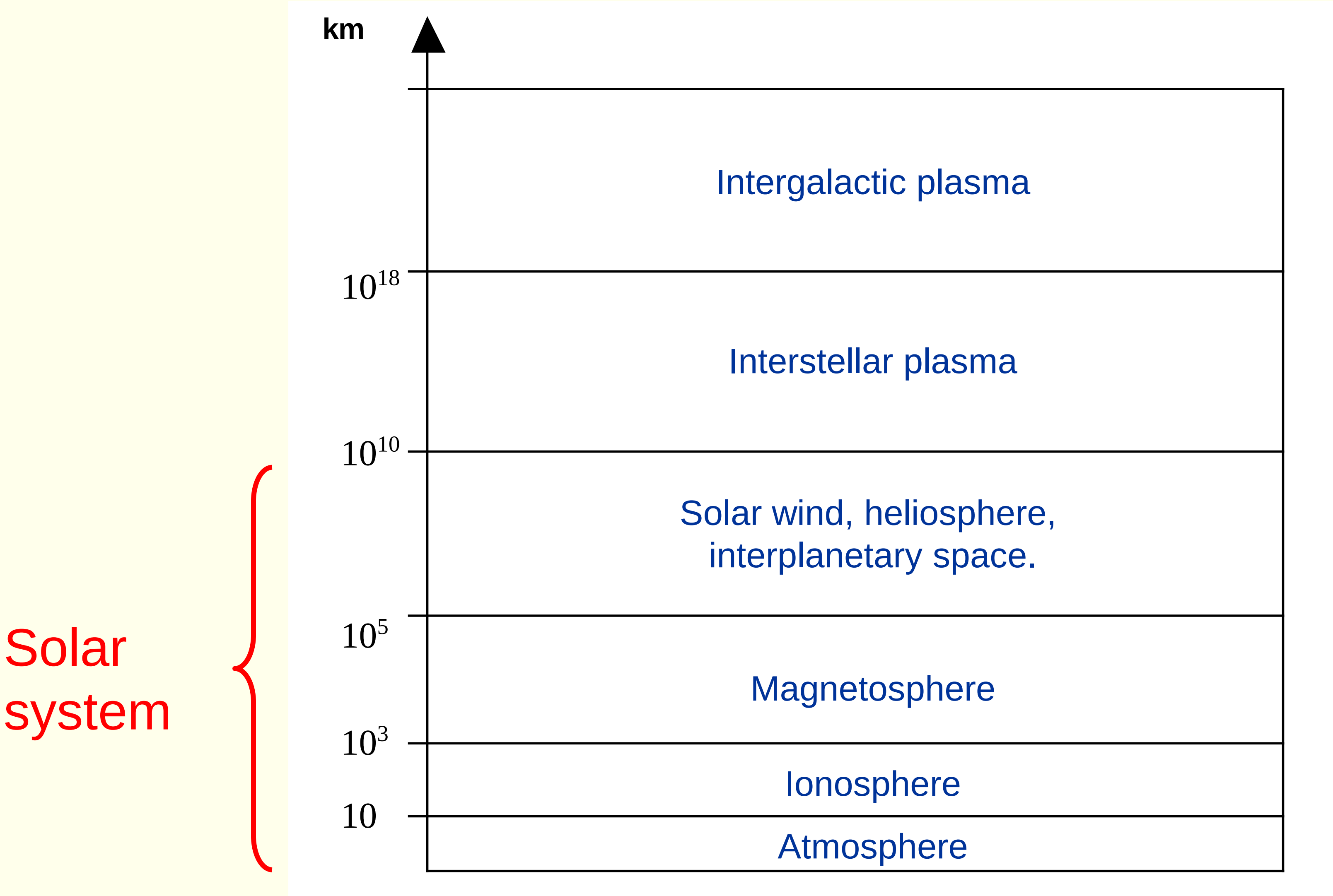


2 August 2014

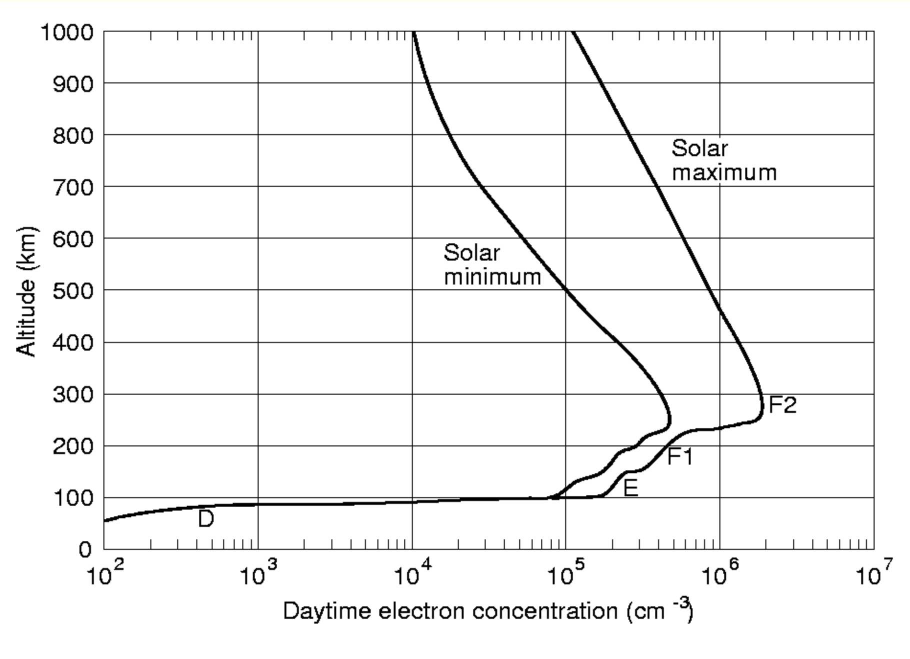
The Rosetta mission to comet 67P



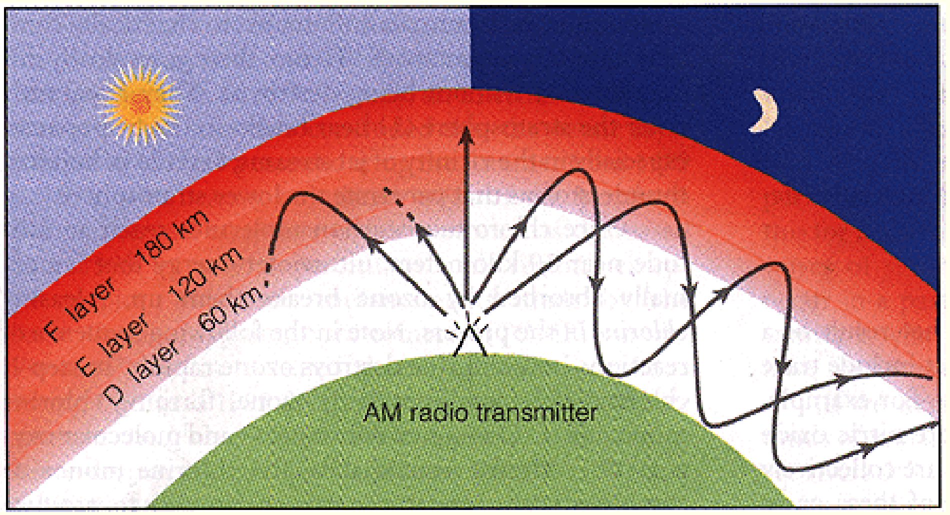
From atmosphere to intergalactic plasma!



Ionosphere

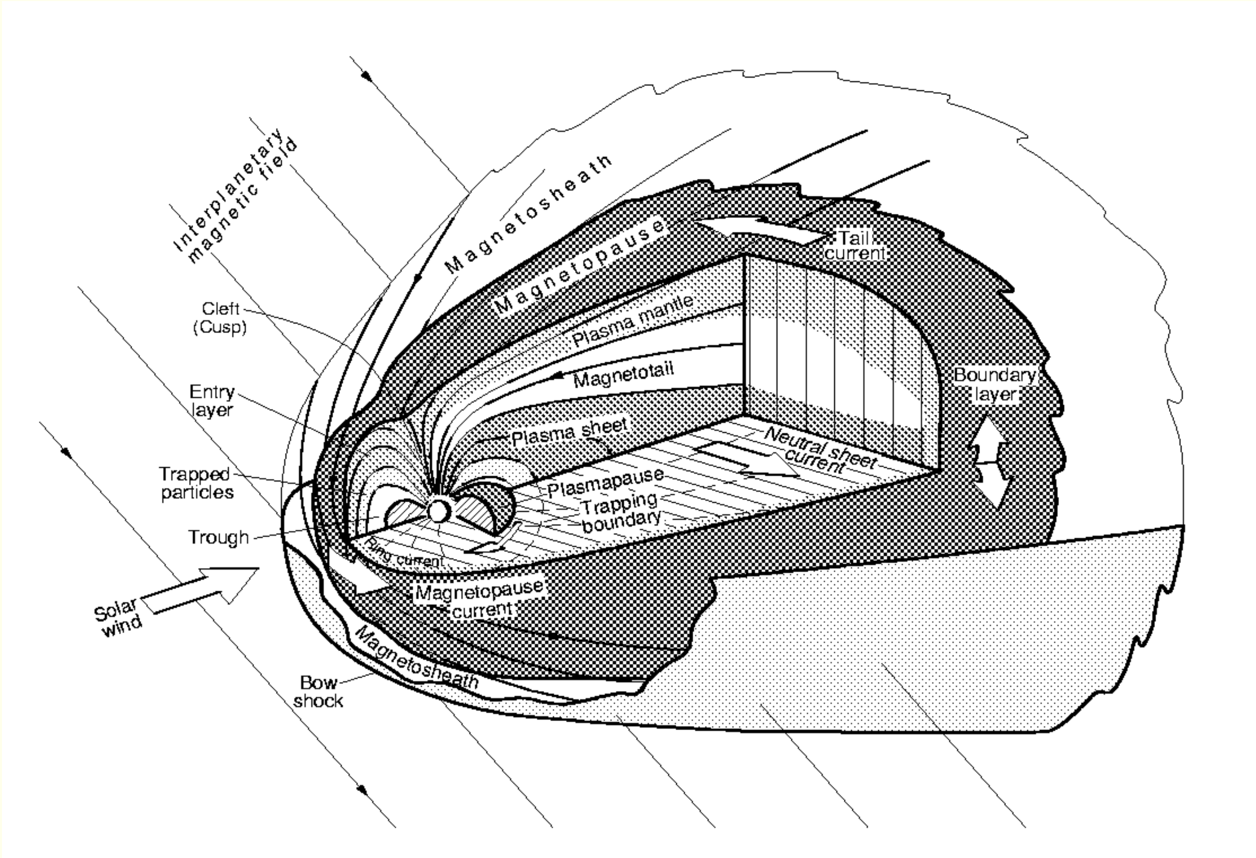


1000
900
800
700
600
500
400
300
200
100
0
Altitude (km)
10^2
10^3
10^4
10^5
10^6
10^7
Daytime electron concentration (cm $^{-3}$)
Solar maximum
Solar minimum
F2
F1
E
D



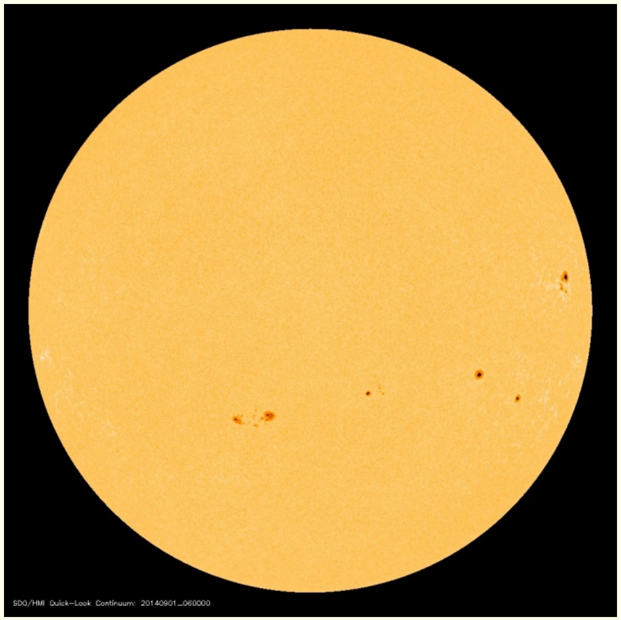
F layer 180 km
E layer 120 km
D layer 60 km
AM radio transmitter

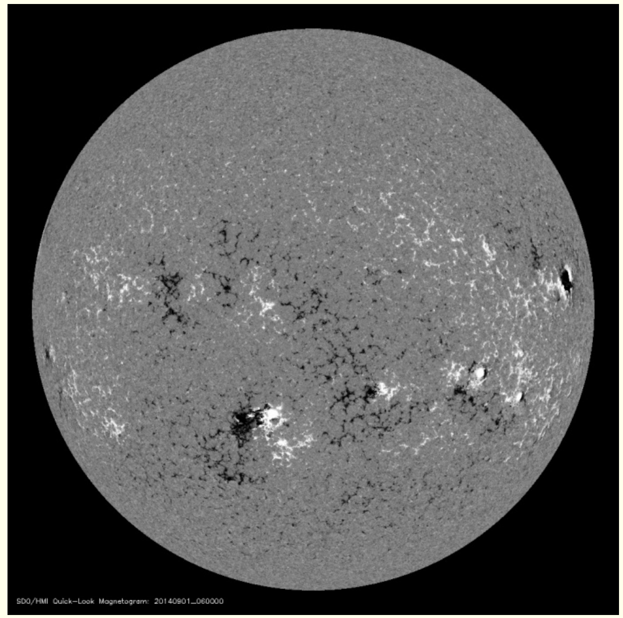
Magnetosphere

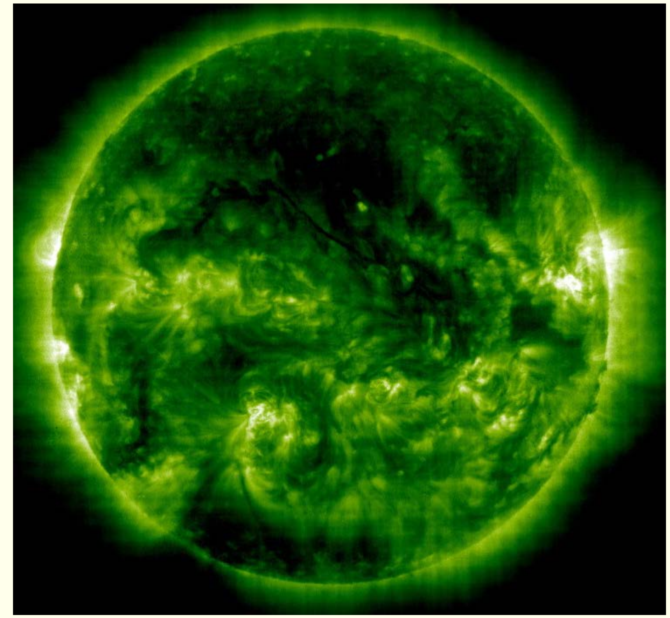


Definition: *That region in space where the geomagnetic field is the dominating magnetic field.*

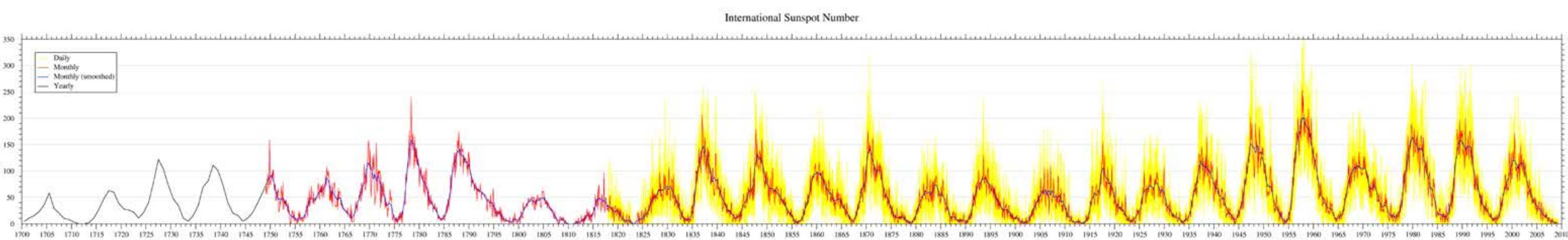
The sun (2014-09-01)
SOHO observations



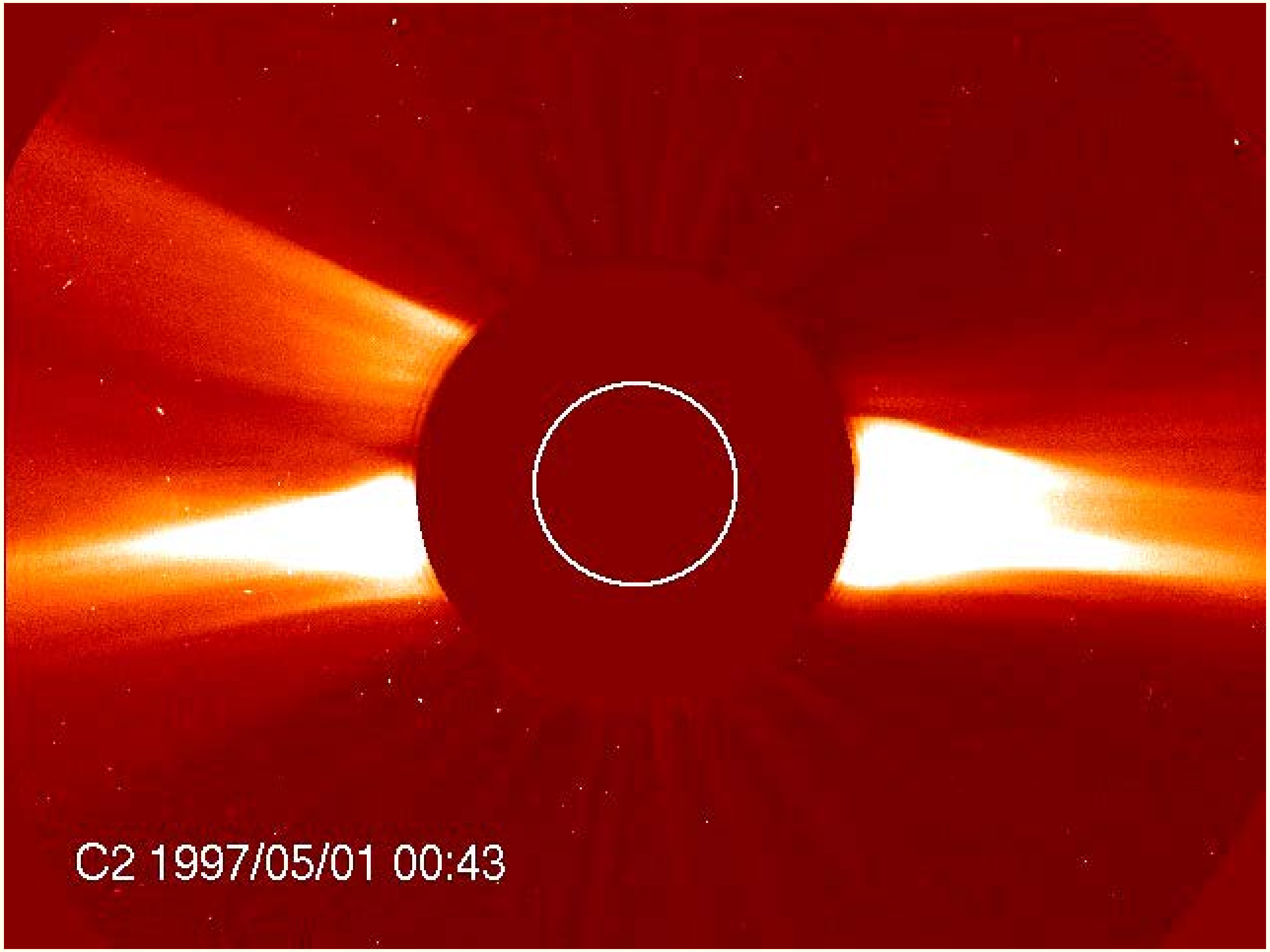




Visible light *Magnetogram* *SOHO EUV (Fe XV)*

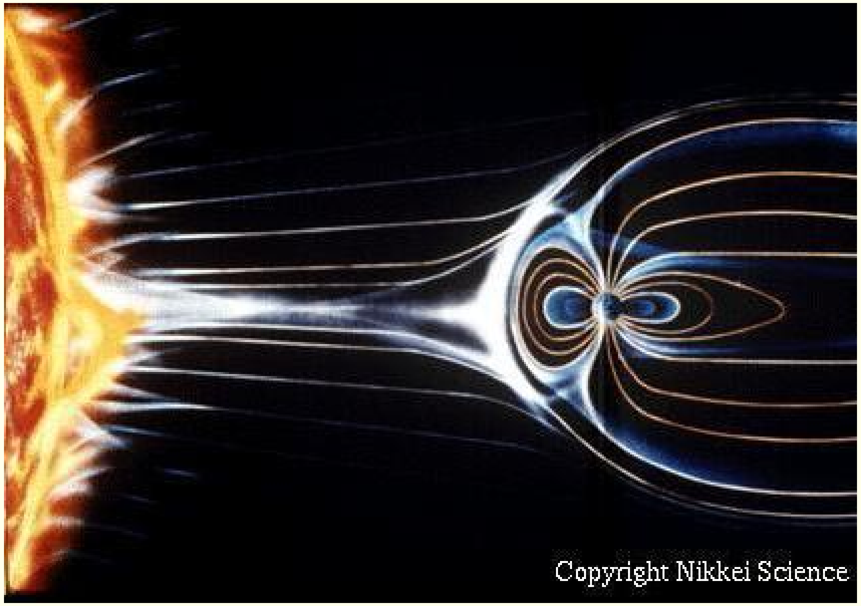


The sun, solar wind

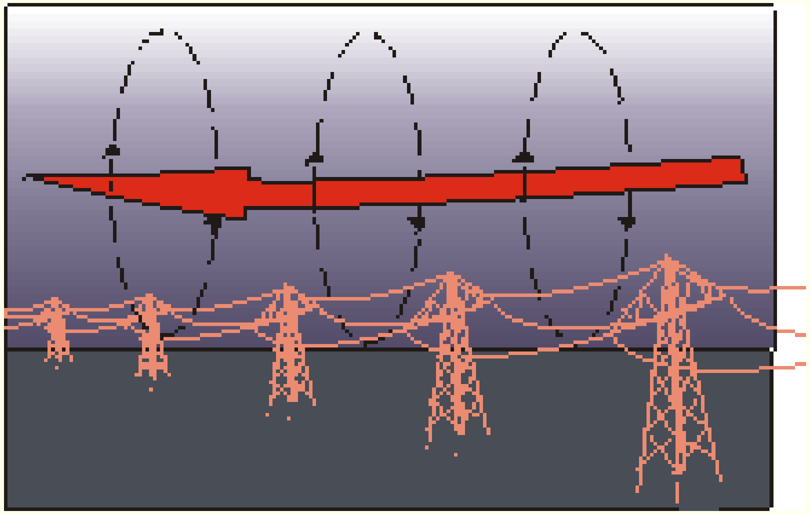


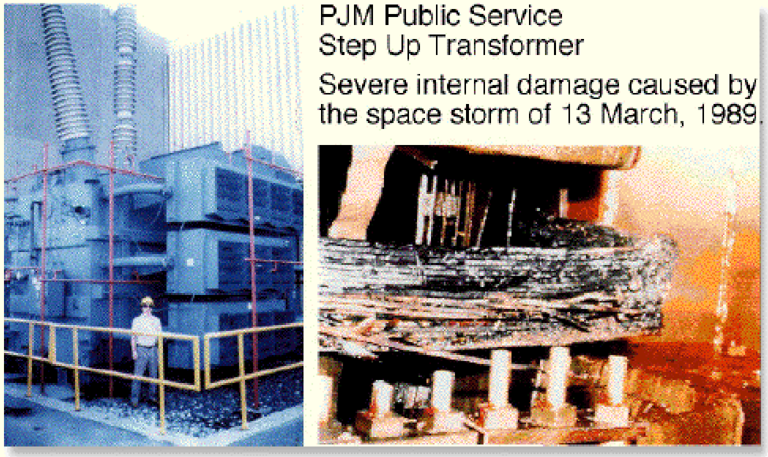
Solar and Heliospheric Observatory (SOHO),
LASCO C2 Coronagraph Movie

Solar-terrestrial interaction



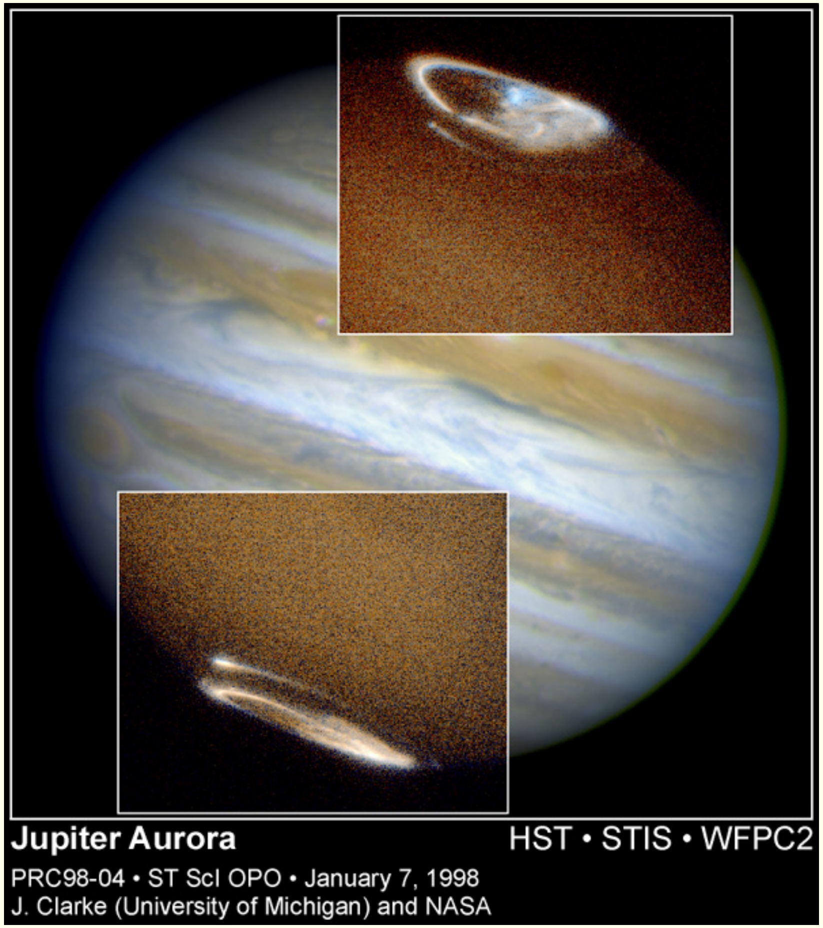
Space weather:
Geomagnetically induced currents (GIC)





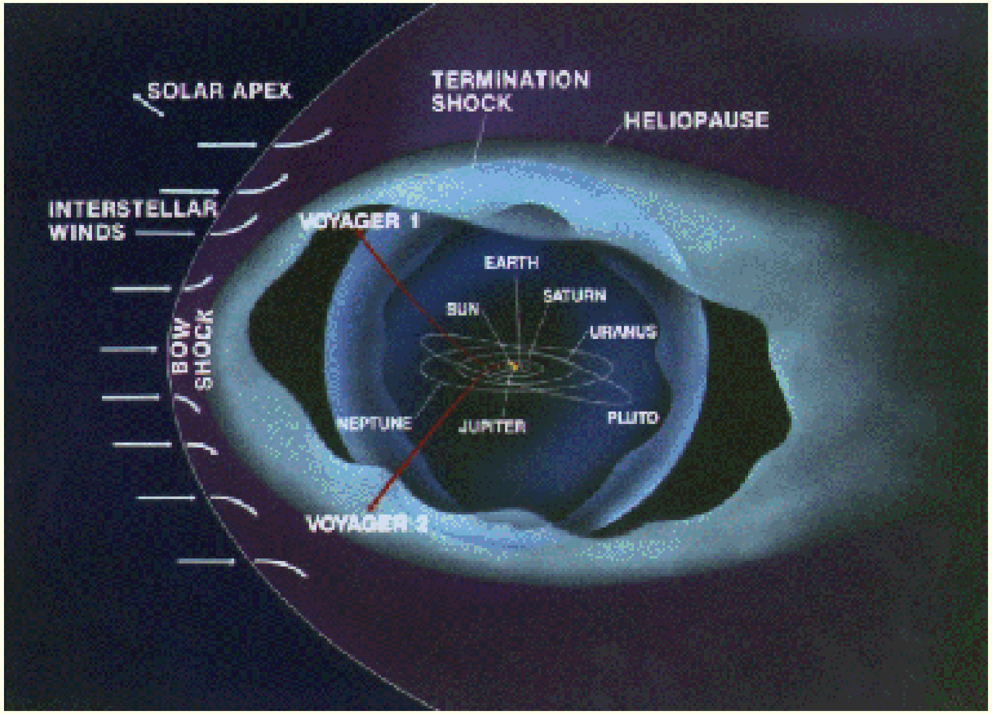
Aurora on Earth

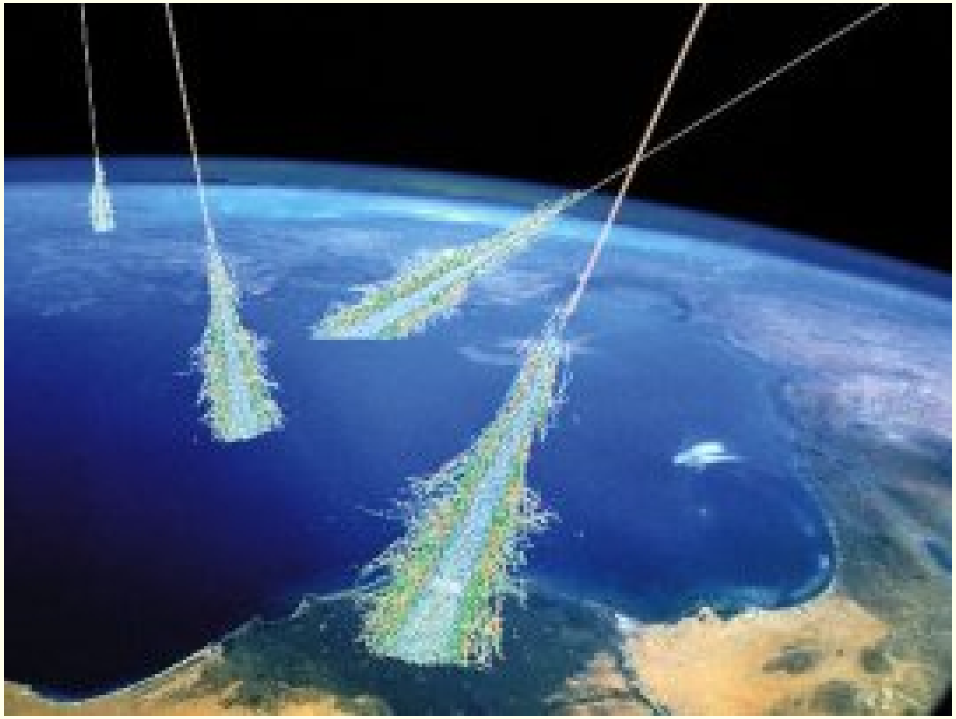
Aurora on other planets



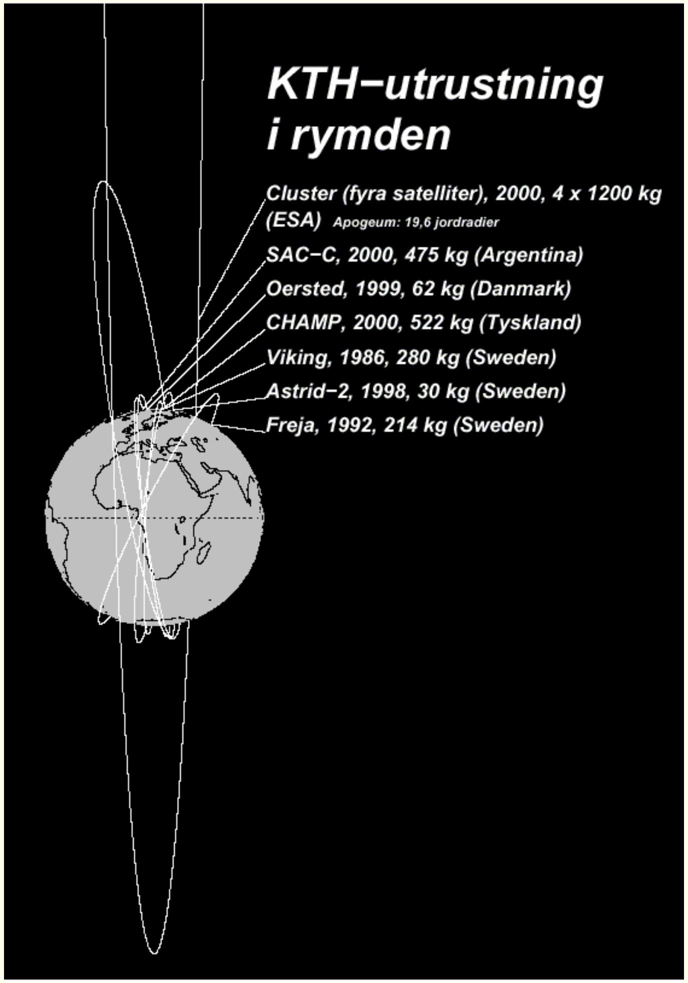
Interstellar and intergalactic plasma

Cosmic radiation

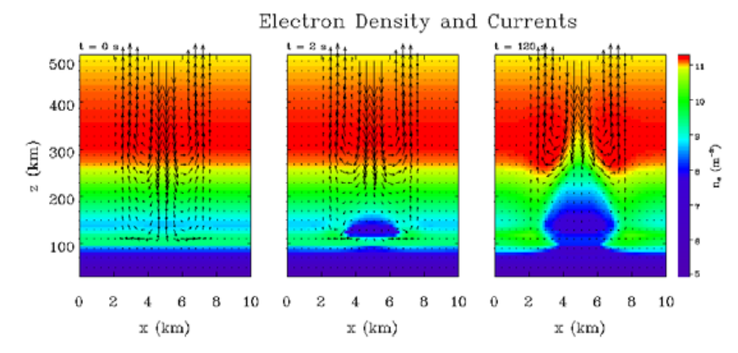




Swedish and international space physics research

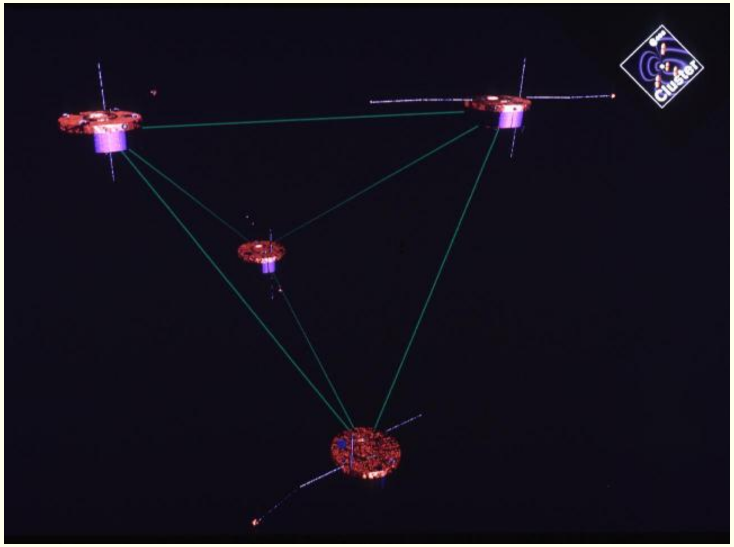


Micro satellite Astrid-2



Simulations

Cassini & Huygens at Saturn



Cluster satellites

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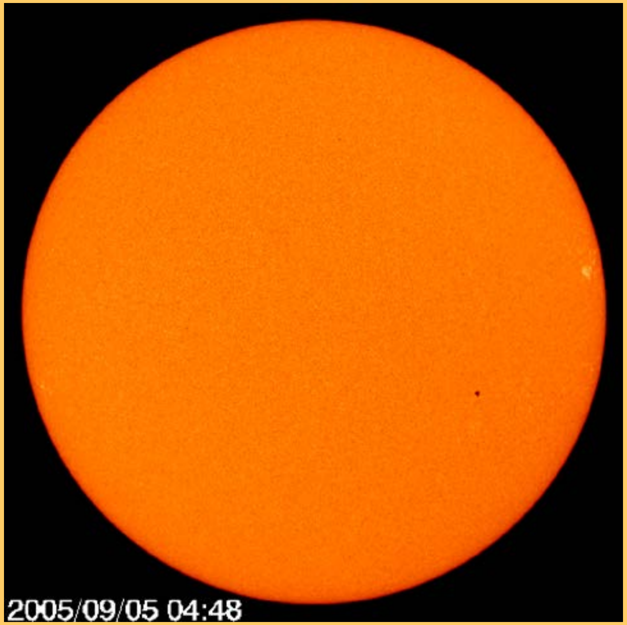
L = Lecture, T = Tutorial

Think about this:



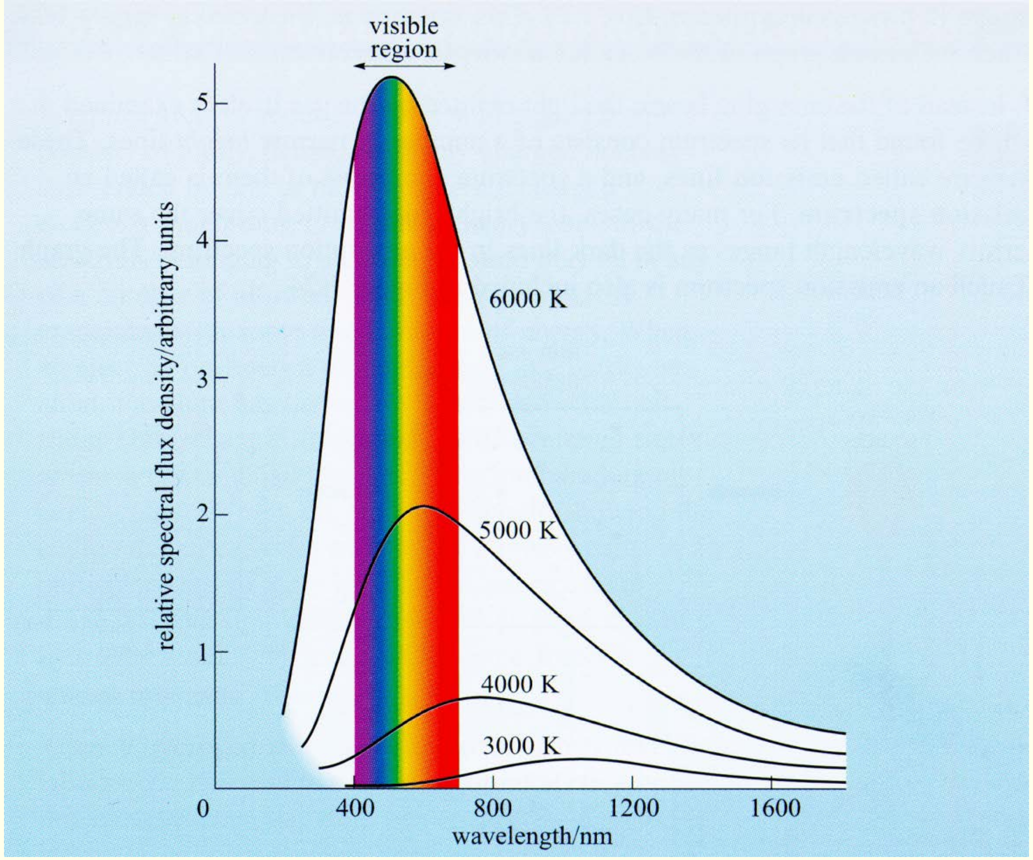
The temperature of the solar surface is approximately 6000 K.

How can we know that ???

Hot steel emitting red light.

Chart to estimate steel temperature in steelworks.

Black-body radiation



Black-body good approximation for opaque bodies where emitted light is much more likely to interact with the material of the source than to escape.

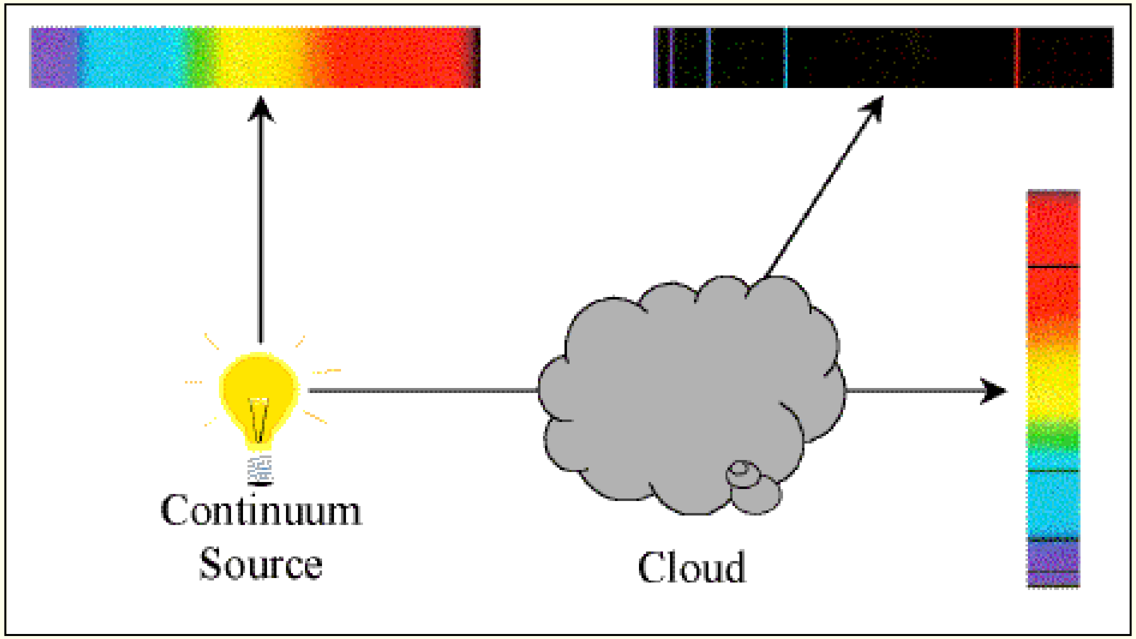
Wien's displacement law

$$\lambda_{peak} = \frac{2.90\times10^{-3}\,\mathrm{m}\cdot\mathrm{K}}{T}$$

Stefan-Bolzmanns law

$$J = \sigma_{SB} T^4$$

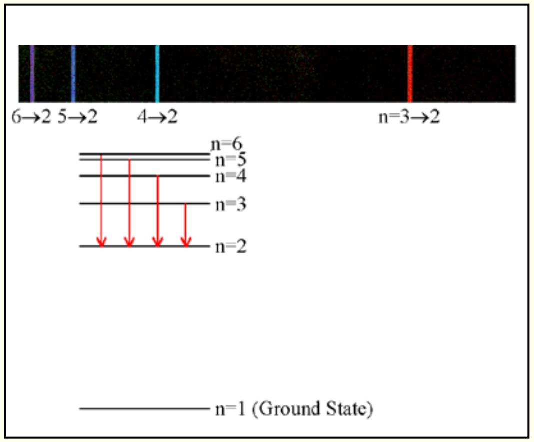
(J = total energy radiated per unit area per unit time)



For non-blackbody thermal light emitter (for example a thin gas) it is more complicated. Spectrum depends on e.g. chemical composition, and how many atoms/molecules happen to be in state with high probability to decay and cause emission.

Black-body radiation

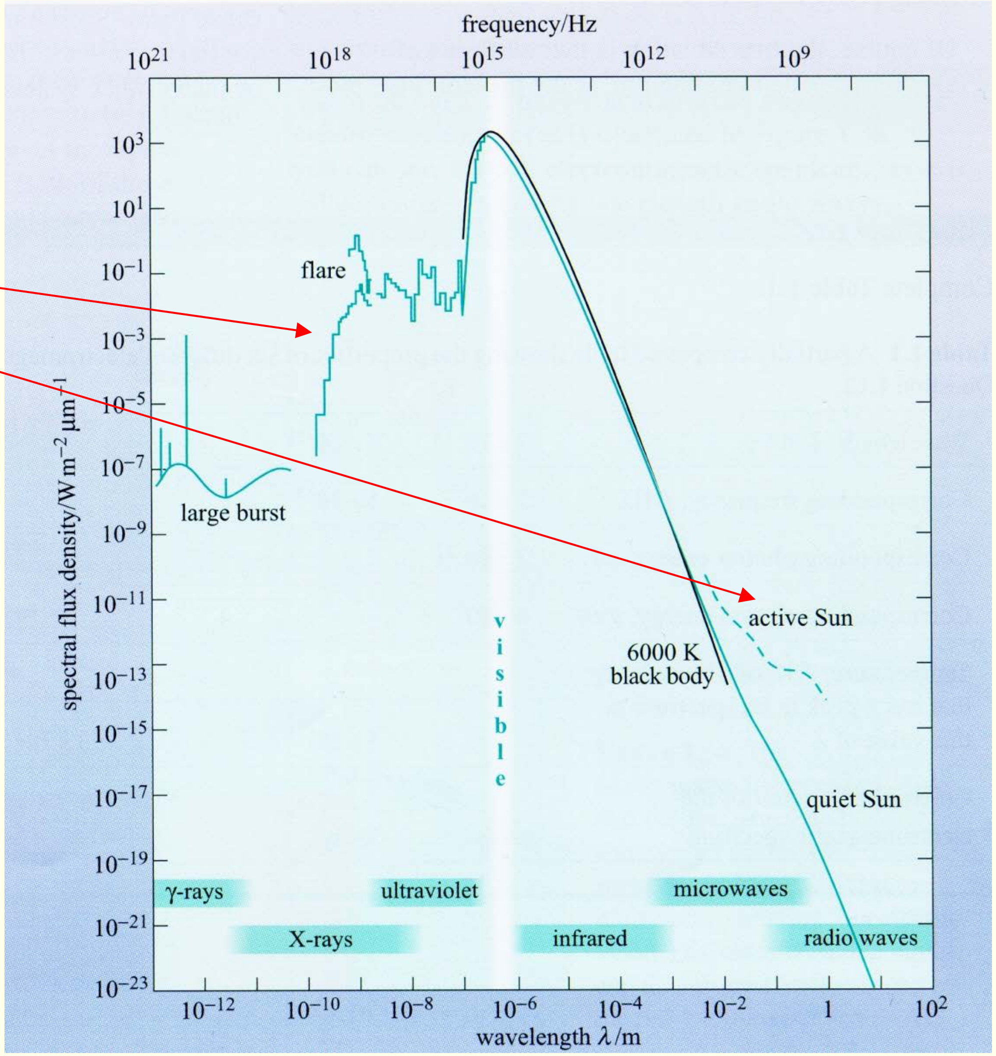
$$\lambda_{peak} = \frac{2.90\times10^{-3}}{T}$$



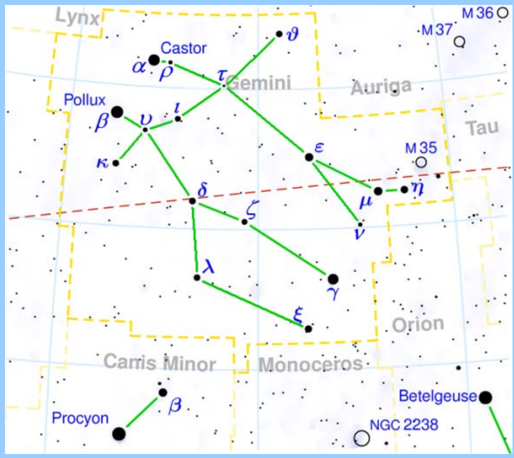
Atomic energy levels

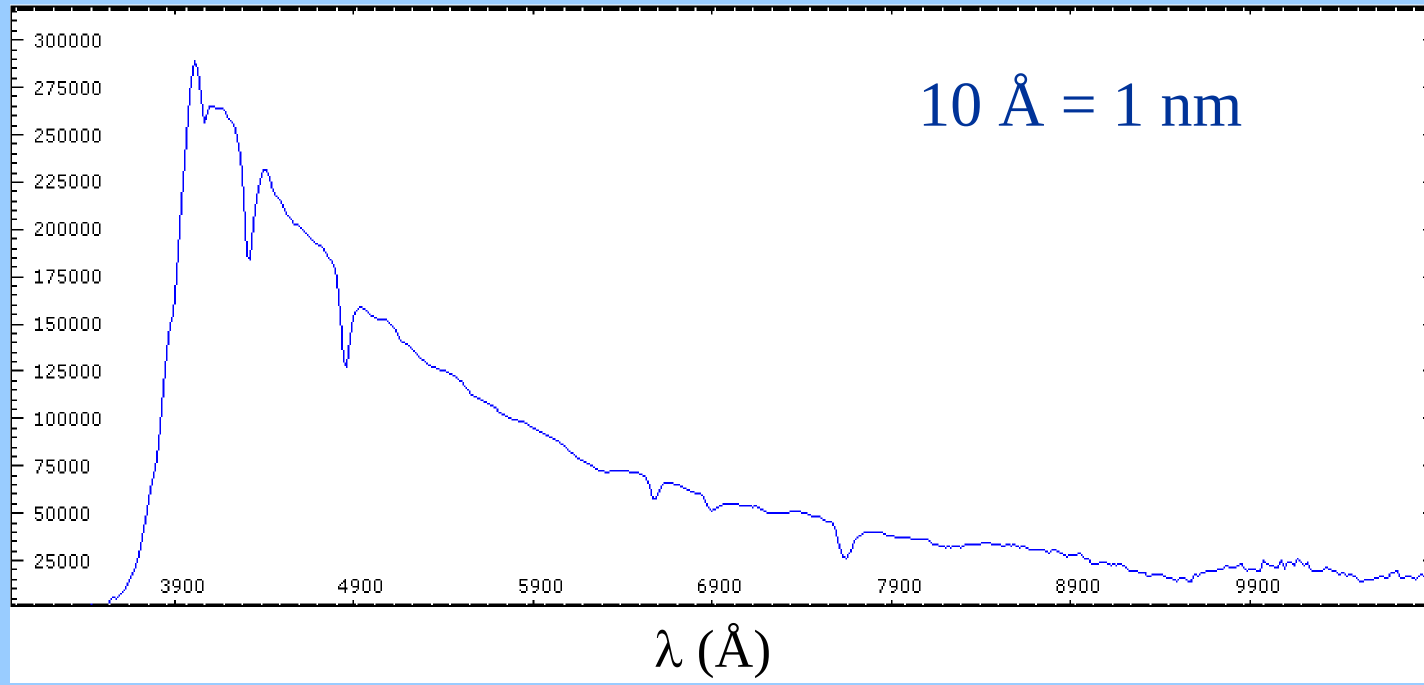
The solar spectrum

Non-blackbody contributions



Estimate the temperature of the star *Gamma Geminorum A0iv* !





Red	T = 2100 K
Green	T = 7200 K
Blue	T = 4700 K
Yellow	T = 9300 K

Estimate the temperature of the star *Gamma Geminorum A0iv* !

Wien's displacement law

$$\lambda_{peak} = \frac{2.90\times10^{-3}}{T}$$

$$T = \frac{2.90\times10^{-3}}{\lambda_{peak}} \approx \frac{2.90\times10^{-3}}{400\times10^{-9}}\,\mathrm{K} \approx 7250\ \mathrm{K}$$

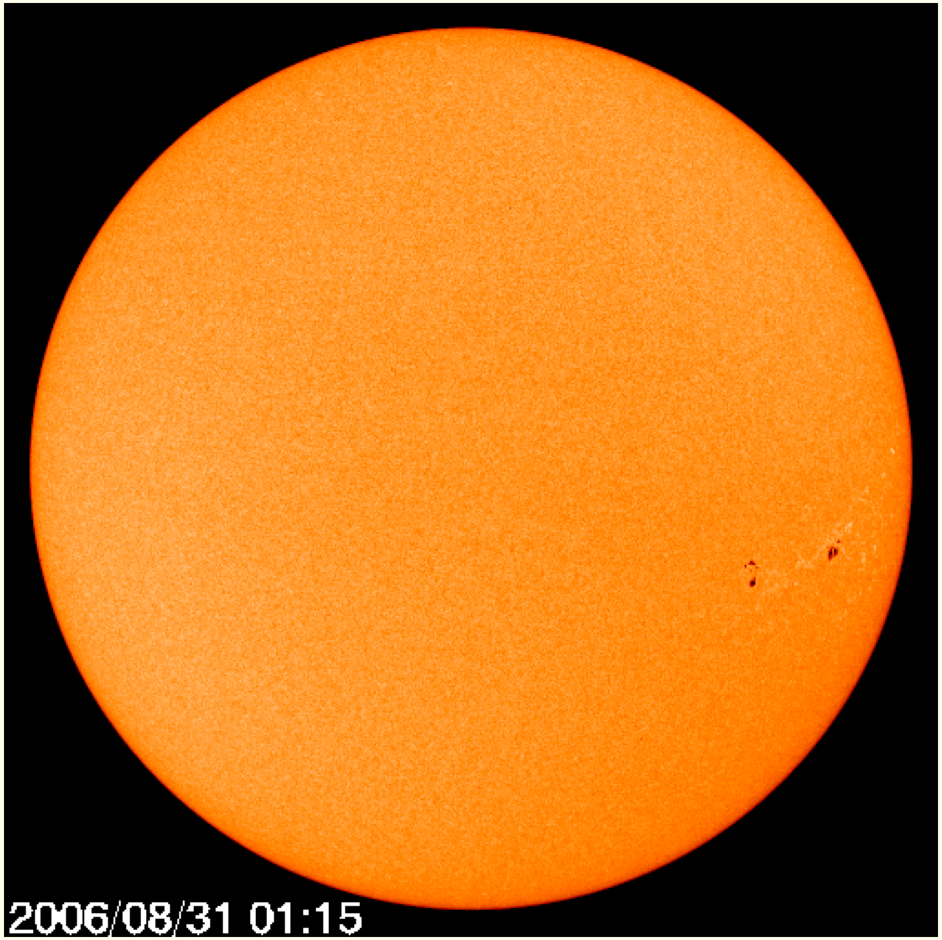
Green T = 7200 K

BREAK!

But think about this:

How can we know anything about the solar interior?

The Sun



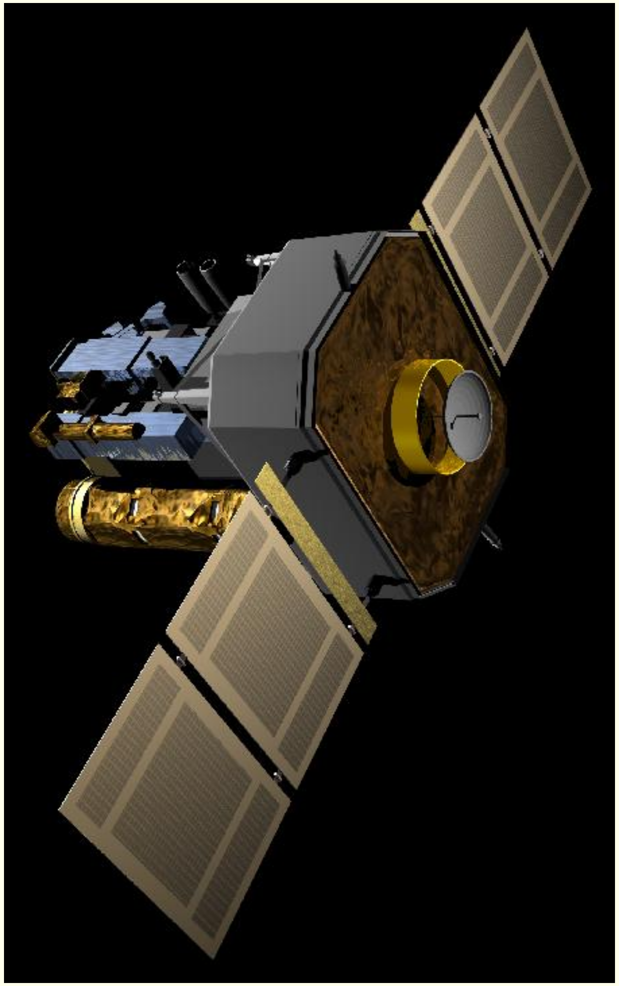
SOHO Michelson Doppler Imager (MDI) 6767 Å continuum images from Stanford University

Basic facts

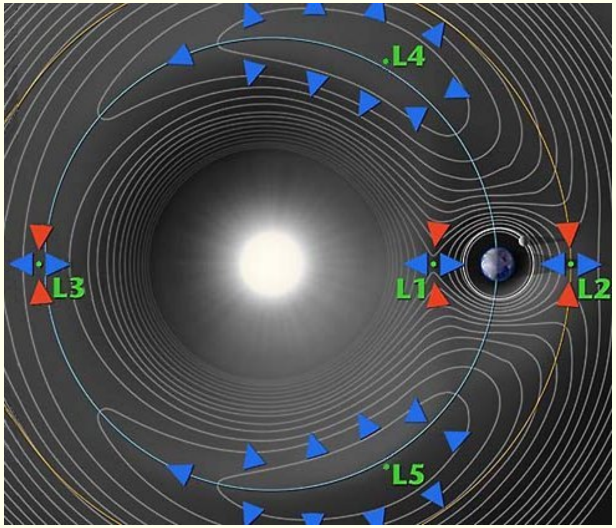
- *diameter:* $1.39 \cdot 10^9\ m \approx 109\ d_E$
- *mass:* $2 \cdot 10^{30}\ kg \approx 333\,000\ m_E$
- *density:* $1.4\ kg/dm^3$
- *radiated effect:* $4 \cdot 10^{26}\ W$
- *age:* $4.5 \cdot 10^9$ *years*

SOHO spacecraft

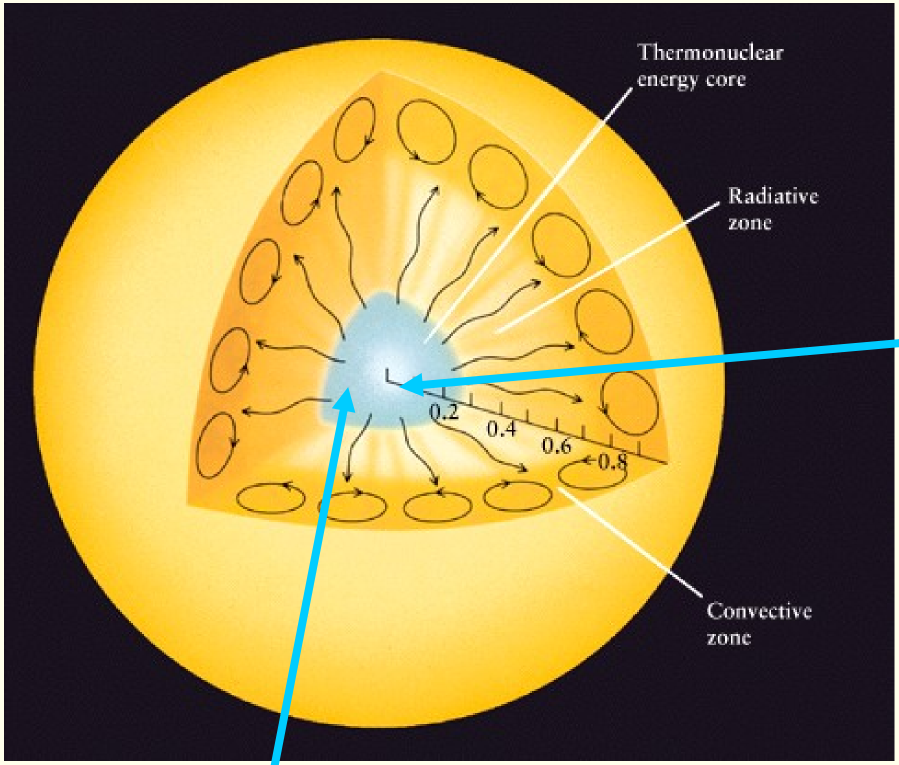
SOlar and Heliospheric Observatory



- Launched 1995
- Orbiting L1
- Collaboration between ESA and NASA
- 12 instruments, including imagers and particle detectors



Sun's interior

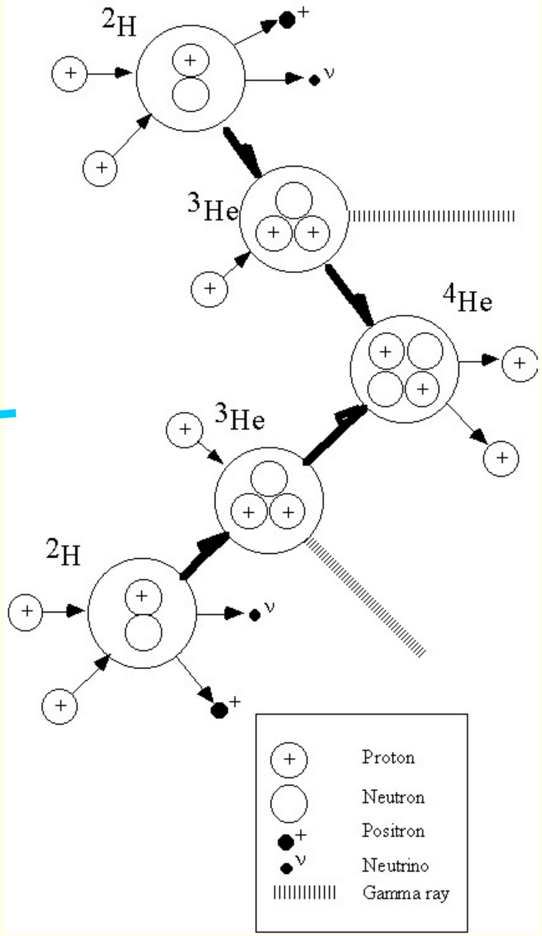


$T = 15 \cdot 10^6$ K

$P = 4 \cdot 10^{26}$ W

(P/m ~ 1mW/kg)

The proton cycle



$$4\,{}^1_1\mathrm{H} \rightarrow {}^4_2\mathrm{He} + 2e^+ + 2\nu_e + 2\gamma$$

Proton cycle

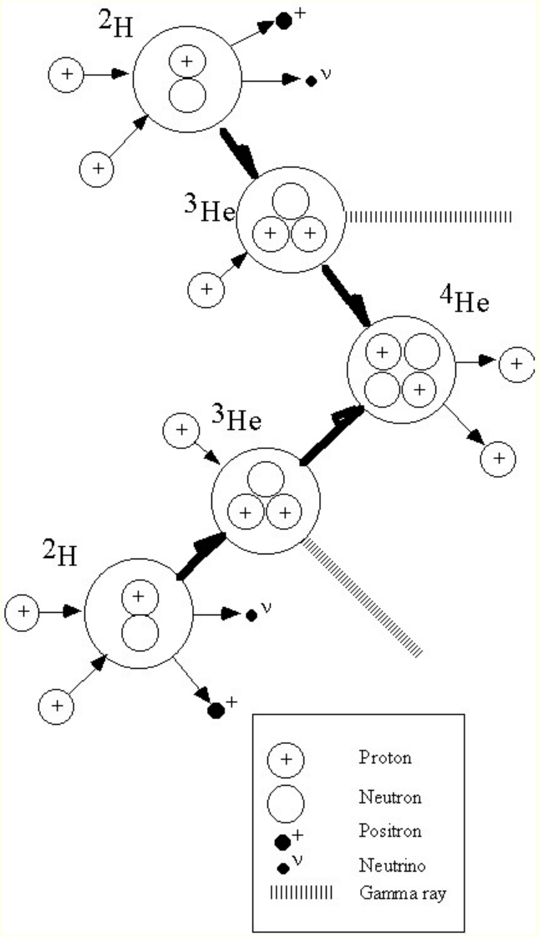
$$2\,{}^1_1\mathrm{H} \rightarrow {}^2_1\mathrm{D} + \mathrm{e}^+ + \nu_e$$

$$2\,{}^1_1\mathrm{H} \rightarrow {}^2_1\mathrm{D} + \mathrm{e}^+ + \nu_e$$

$${}^2_1\mathrm{D} + {}^1_1\mathrm{H} \rightarrow {}^3_2\mathrm{He} + \gamma$$

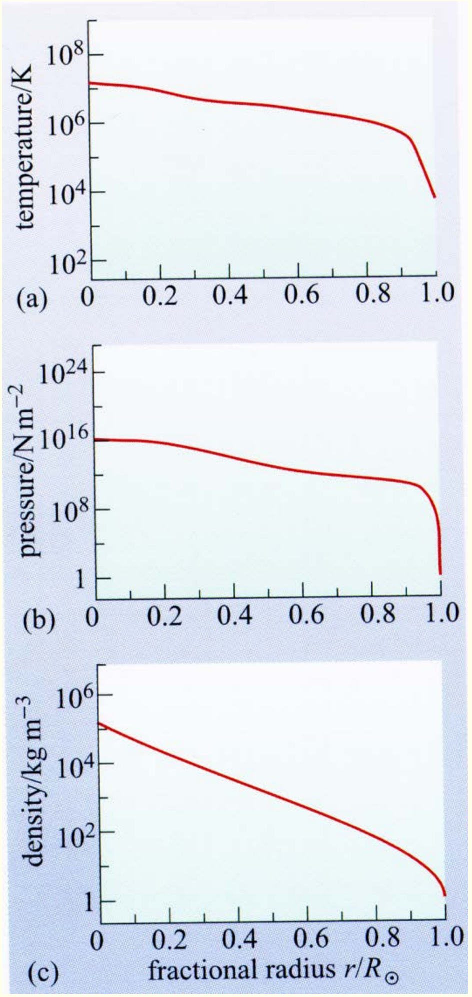
$${}^2_1\mathrm{D} + {}^1_1\mathrm{H} \rightarrow {}^3_2\mathrm{He} + \gamma$$

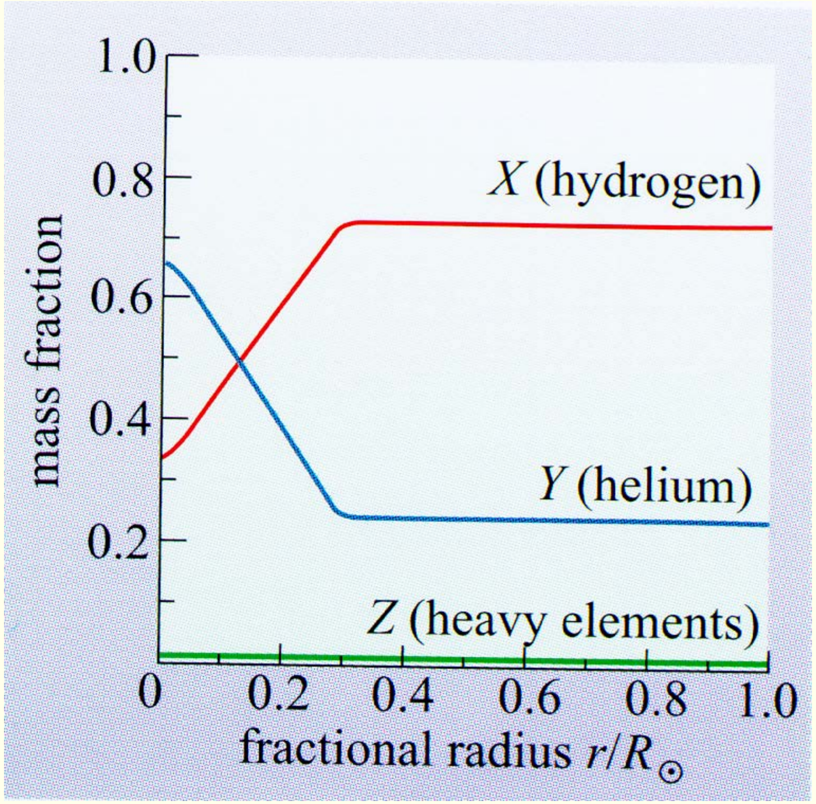
$$2\,{}^3_2\mathrm{He} \rightarrow {}^4_2\mathrm{He} + 2\,{}^1_1\mathrm{H}$$



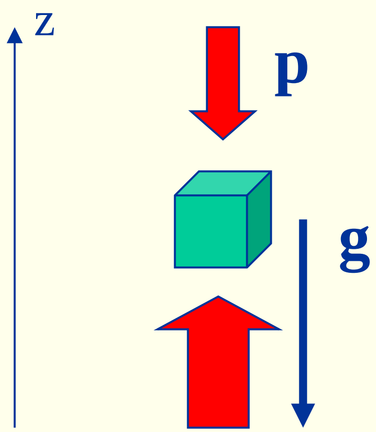
$$4\,{}^1_1\mathrm{H} \rightarrow {}^4_2\mathrm{He} + 2\mathrm{e}^+ + 2\nu_e + 2\gamma$$

Sun's interior





Atmospheric scale height



$$-\frac{dp}{dz} = g\rho$$

hydrostatic equilibrium for a volume element

$$p = nk_B T = \frac{\rho k_B T}{m}$$

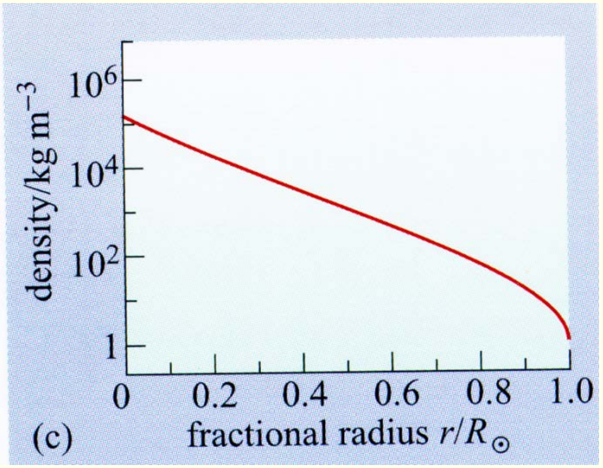
ideal gas law

$$-\frac{k_B T}{m}\frac{d\rho}{dz} = g\rho$$

if T is constant

$$\rho = const \cdot e^{-z/(k_B T/gm)} = const \cdot e^{-z/H}$$

$$\log\rho = const - \frac{z}{H}$$

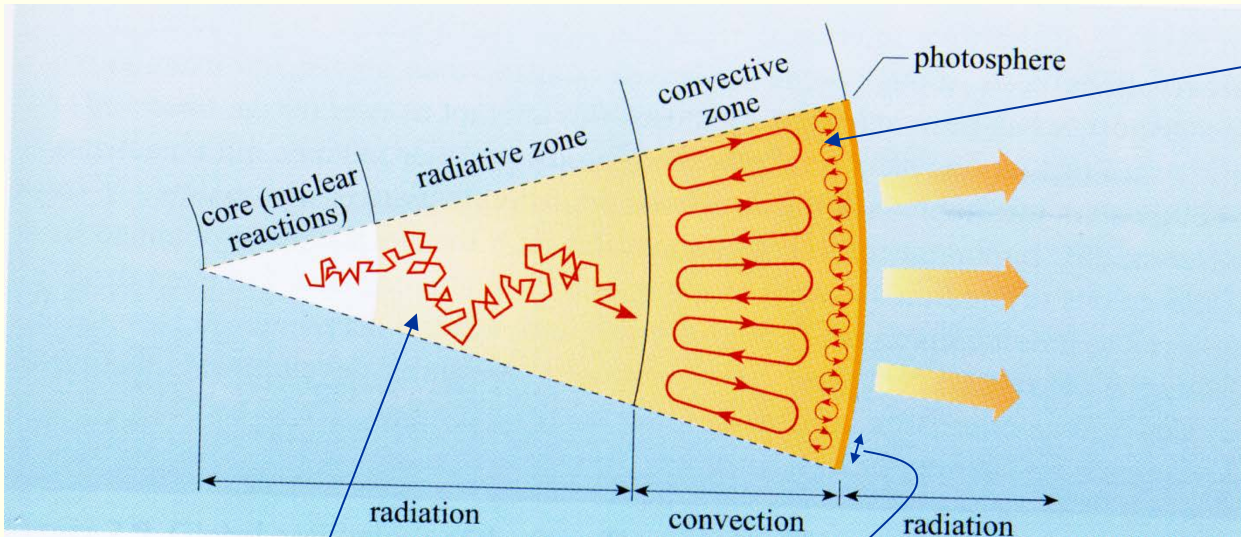


Scale height

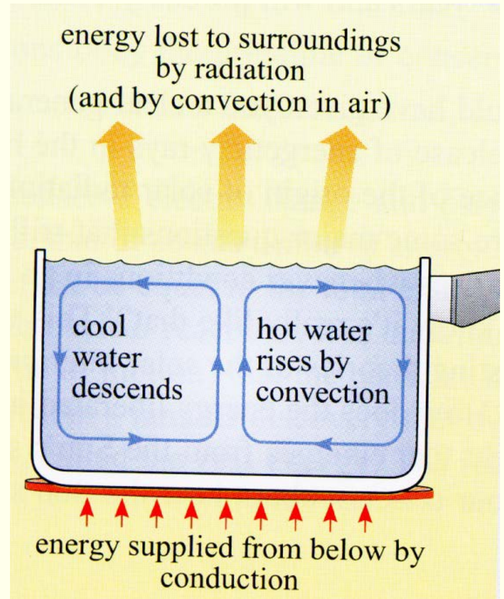
$H = k_B T/gm$

Energy transport in the sun

Transport by convection



~1000 km



Transport by radiation, which interacts with the dense solar matter (scattering and absorption/re-emission).

I takes on average 200 000 years for a photon to reach the photosphere!

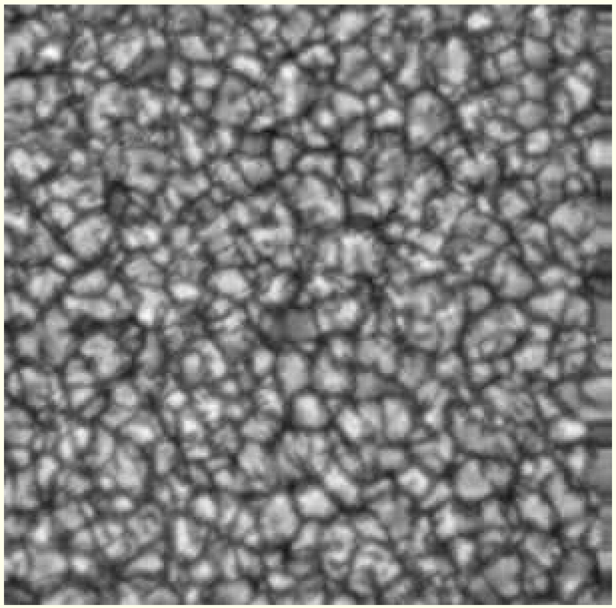
These convection cells are called *granulation*.

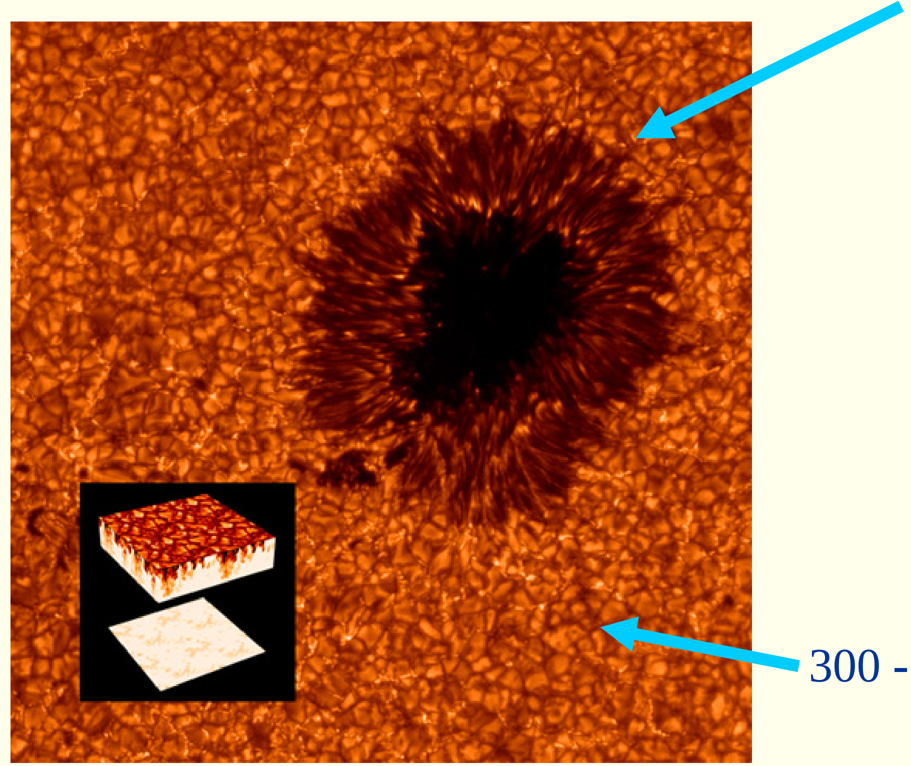
At the photosphere the mean free path of the photons becomes so large that they can reach directly out into space.

Granulation in photosphere

T = 35 min

1000 - 50 000 km

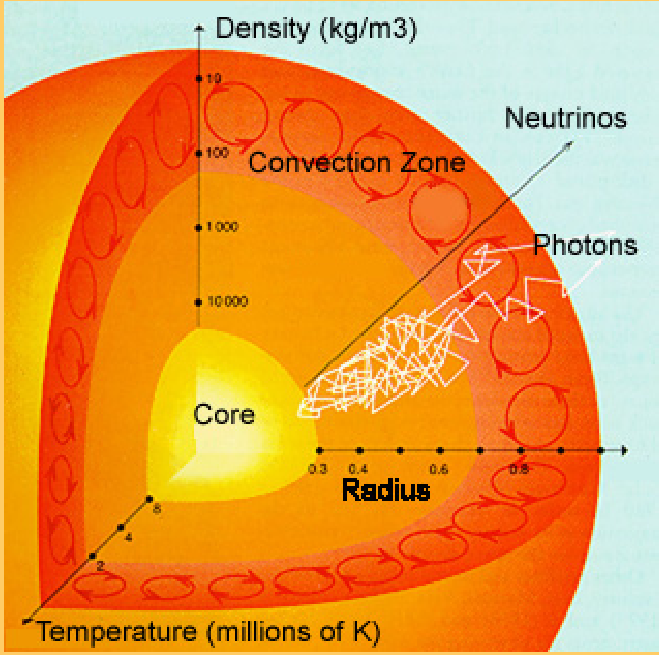




300 - 2000 km

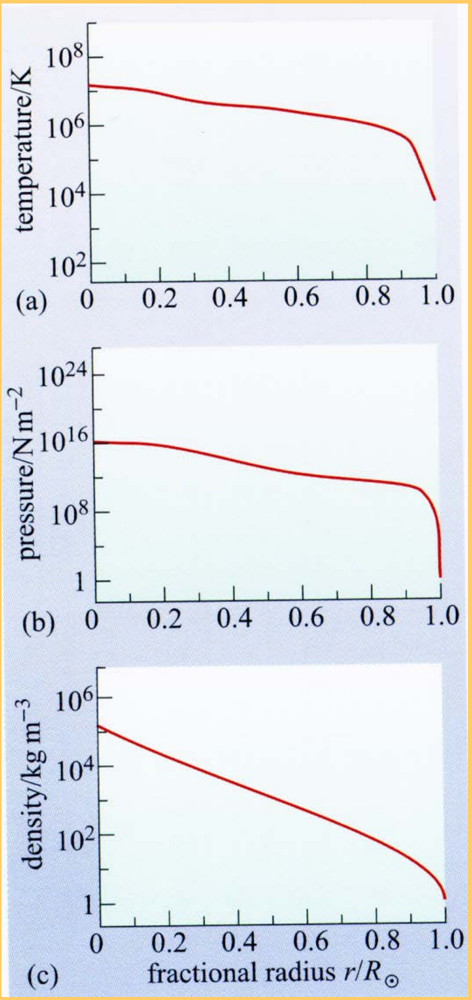
~27 000 km

Life time ~10 min



So how can we know all these details about the solar interior?

1. Solar models



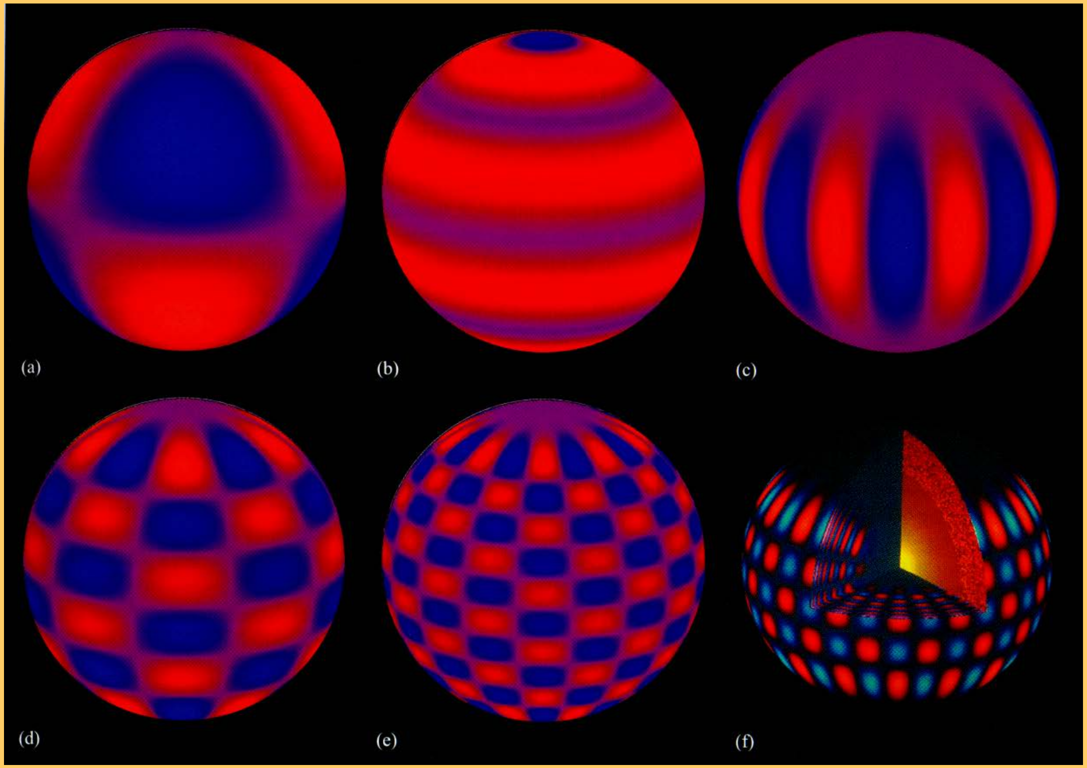
$$4\,{}^1_1\mathrm{H} \rightarrow {}^4_2\mathrm{He} + 2\mathrm{e}^+ + 2\nu_e + 2\gamma$$

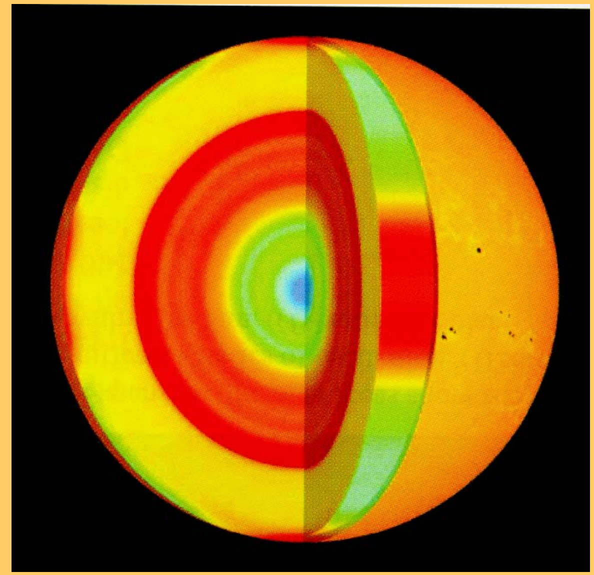
Models of nuclear reactions etc are tuned to *boundary conditions.*

These are e.g.

- *sun's radius R*
- *total mass M*
- *luminosity, L,*
- *surface temperature T,*
- *chemical composition etc.*

2. Helioseismology



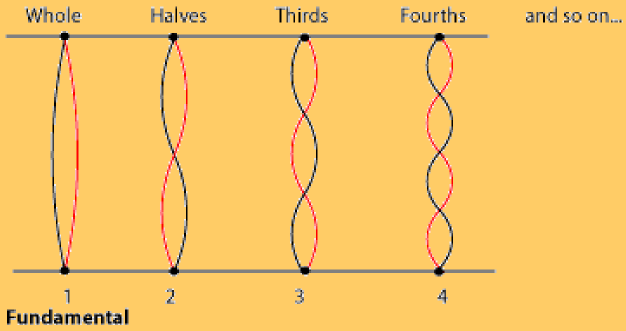


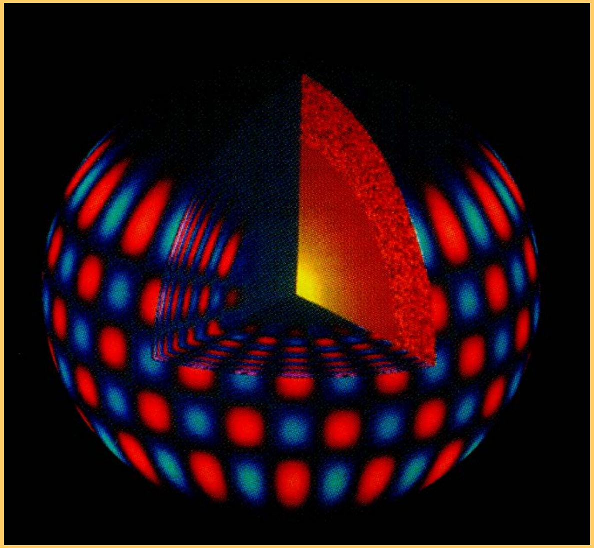
Sound speed.

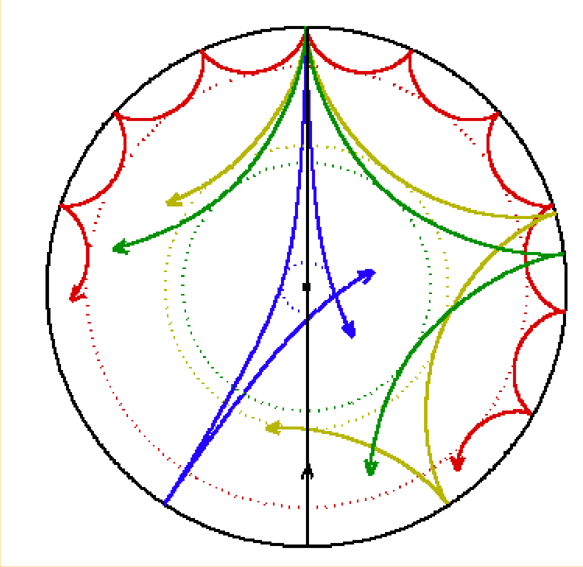
Pressure waves (”sound waves”) on the solar surface can give information of e.g. sound speed in solar interior, which depends on temperature and density.

Helioseismology

String







Fundamental: $\lambda = 2L$

$$\lambda f = c_S$$

$$c_S = 2fL$$

$$c_S = \sqrt{\frac{5p}{3\rho}}$$

Last Minute!

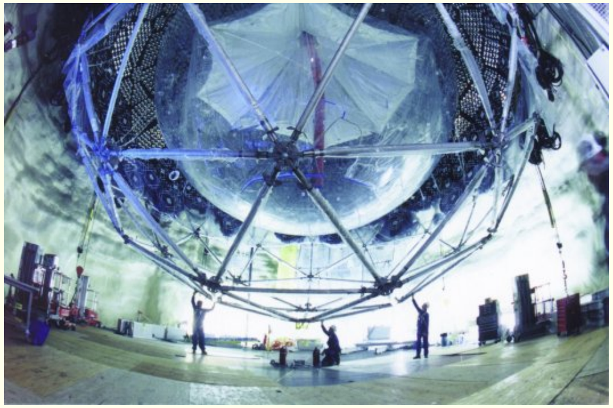
Last Minute!

- What was the most important thing of today’s lecture? Why?
- What was the most unclear or difficult thing of today’s lecture, and why?
- Other comments

KTH
VETENSKAP
OCH KONST

Solar neutrino problem

$$4\,{}^1_1\mathrm{H} \rightarrow {}^4_2\mathrm{He} + 2\mathrm{e}^+ + 2\nu_e + 2\gamma$$



Sudbury Neutrino Detector (Canada)

Neutrino detectors only gave 1/3 of expected neutrino flux.

Explanation:

New theories have shown that different ”flavours” of neutrions can change into each other:

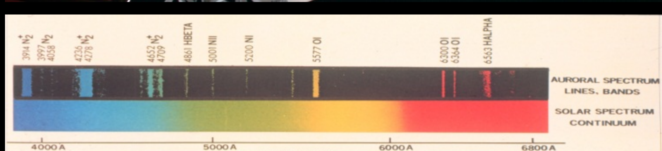
$$\nu_e \leftrightarrow \nu_\mu \leftrightarrow \nu_\tau \leftrightarrow \nu_e$$

During travel from sun to earth, only 1/3 probability that neutrino is electron neutrino, which was the only one that could be detected by early neutrino detectors.

Courses at the Alfvén Laboratory

EF2230 EXPERIMENTAL METHODS IN SPACE PLASMA PHYSICS , 6 ECTS *credits, period 2*

- operation principles of experimental techniques in space plasma physics
- interpretation of measurements
- technical implementations of various measurements techniques
- identify major limitations, and order of magnitude estimate of performance



Hands-on:

- Critical analysis and oral presentation
- Data aquisition or analysis using commercial software